

Local Group Analogs in a cosmological context:

Relating the velocity structure to the cosmic web

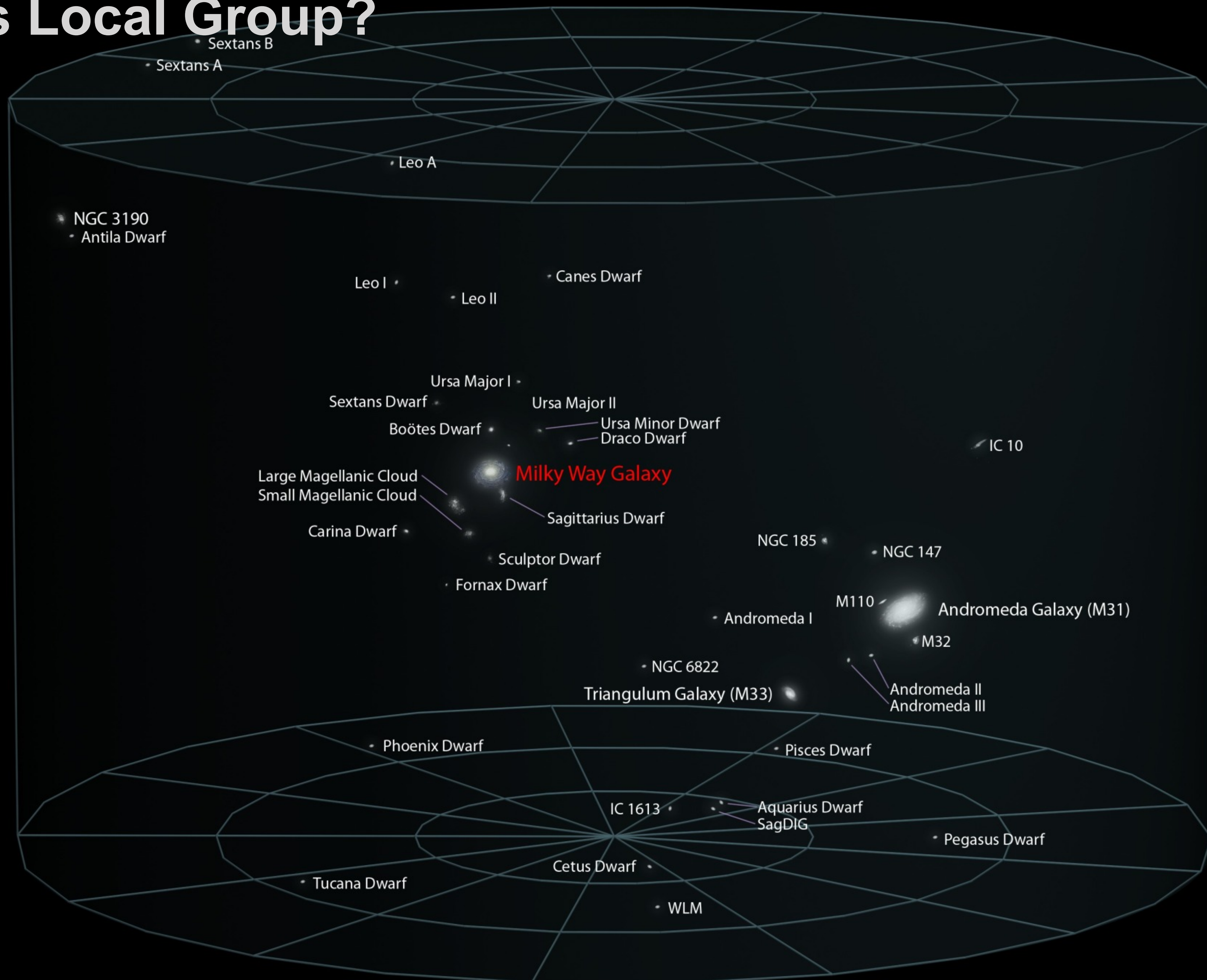
KAI WANG | 王凯

ICC & CEA, Durham University

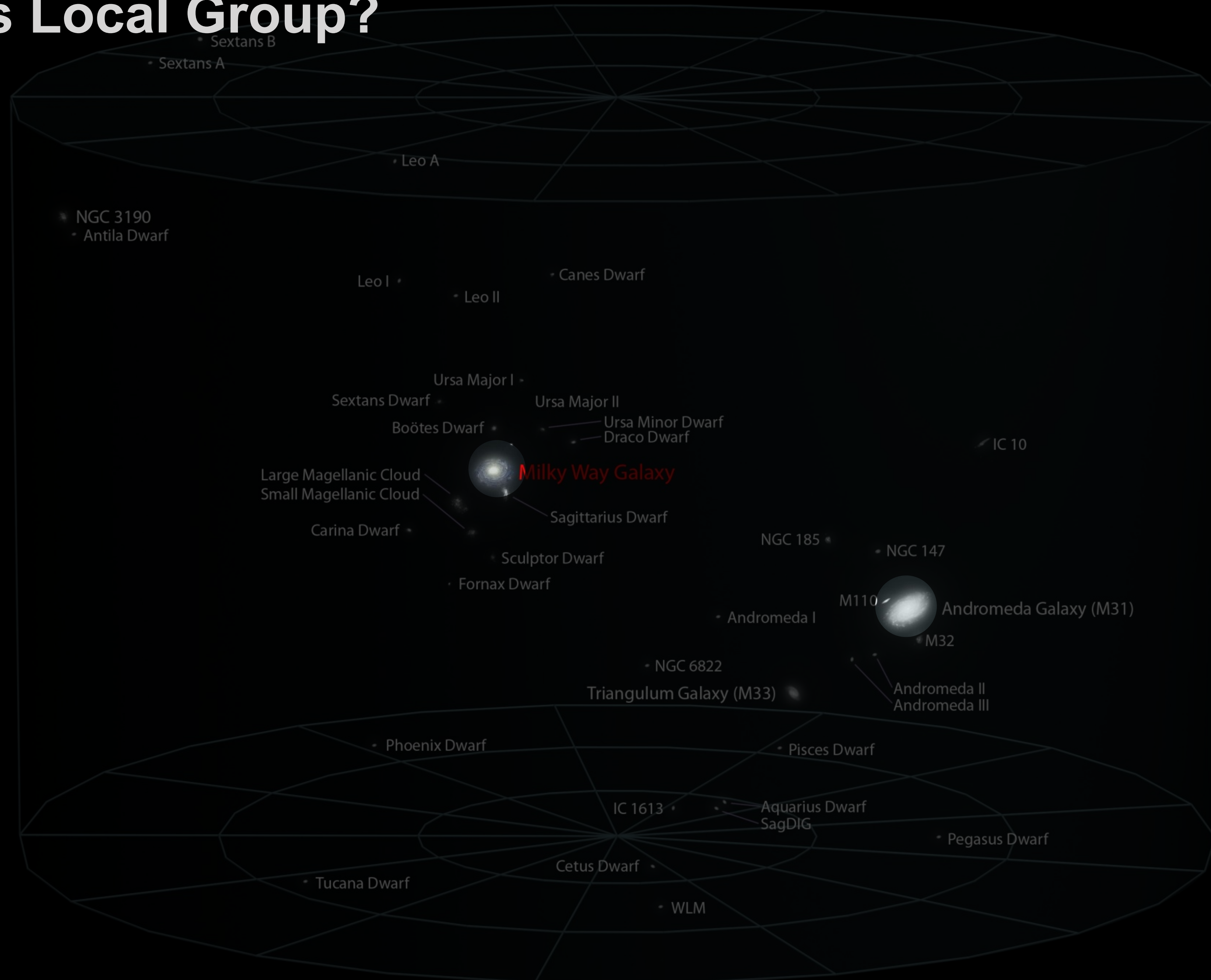
with Peder Norberg, Azadeh Fattahi, Louis E. Strigari

(13mins) **National Astronomical Meeting** Jul. 8, 2025 @ Durham

• What is Local Group?



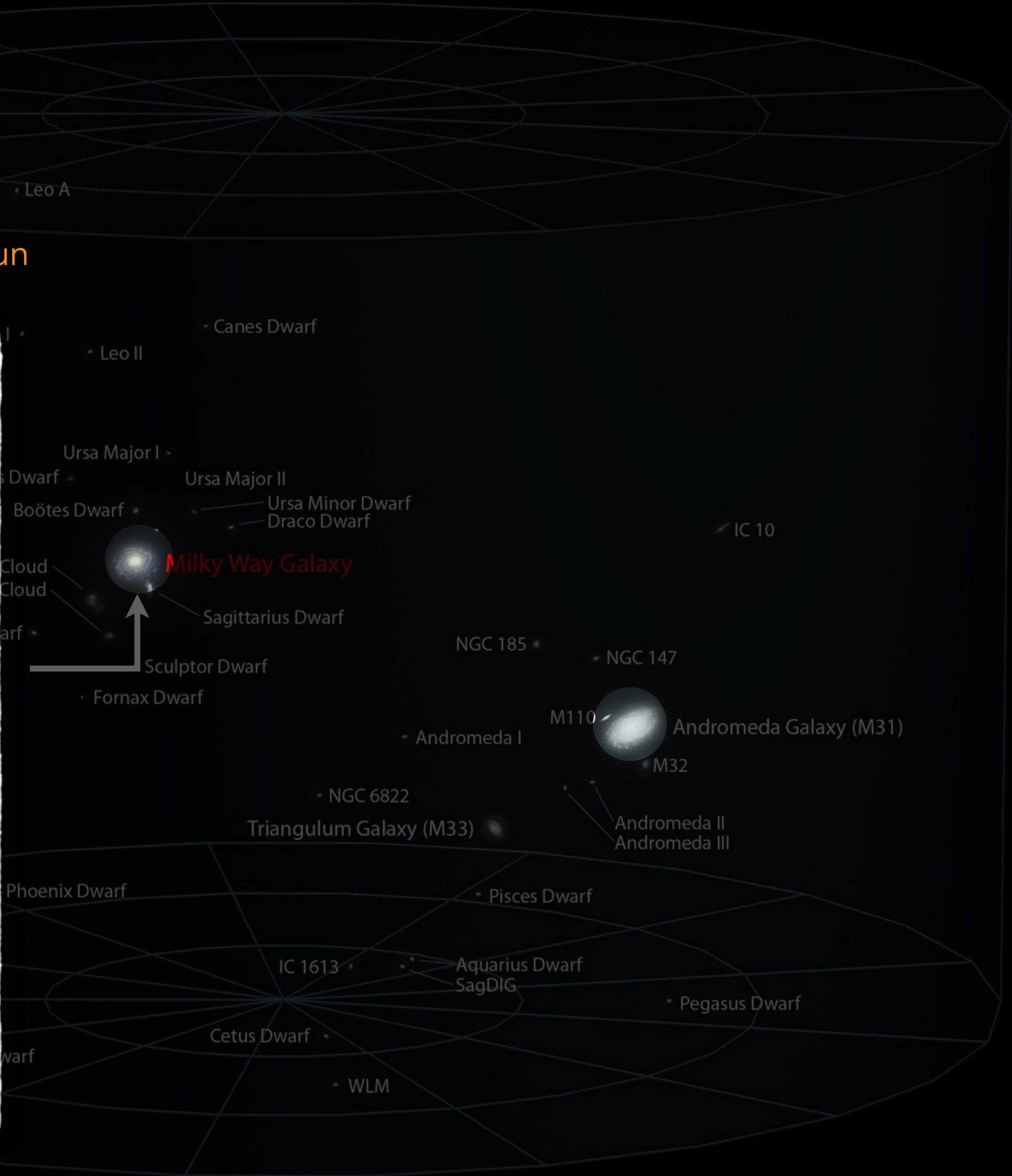
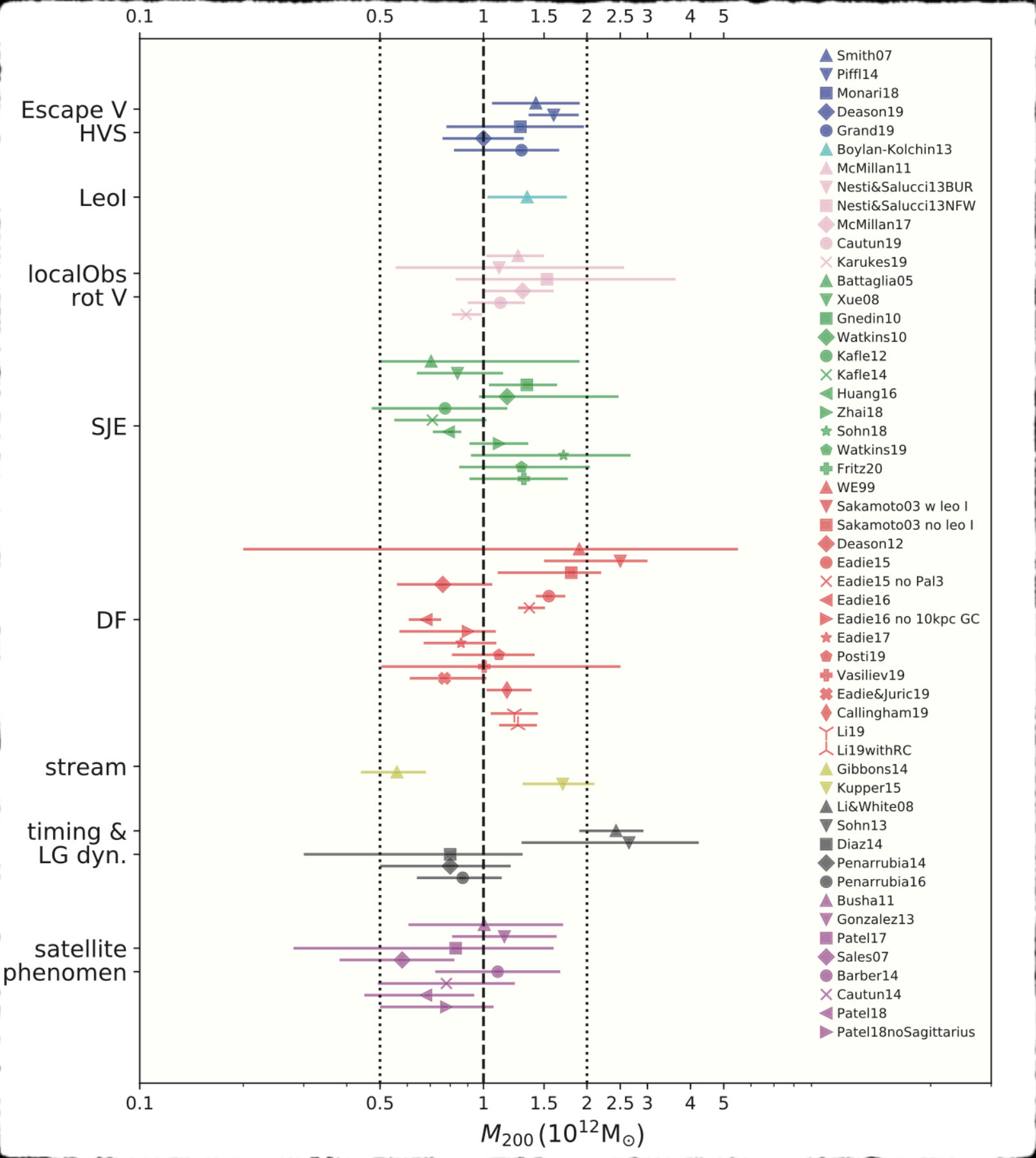
● What is Local Group?



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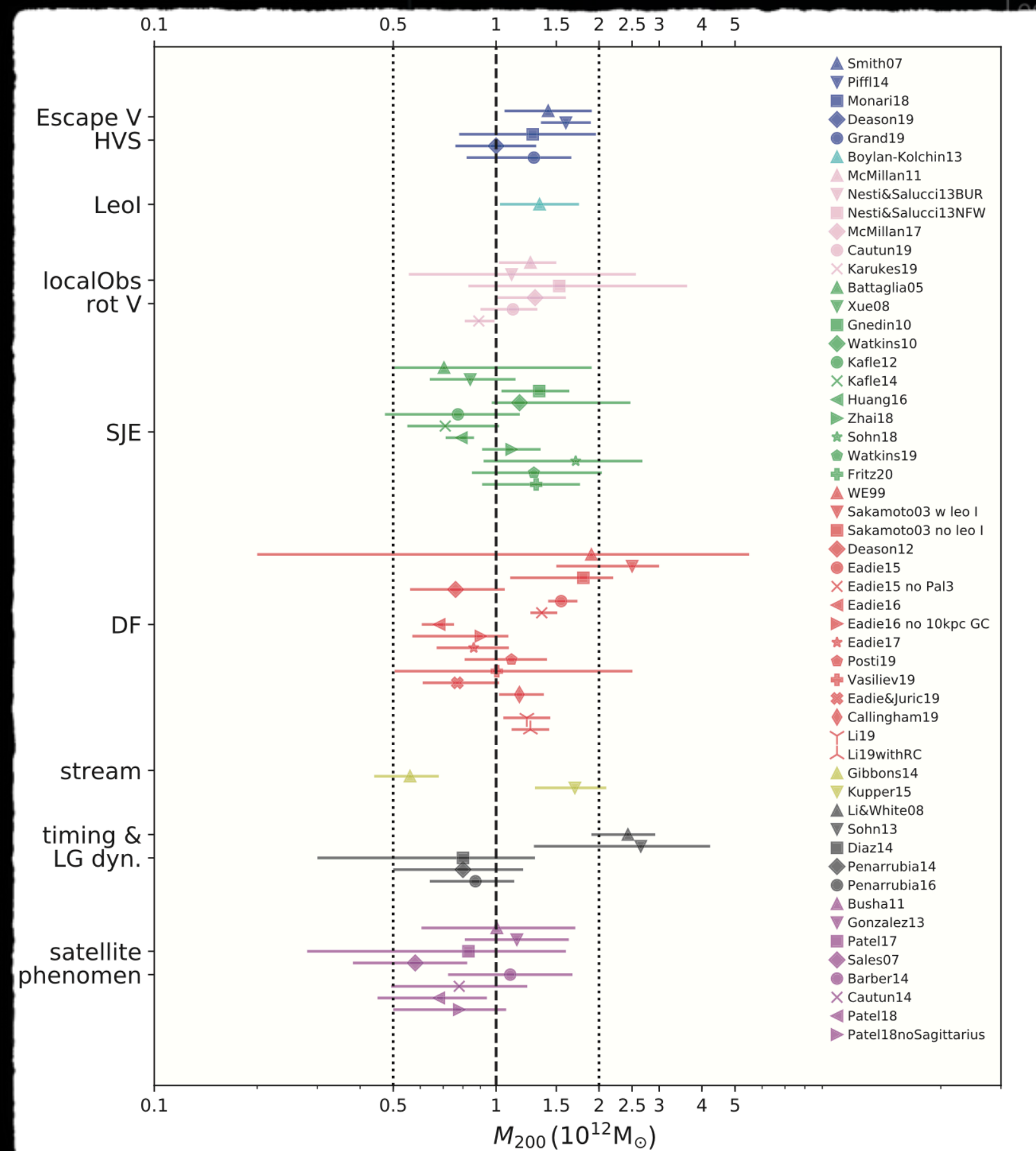
◆ Milky Way mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

e.g. Wang+20



✦ Milky Way mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

e.g. Wang+20



◆ M31 mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

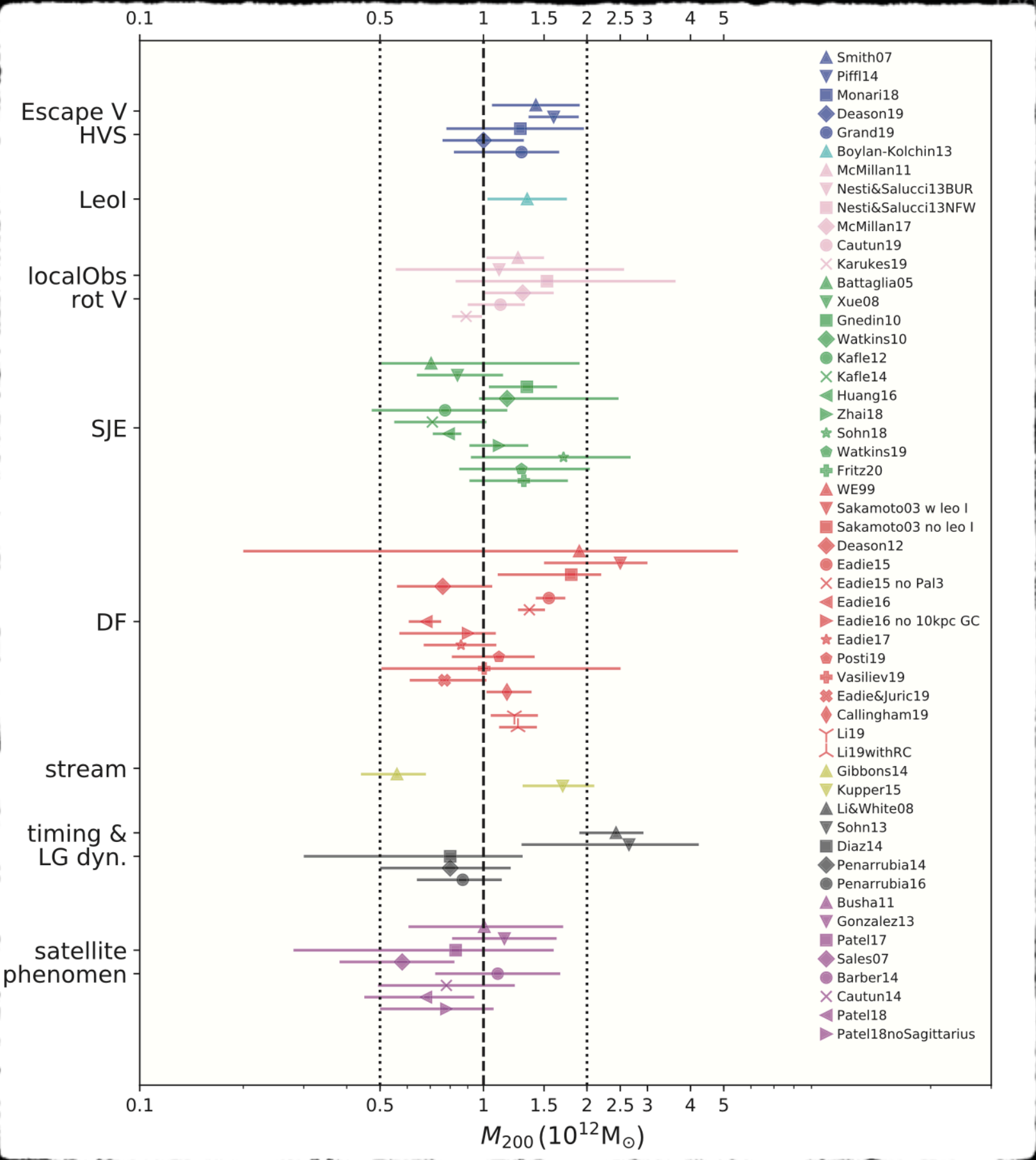
from the timing argument

e.g. Kahn & Woltjer 1959, van der Marel+08, 12

• What is Local Group?

◆ Milky Way mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

e.g. Wang+20



◆ $V_{\text{M31} \rightarrow \text{MW}, r} \sim -110 \text{ km/s}$

◆ $V_{\text{M31} \rightarrow \text{MW}, t} \sim 80 \text{ km/s}$

from redshift and
proper motion measurement
e.g. van der Marel+08, 12, Salomon+21

◆ M31 mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

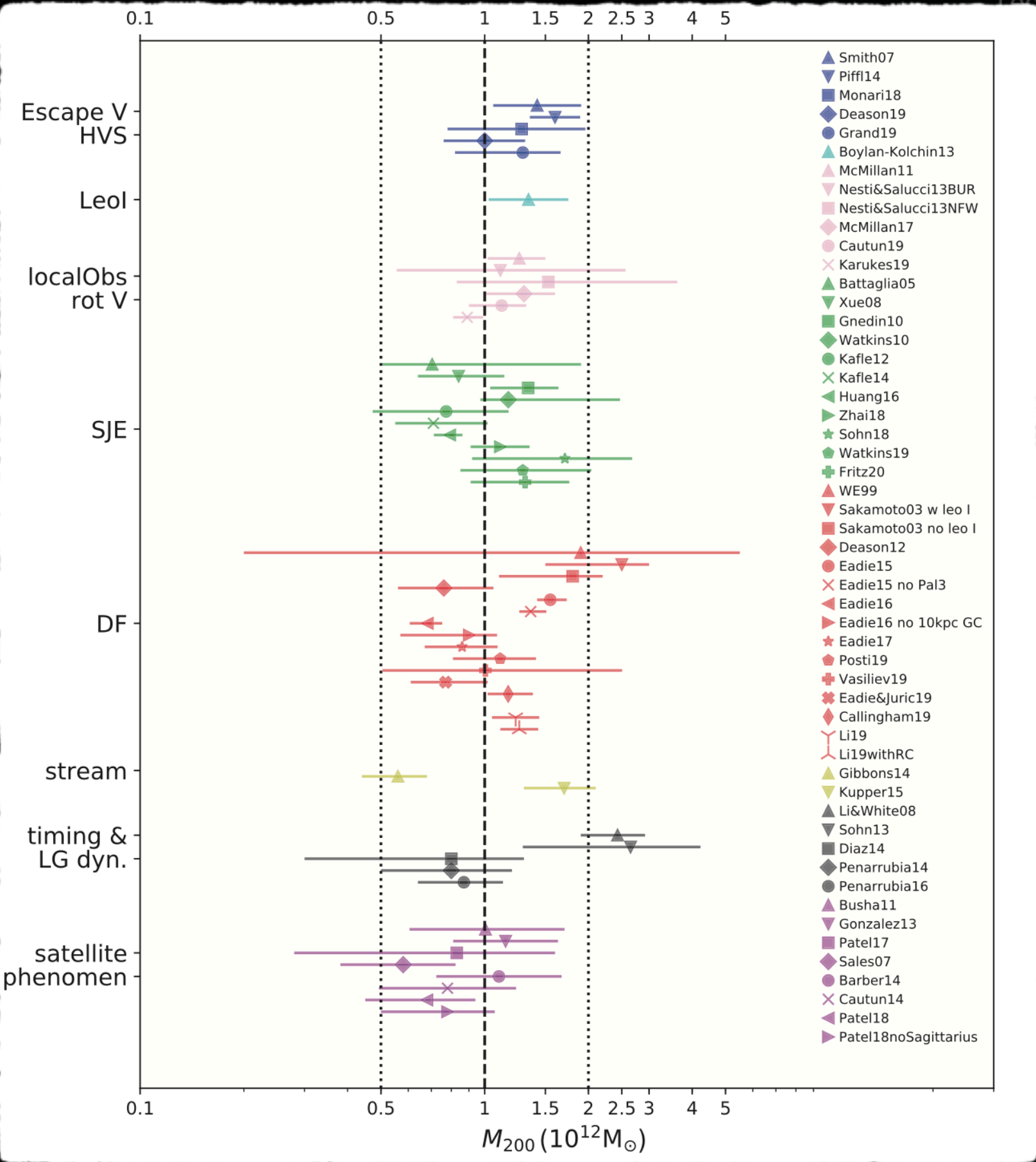
from the timing argument

e.g. Kahn & Woltjer 1959, van der Marel+08, 12

• What is Local Group?

◆ Milky Way mass $\sim 10^{12} h^{-1} M_{\text{sun}}$

e.g. Wang+20



◆ MW-M31 separation: $\sim 785 \text{ kpc}$

from the Tip of the Red Giant Branch (TRGB)

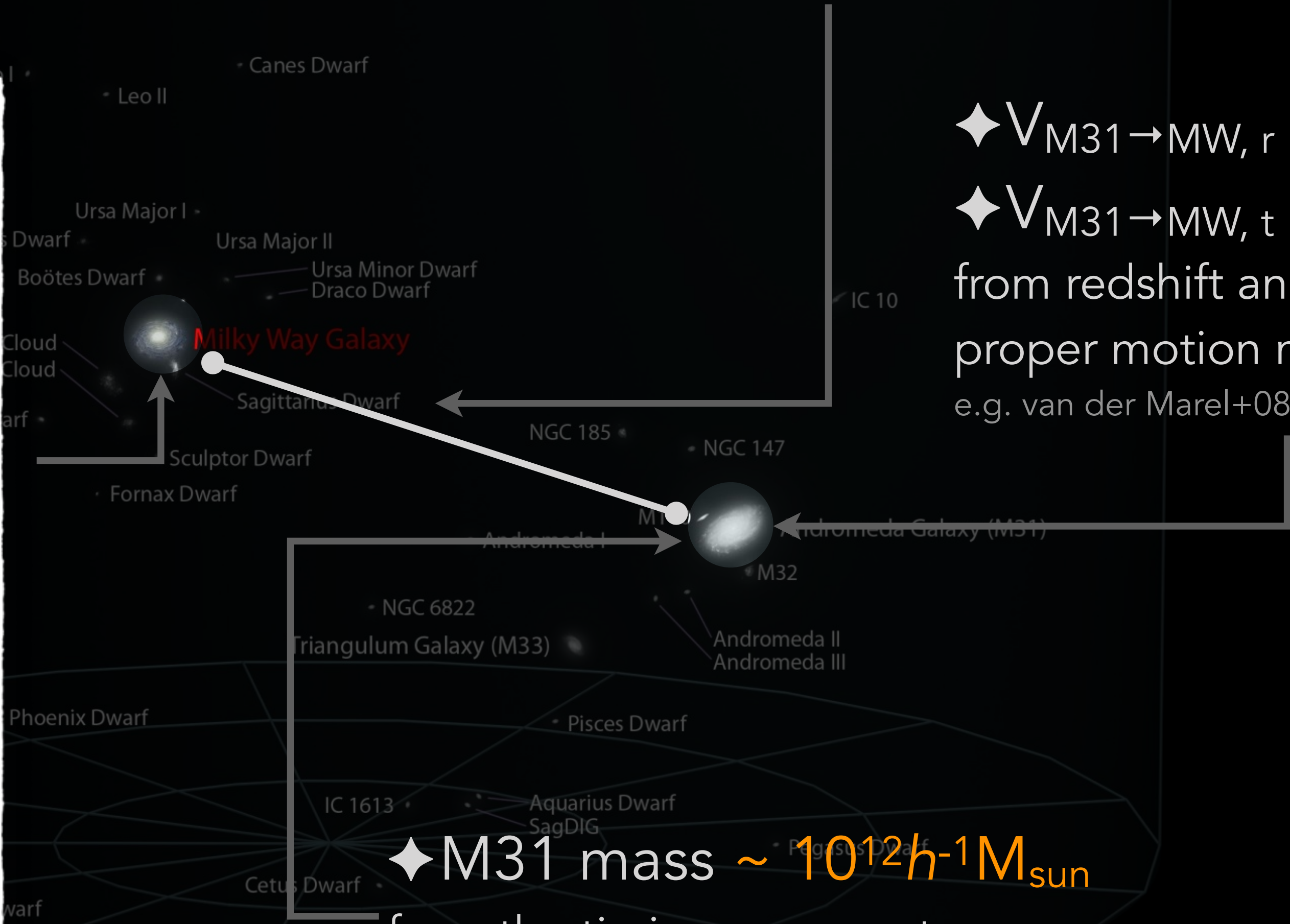
e.g, McConnachie+05

◆ $V_{\text{M31} \rightarrow \text{MW}, r} \sim -110 \text{ km/s}$

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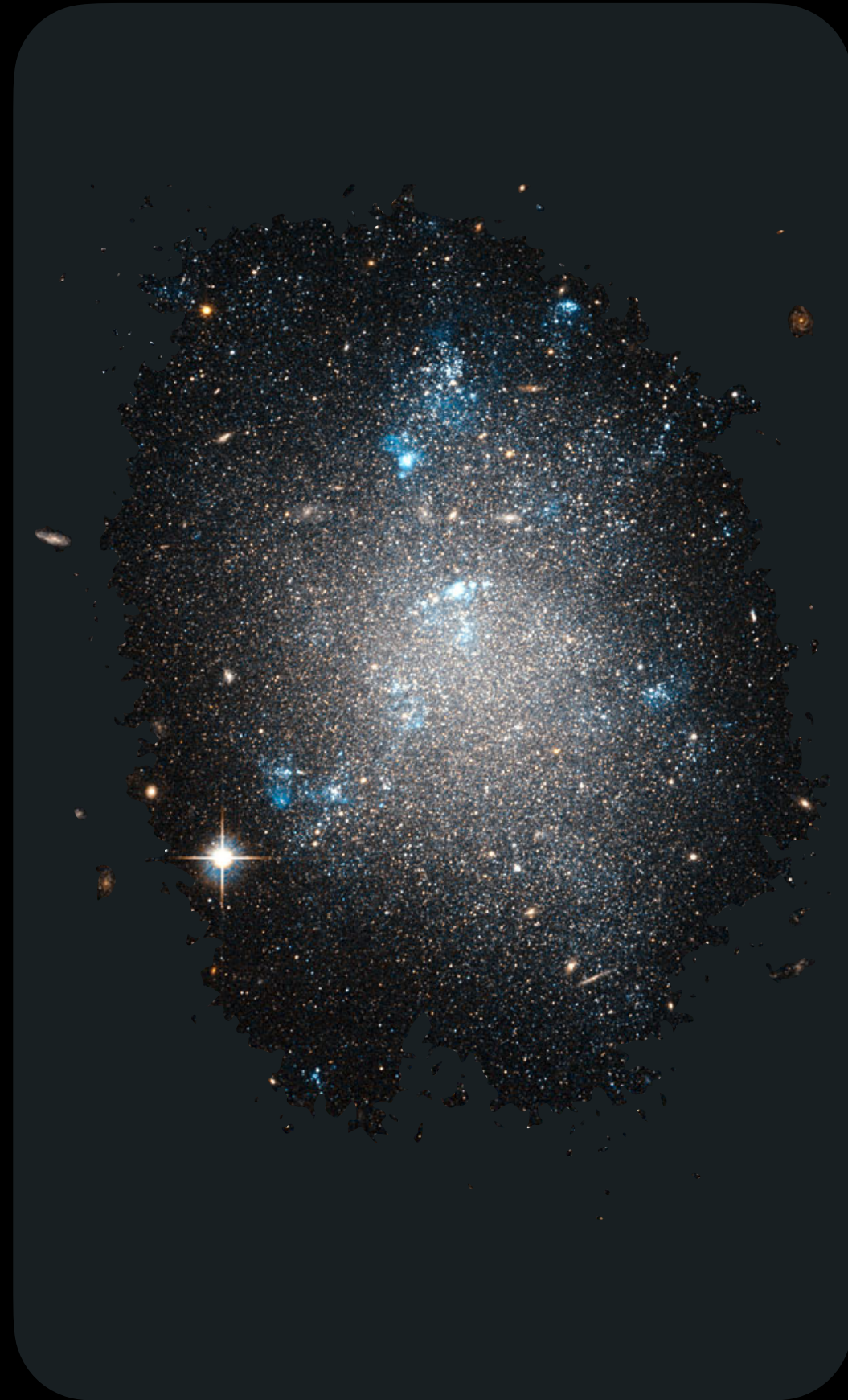
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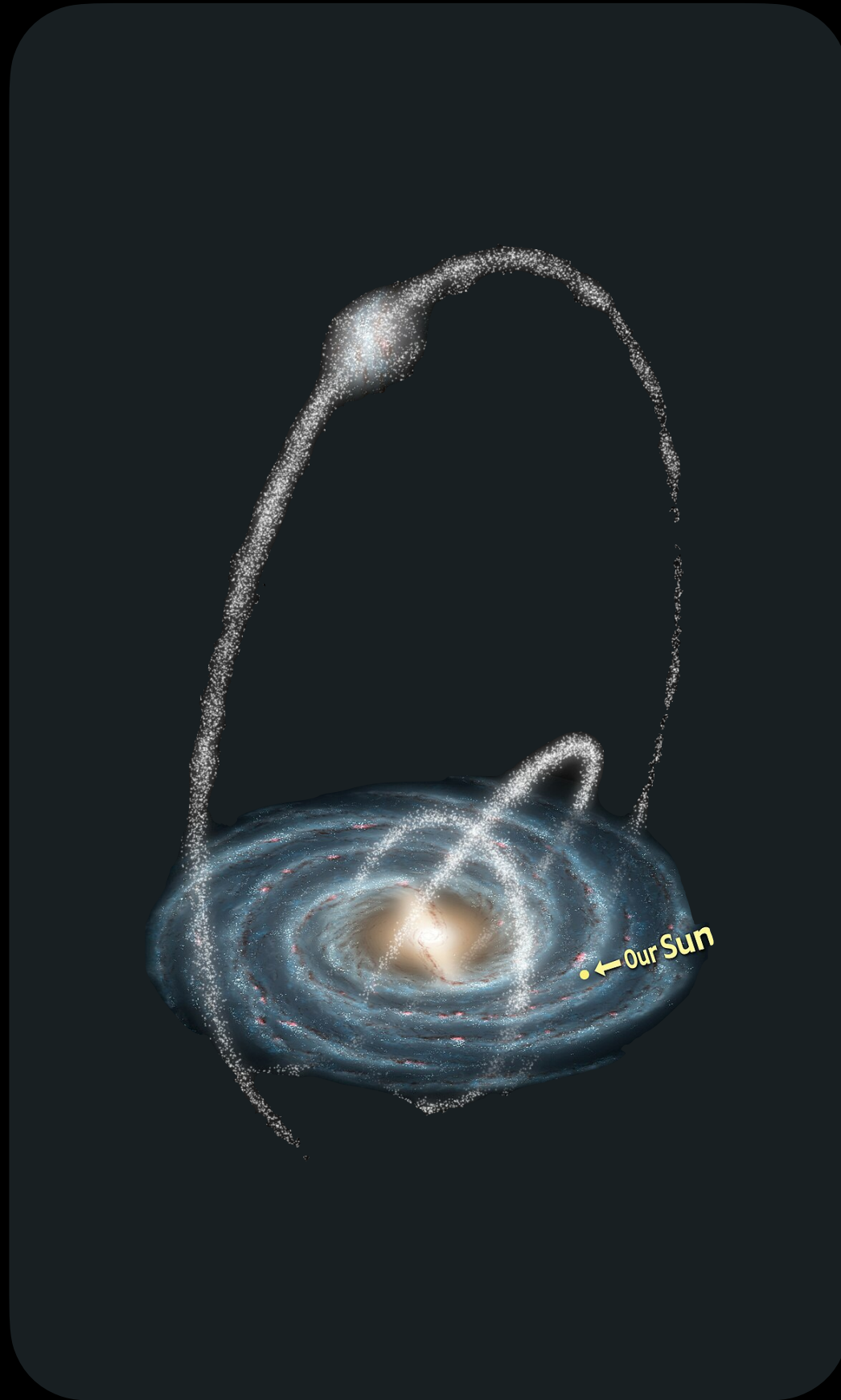
e.g. Kahn & Woltjer 1959, van der Marel+08, 12

- What can we learn from the Local Group

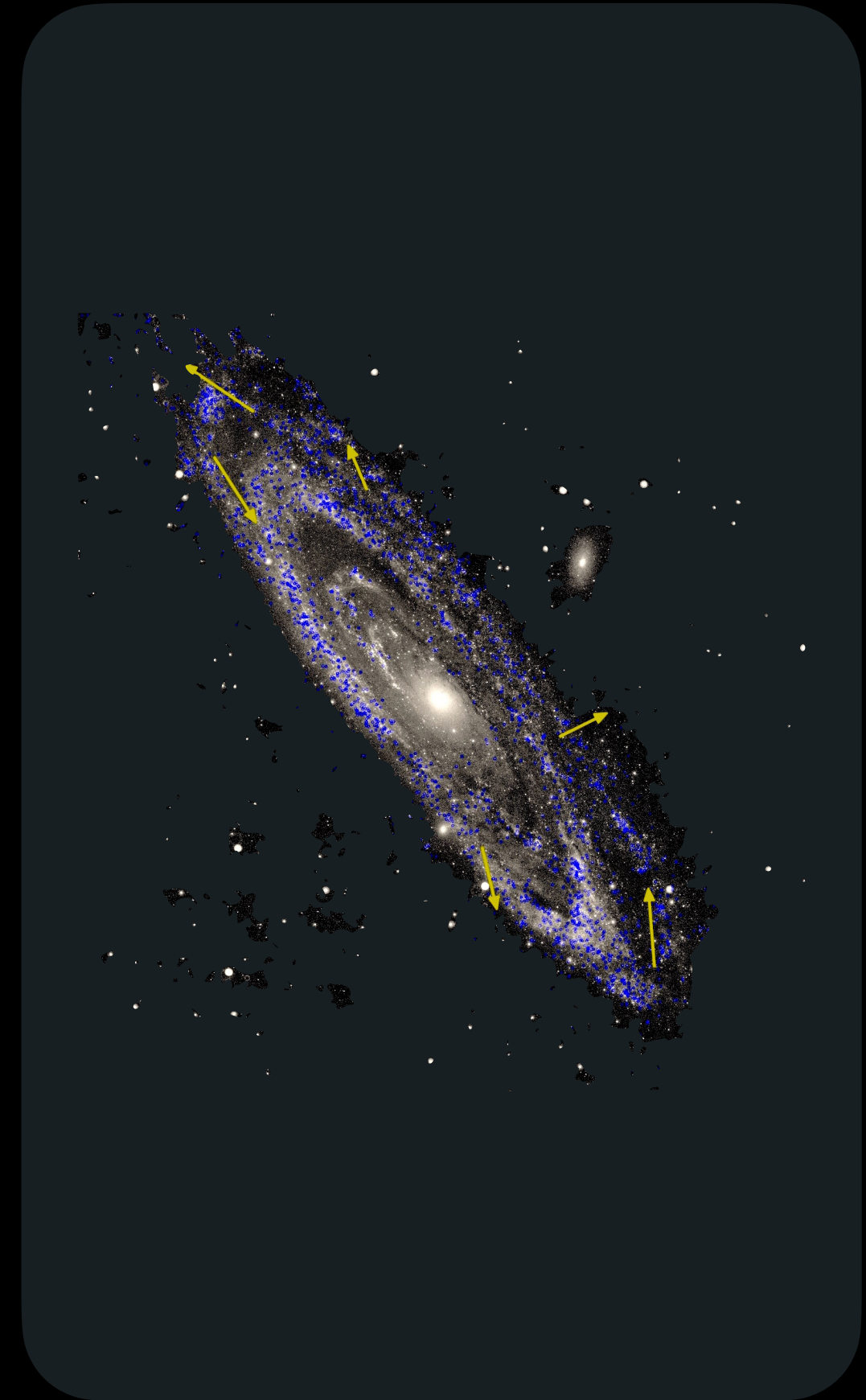
Many things that are **only observable** in our Local Group:



Dwarf galaxies



Stellar streams

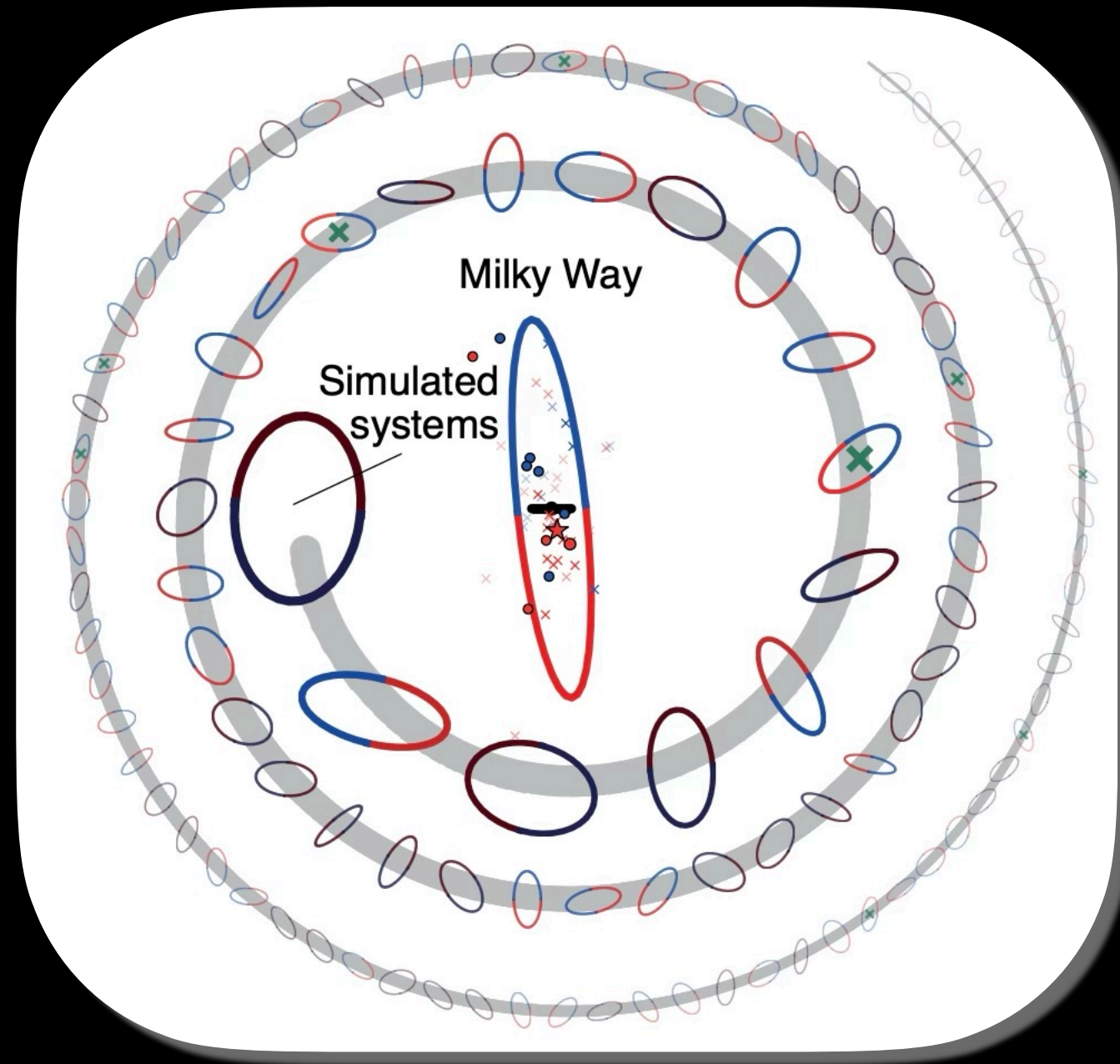


Proper motion

- **New universe crisis from our Local Group**

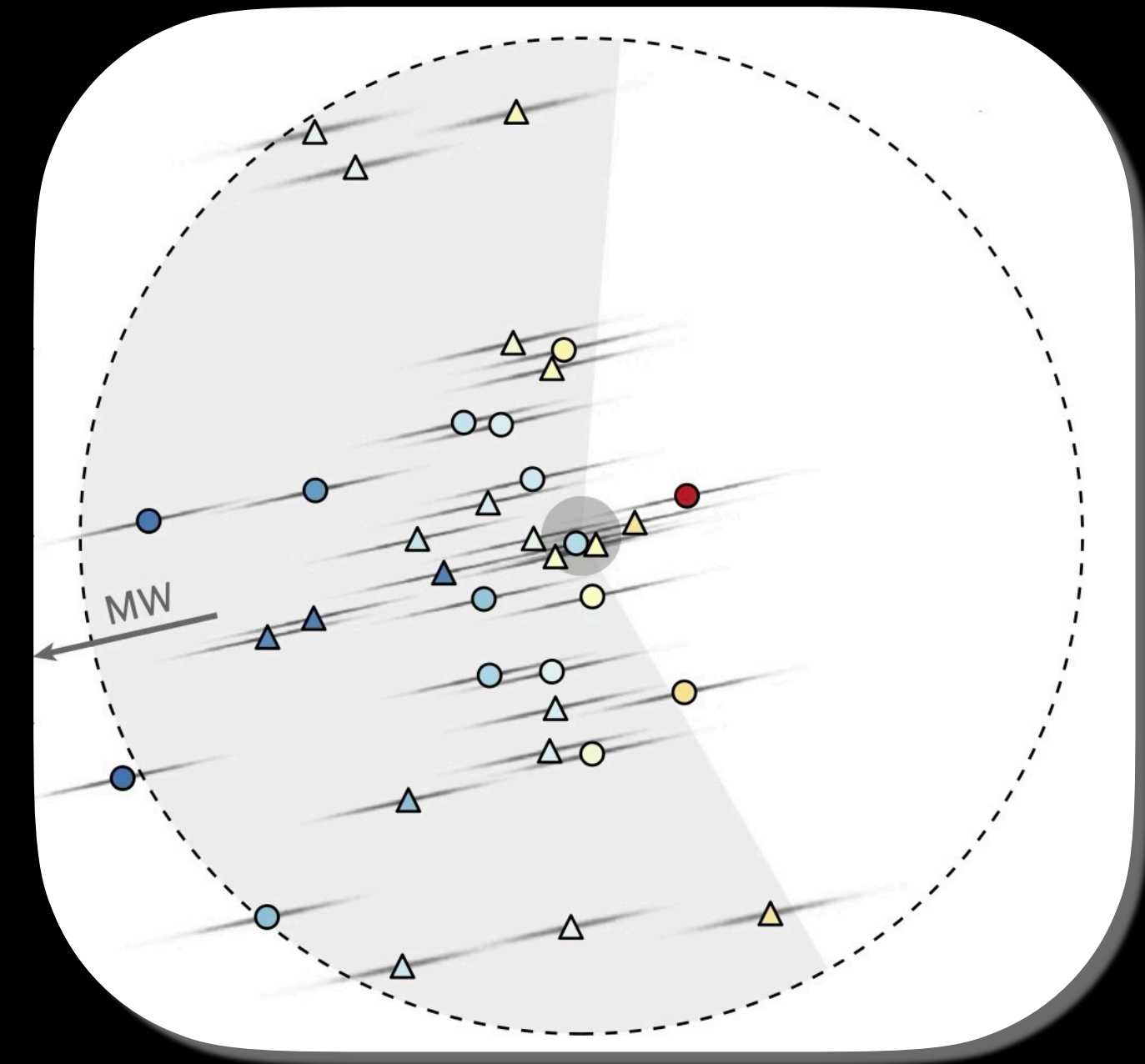
Satellite Plane

Kroupa+2005, Pawlowski+2012, Pawlowski & Kroupa 2013, Pawlowski+2015, Libeskind+2015, Pawlowski 2021, etc.



Satellite Lopsidedness

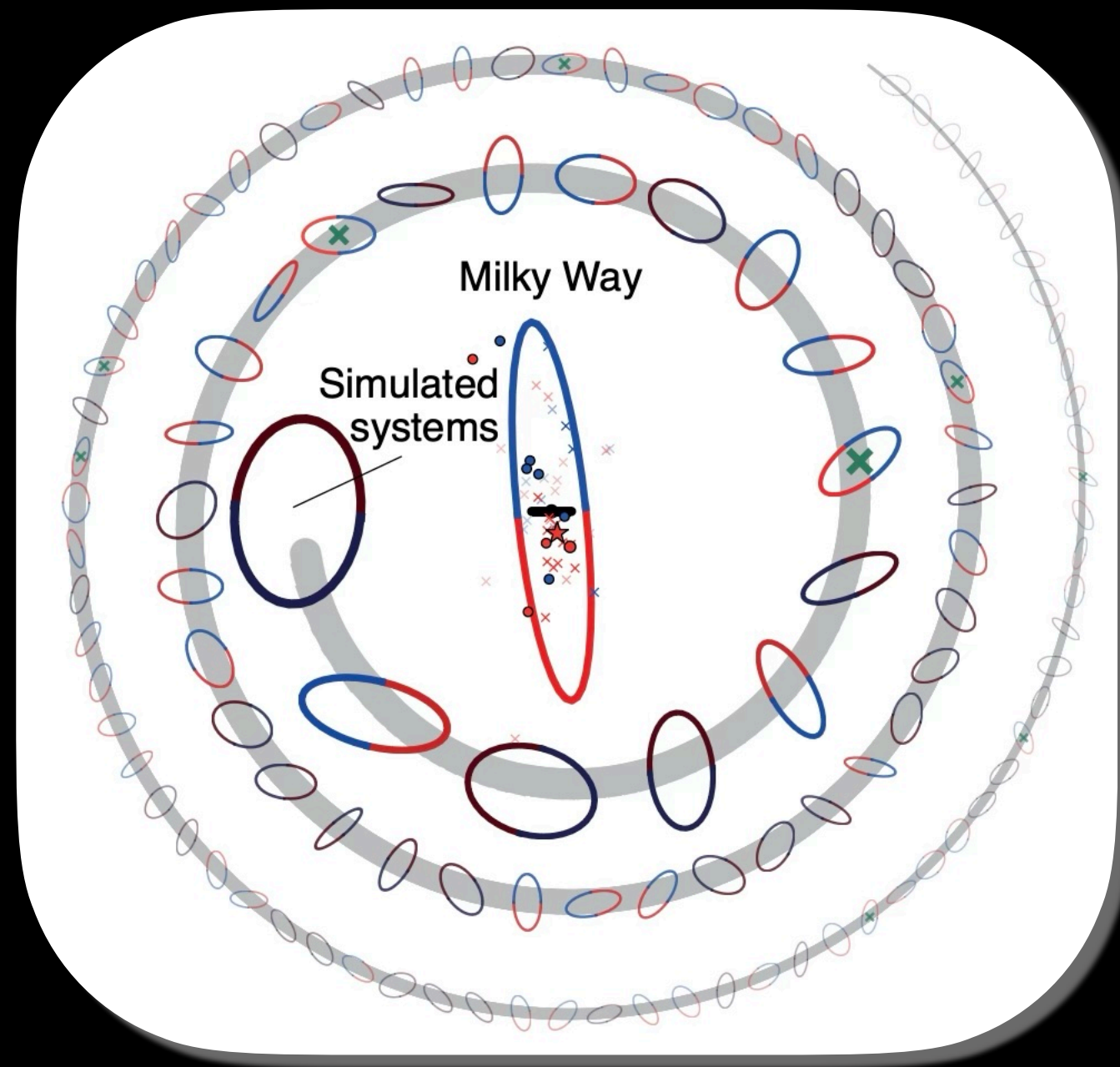
McConnachie & Irwin 2006, Conn+2012, Savino+2022, Kanehisa+2025, etc.



- **New universe crisis from our Local Group**

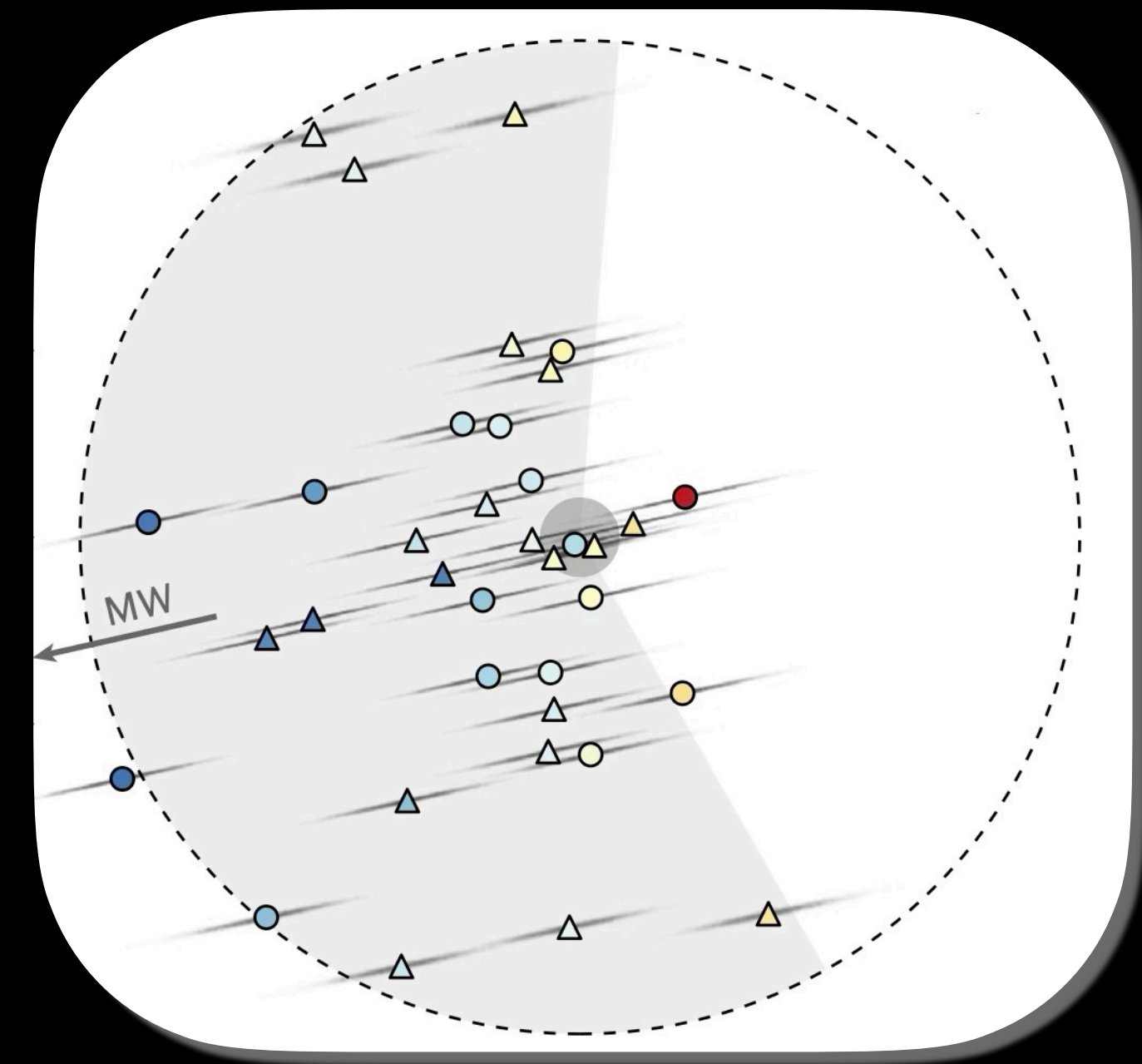
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Satellite Lopsidedness

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These peculiar properties of our Local Group made it either

- a **challenge** to LCDM cosmology and our galaxy formation paradigm
- or just an unfortunate **outlier**

- **Our Local Group shaped by the large-scale cosmic web**

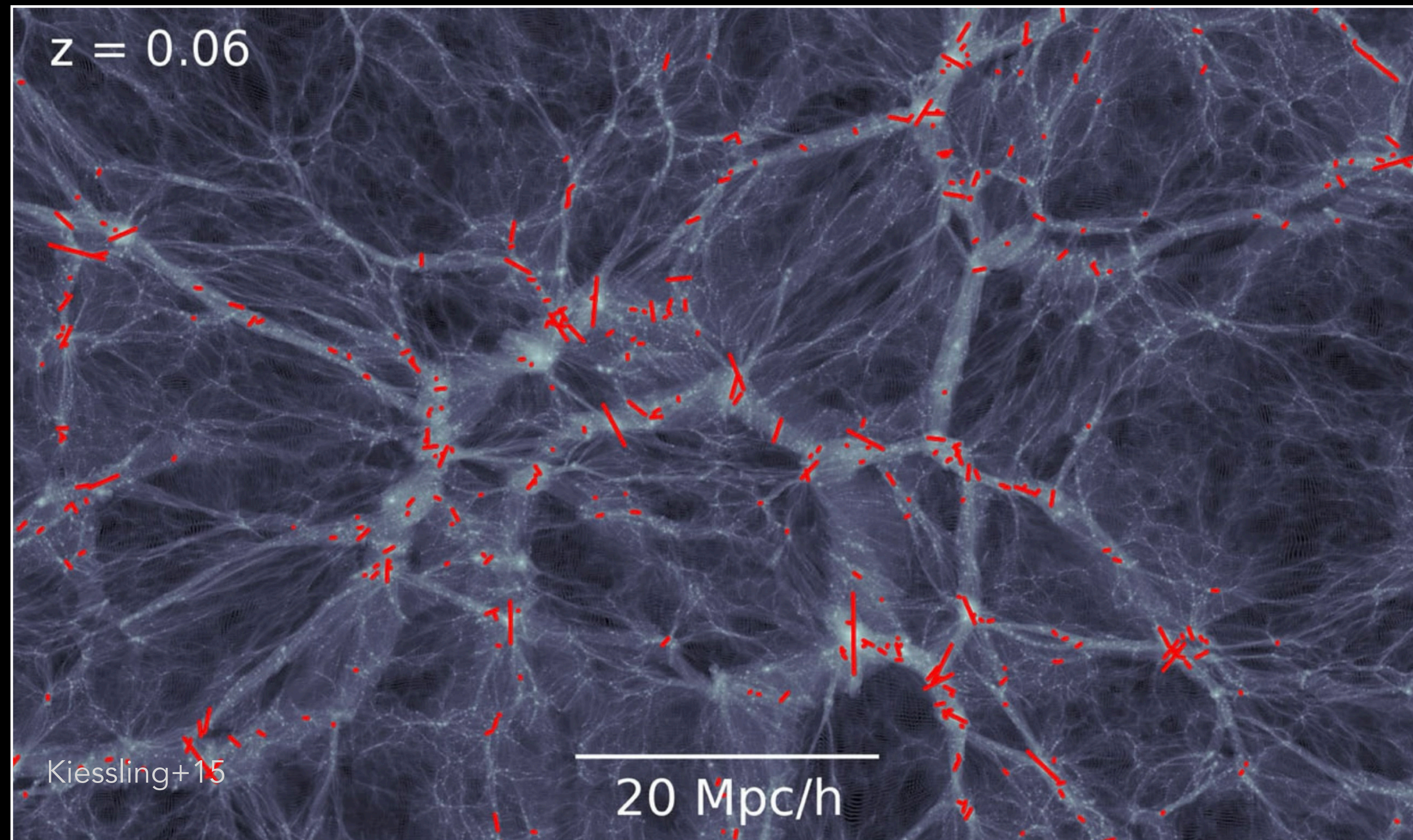
There are tons of evidence that

small-scale structures are related to **Large-scale Cosmic Web**.

- **Our Local Group shaped by the large-scale cosmic web**

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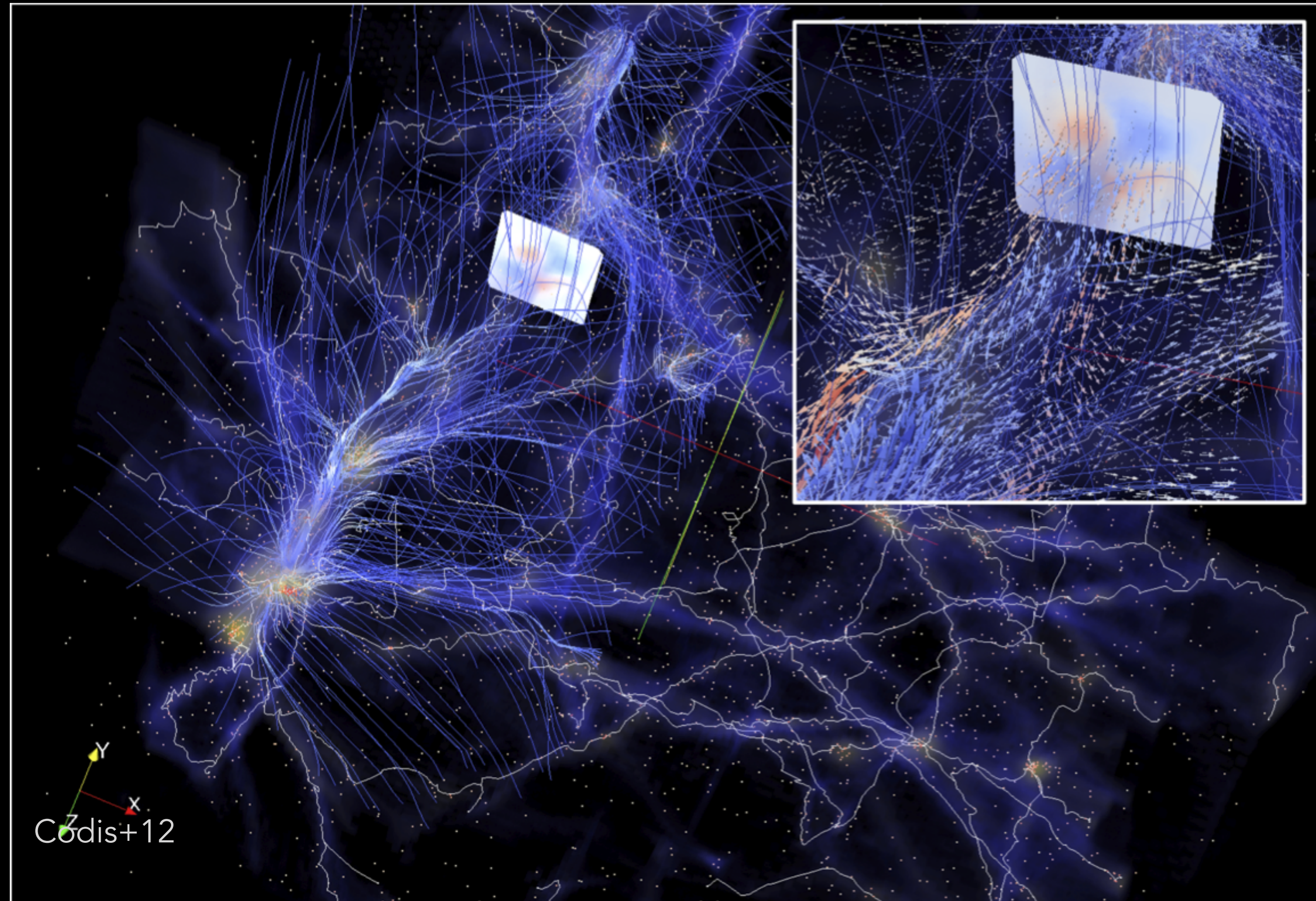
~~small scale structures~~ are related to **Large-scale Cosmic Web.**
galaxy alignment



- Our Local Group shaped by the large-scale cosmic web

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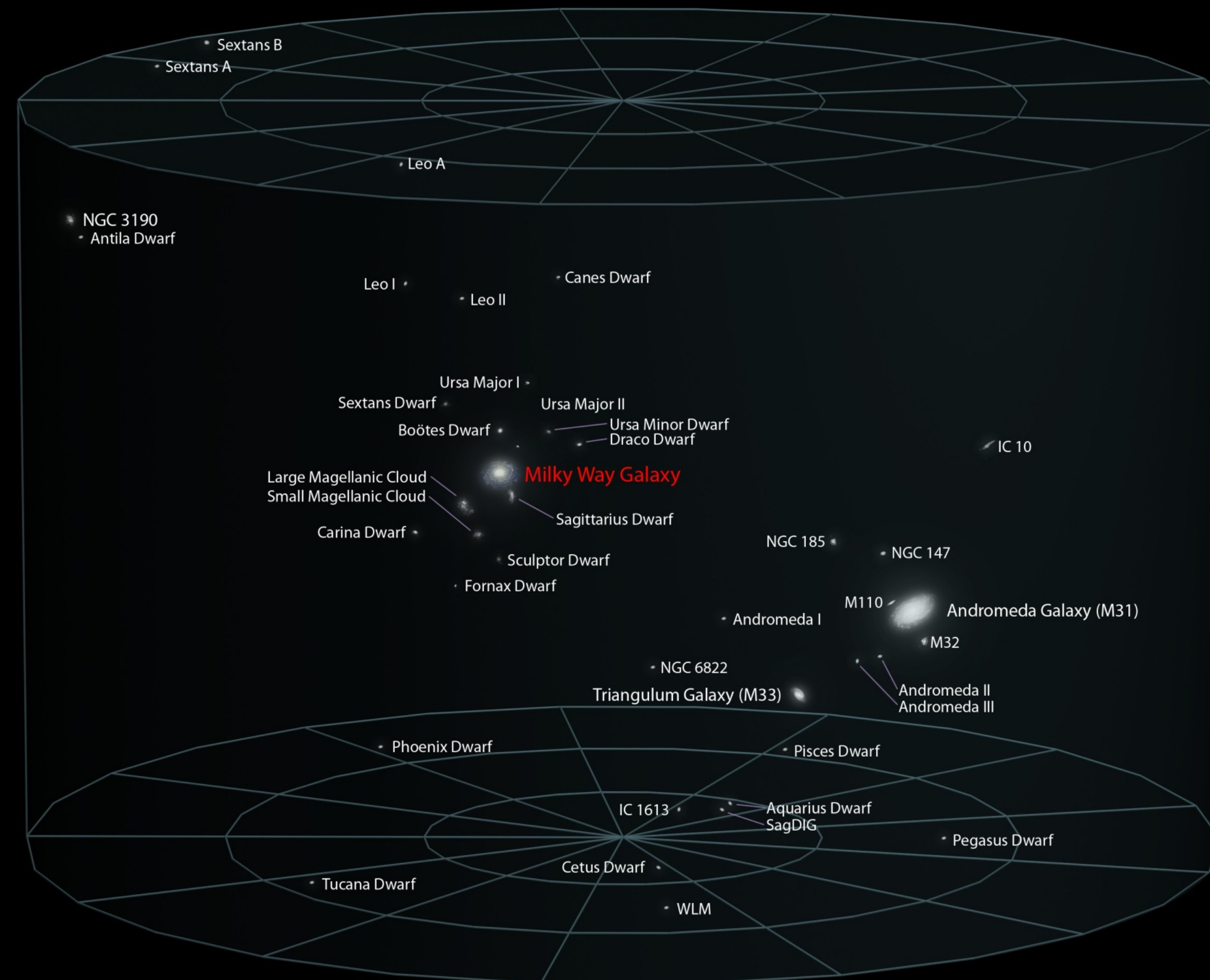
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halo spin



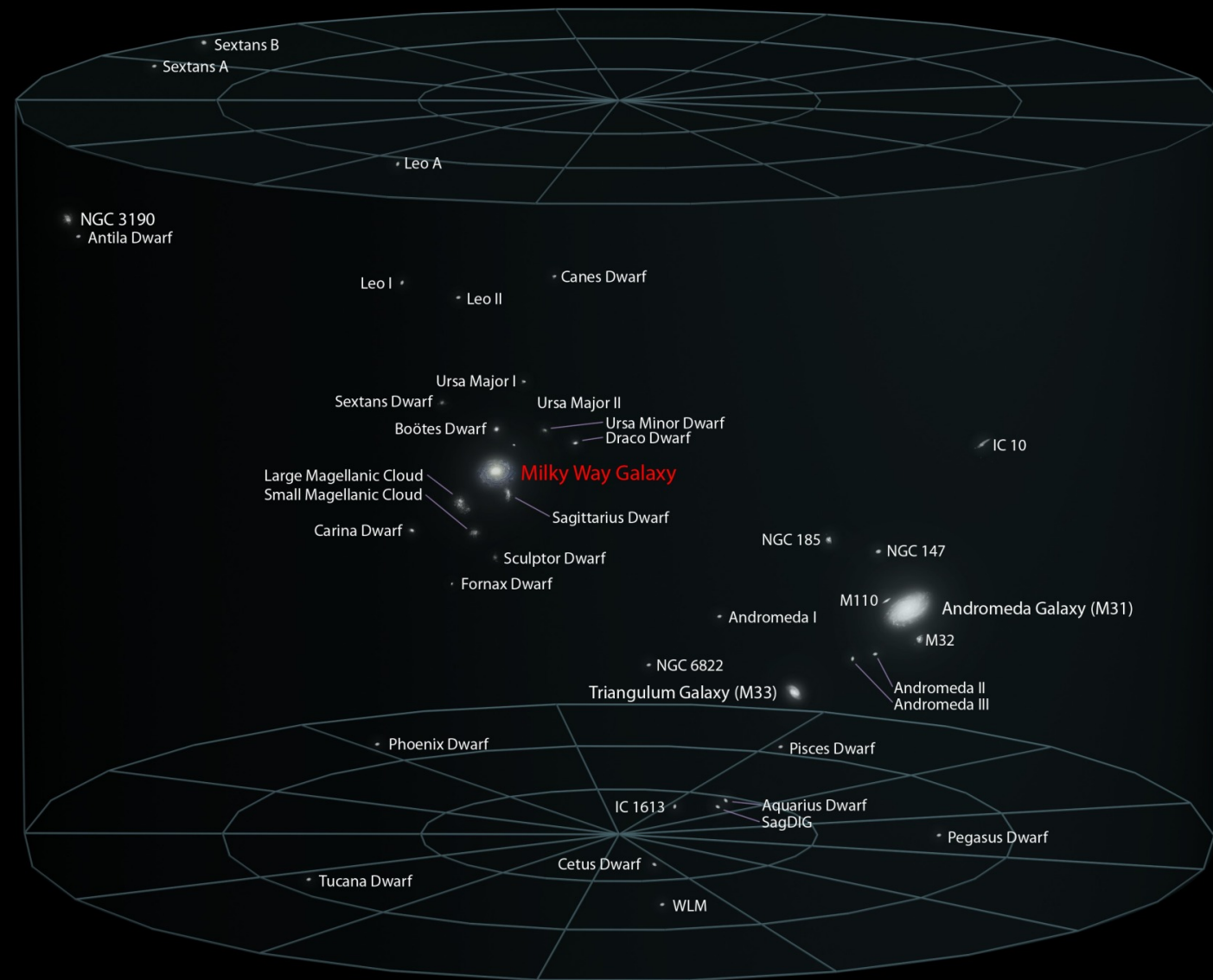
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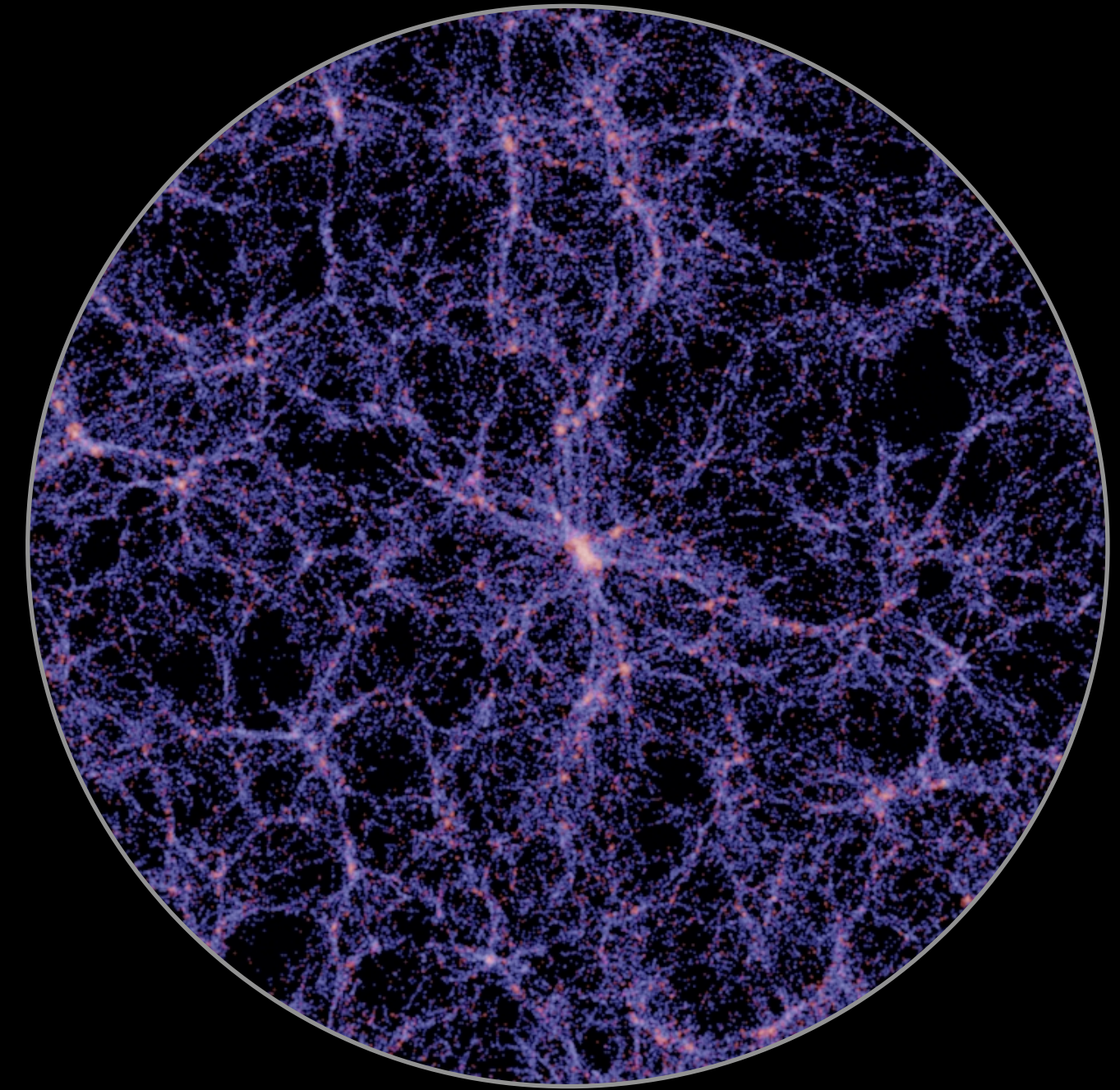
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Local Group Analogue?



- **Our Local Group shaped by the large-scale cosmic web**



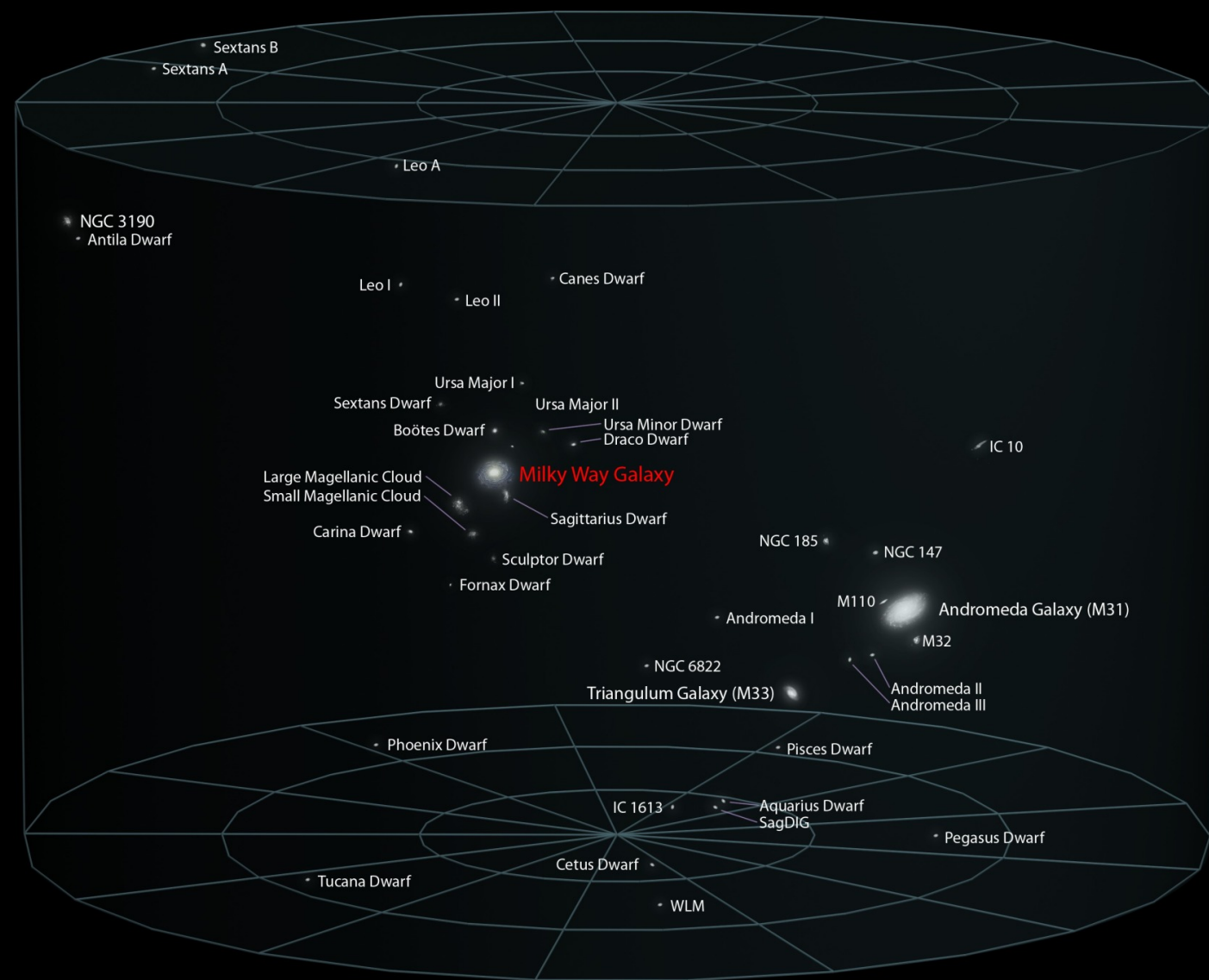
Local Group Analogue



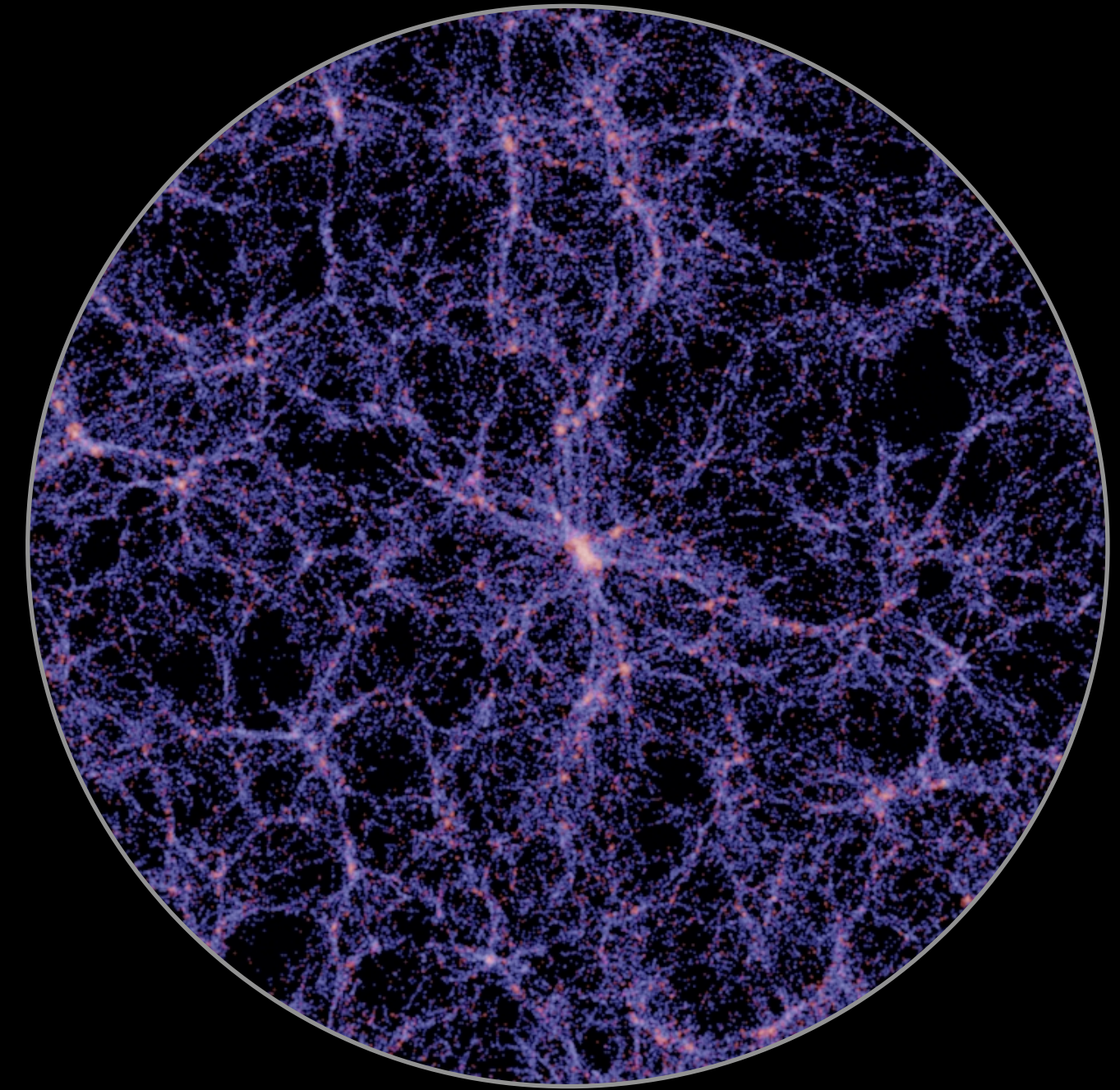
Large-scale Cosmic Web

Our peculiar Local Group caused by **peculiar Large-scale environment?**

- Our Local Group shaped by the large-scale cosmic web



Local Group Analogue



Large-scale Cosmic Web

Yes! We are living in a peculiar large-scale environment.
(see Libeskind+2015)

Our peculiar Local Group caused by peculiar Large-scale environment?

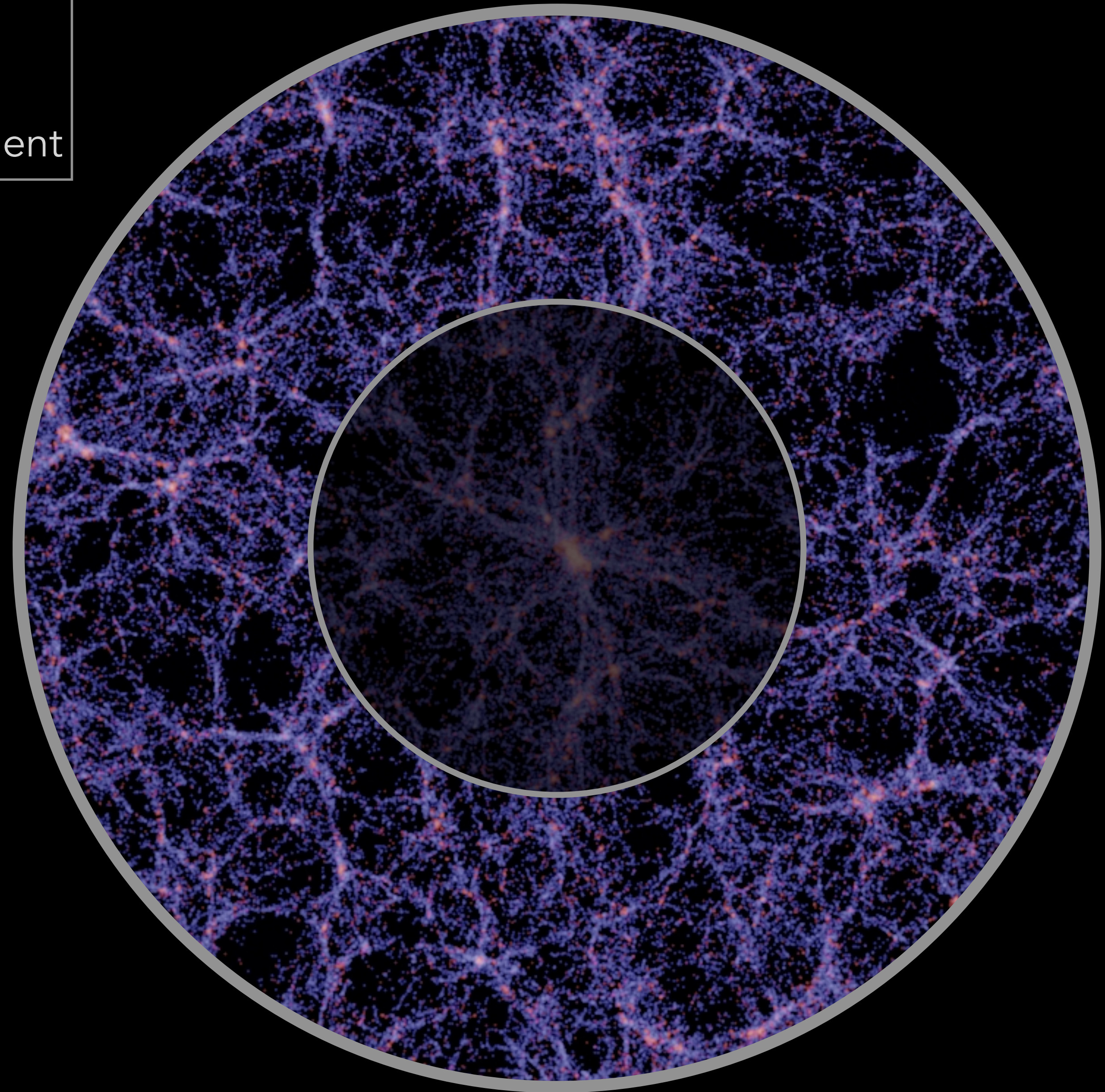
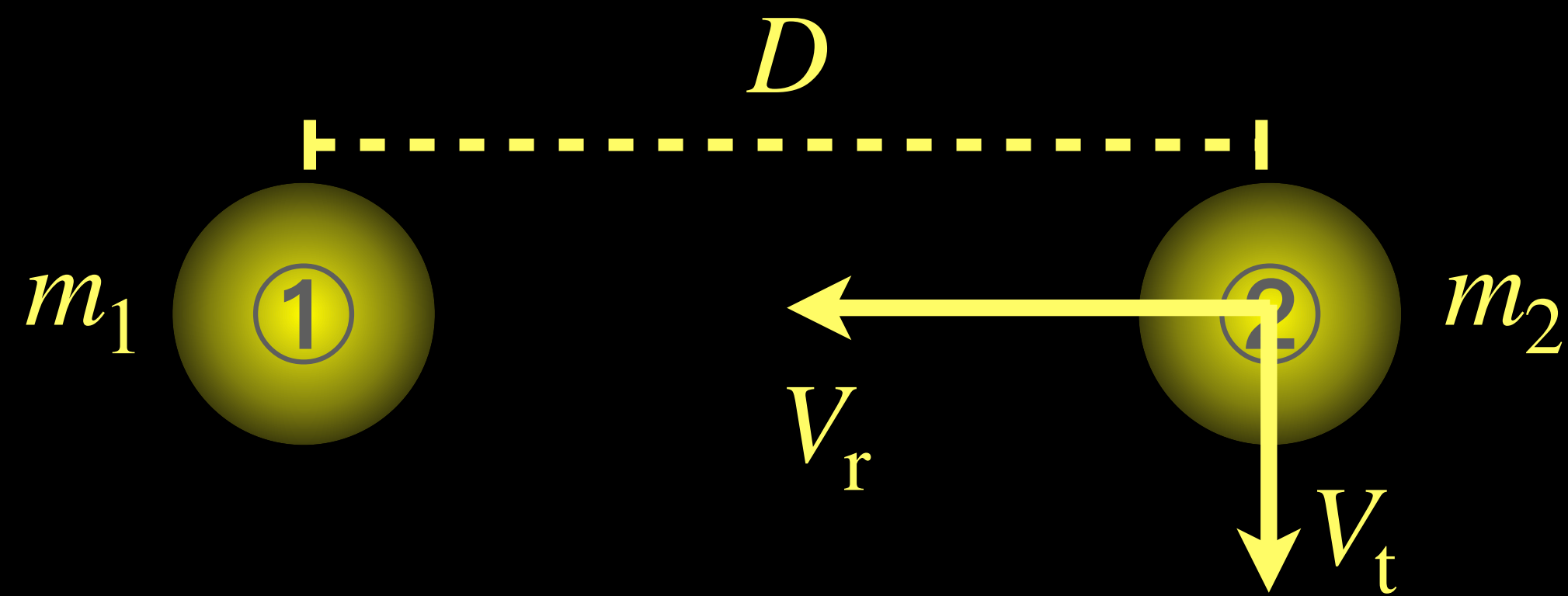
- Quantifying **Local Group Analogue** and **Cosmic Web**

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- two $10^{12}M_{\text{sun}}/h$ halos
- separated by $\sim 500\text{kpc}/h$
- in a relatively isolated environment

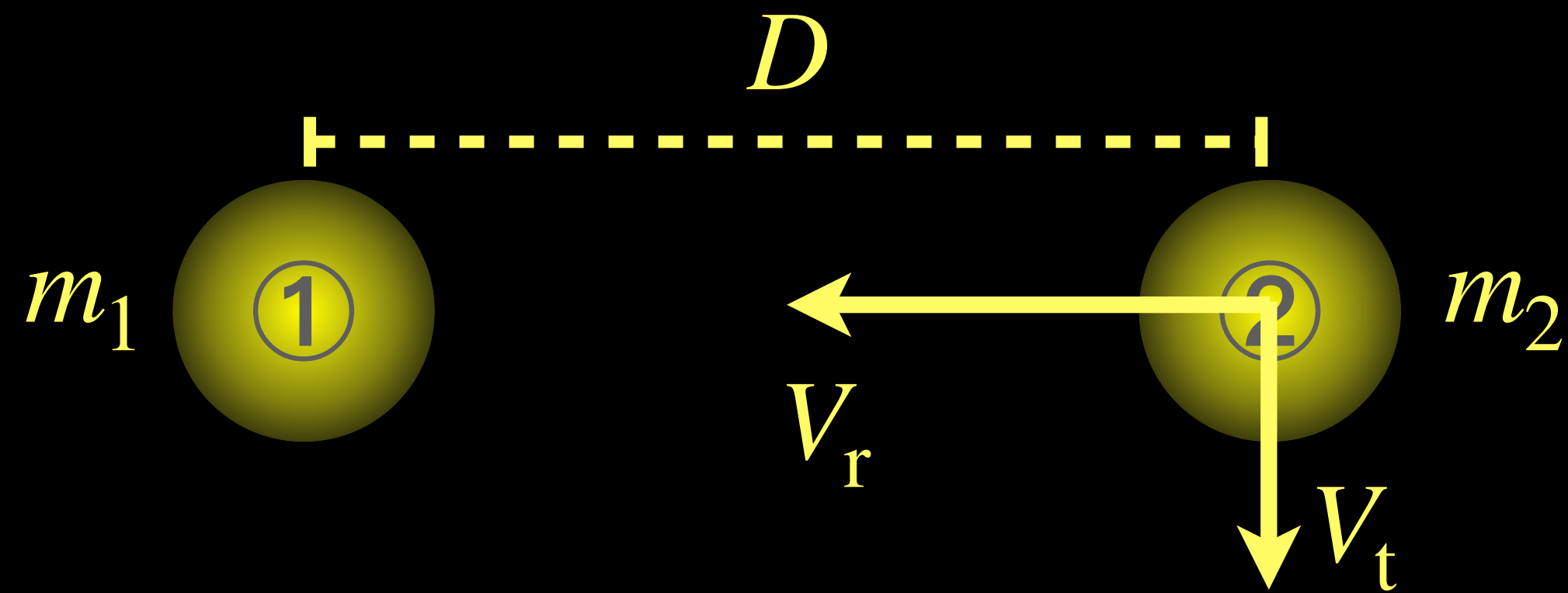
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Local Group Analogues property:

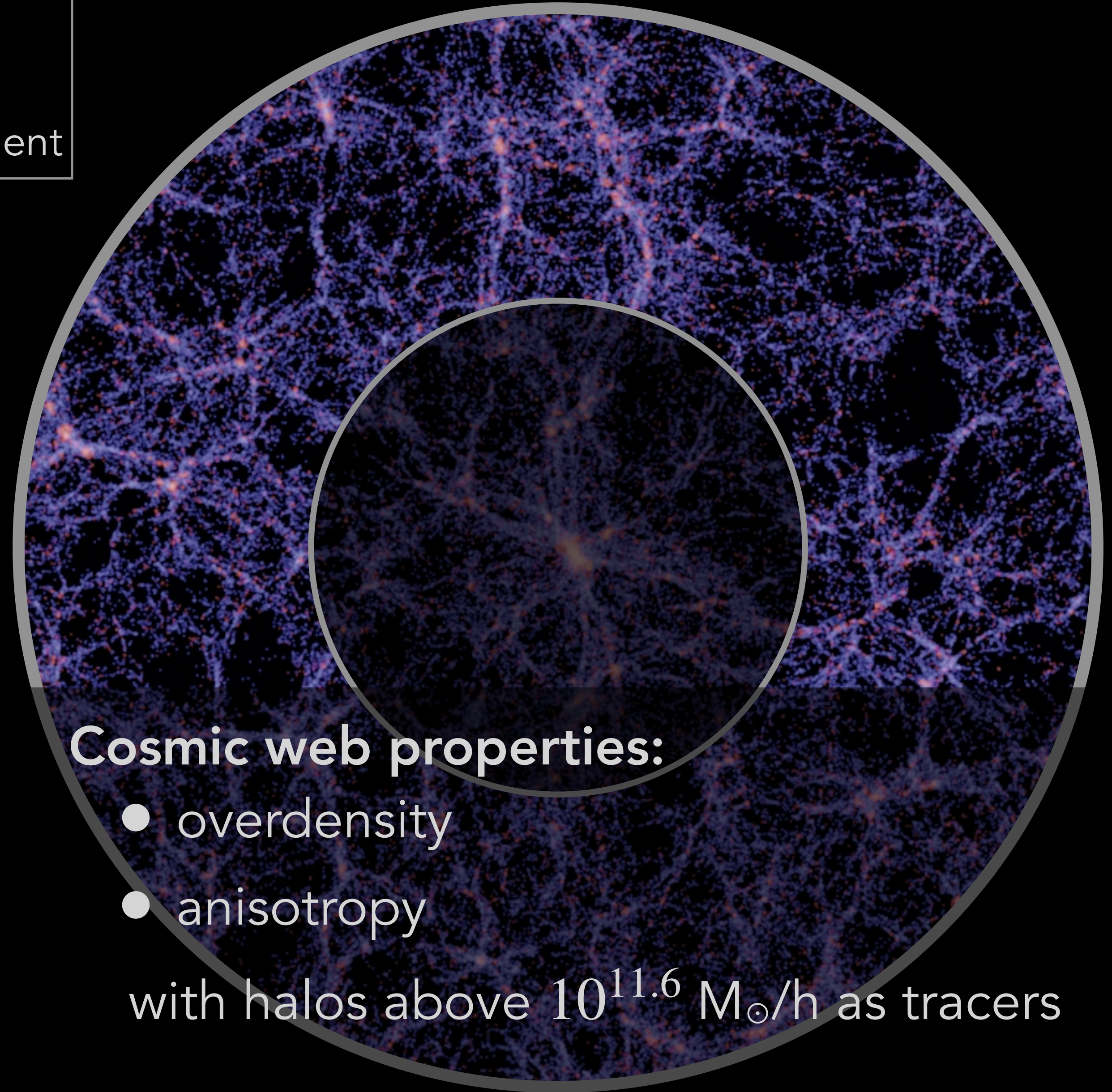
- Gravitational coupling energy

$$E_{\text{coupling}} = \frac{1}{2} \frac{m_1 m_2}{m_1 + m_2} V^2 - \frac{G m_1 m_2}{D}$$

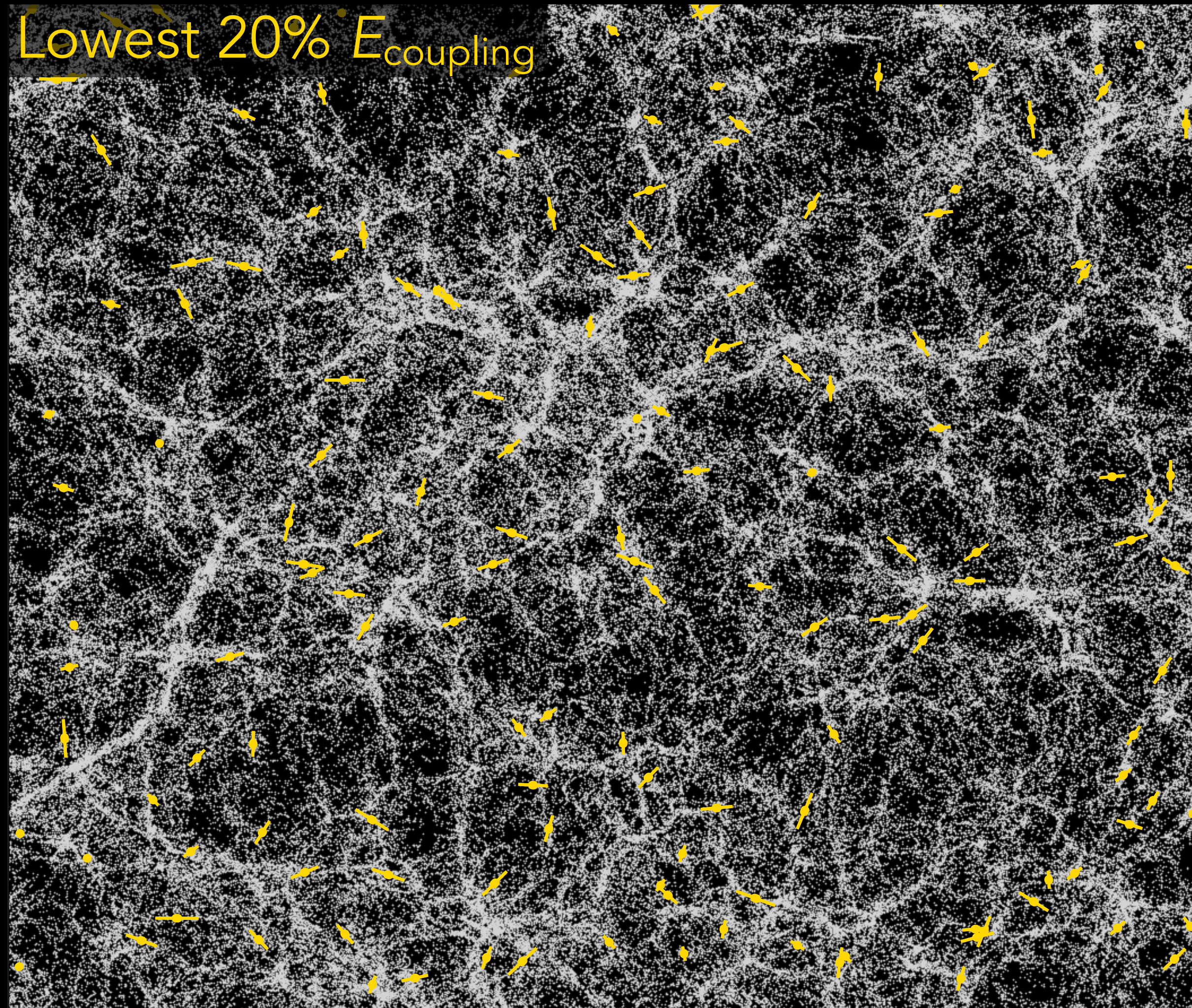
Cosmic web properties:

- overdensity
- anisotropy

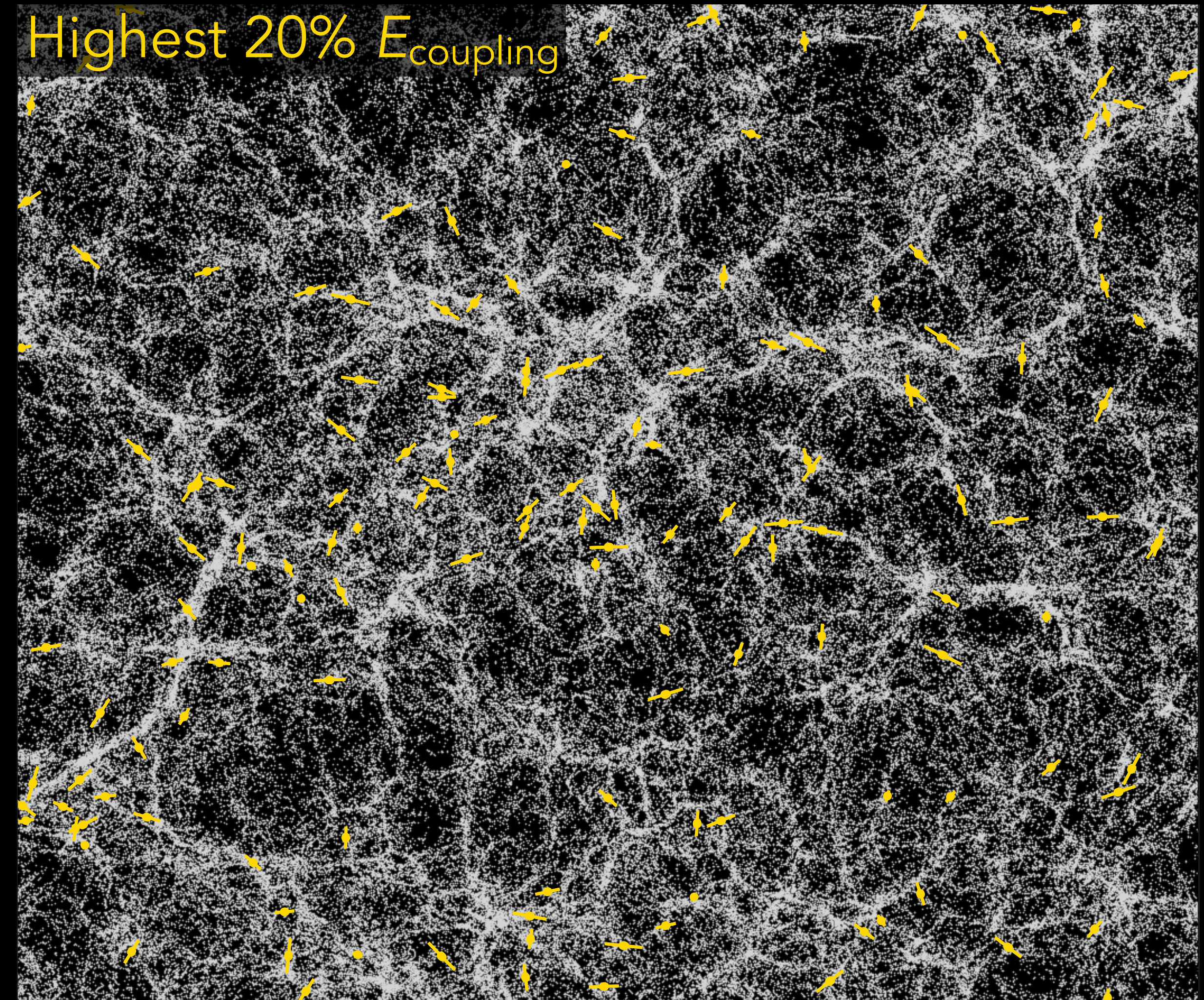
with halos above $10^{11.6} M_{\odot}/h$ as tracers



- **Correlations between Local Group Analogue and Cosmic Web**

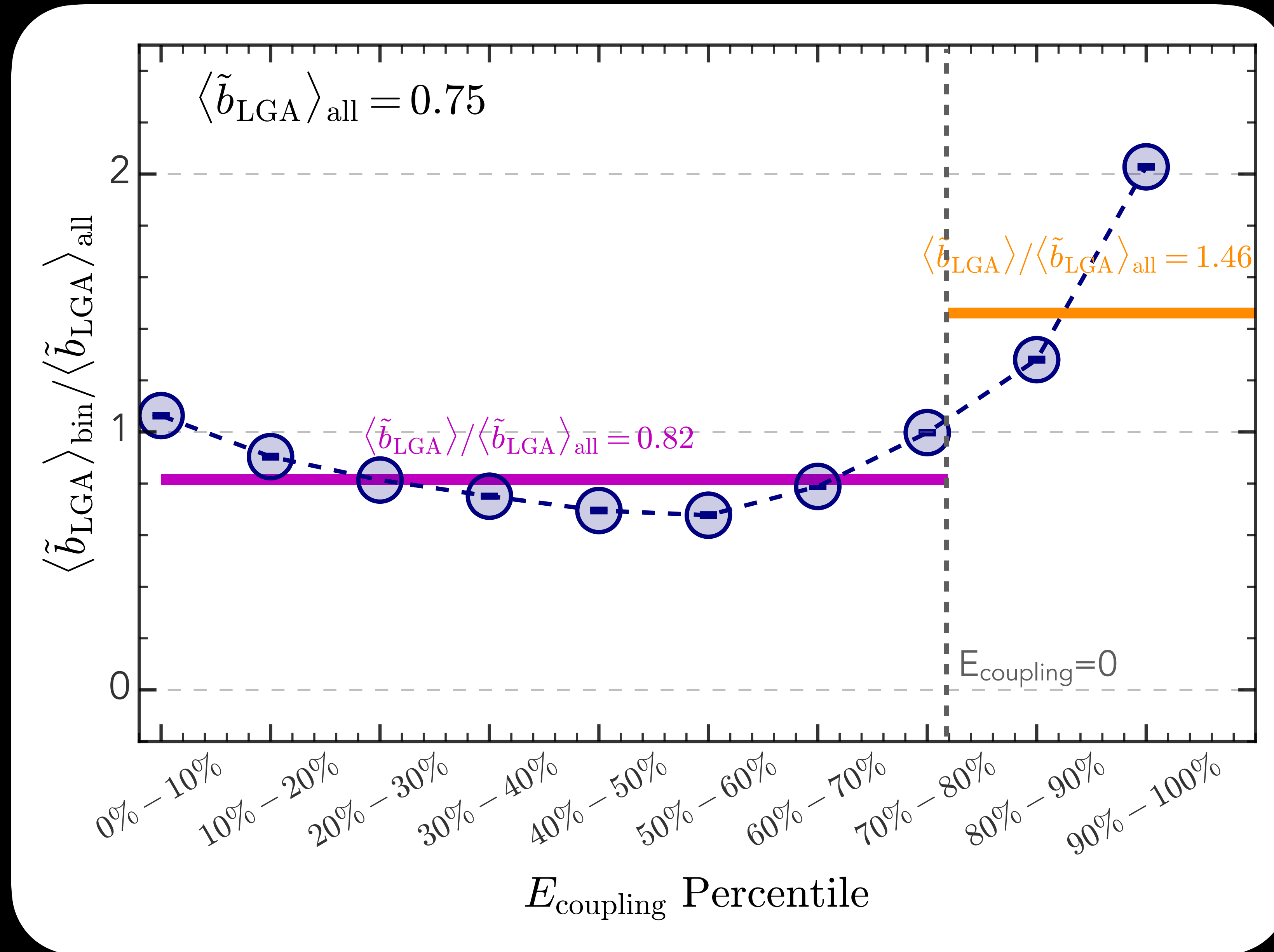


Weak clustering

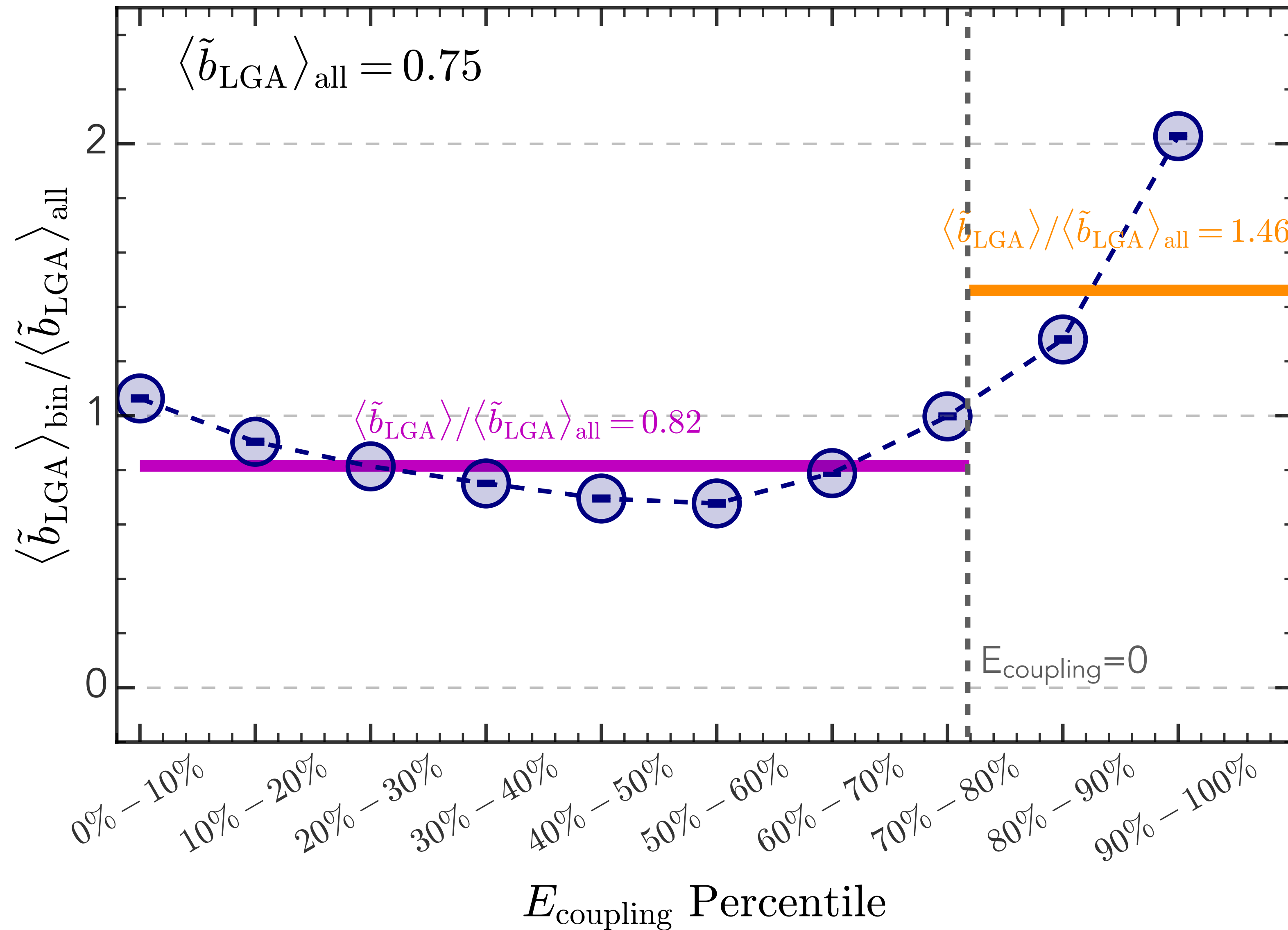


Strong clustering

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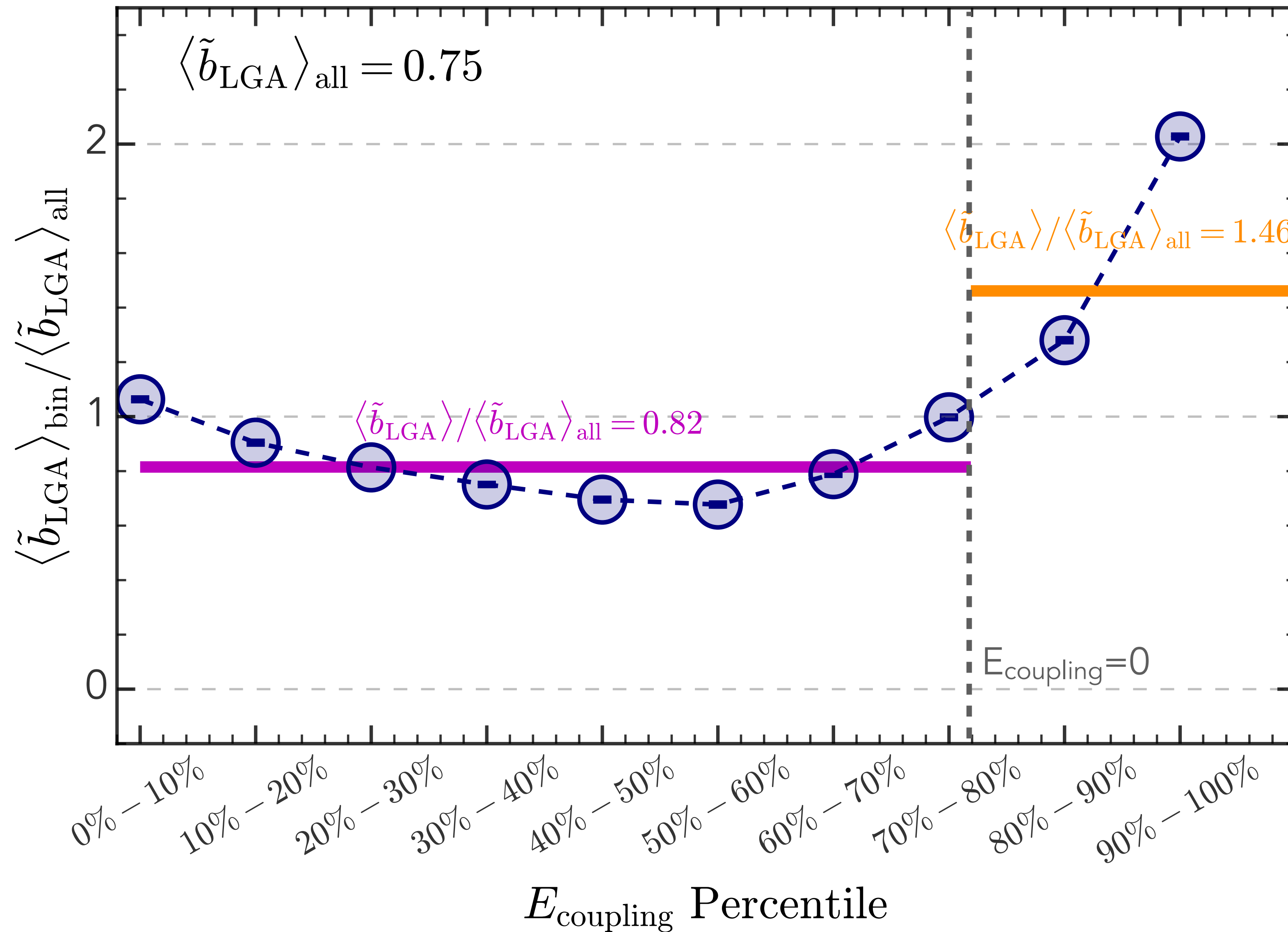


- Correlations between Local Group Analogue and Cosmic Web



~80%

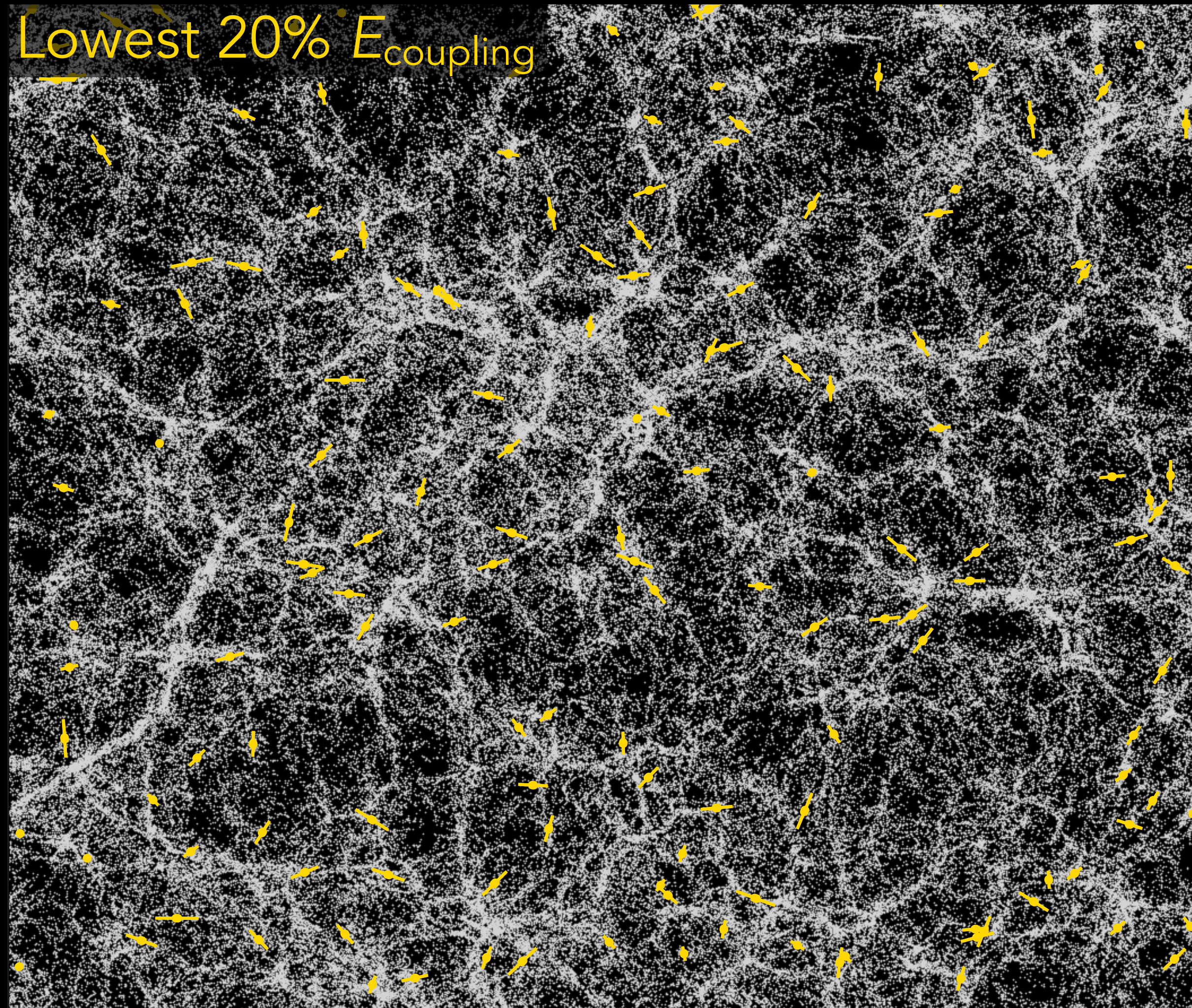
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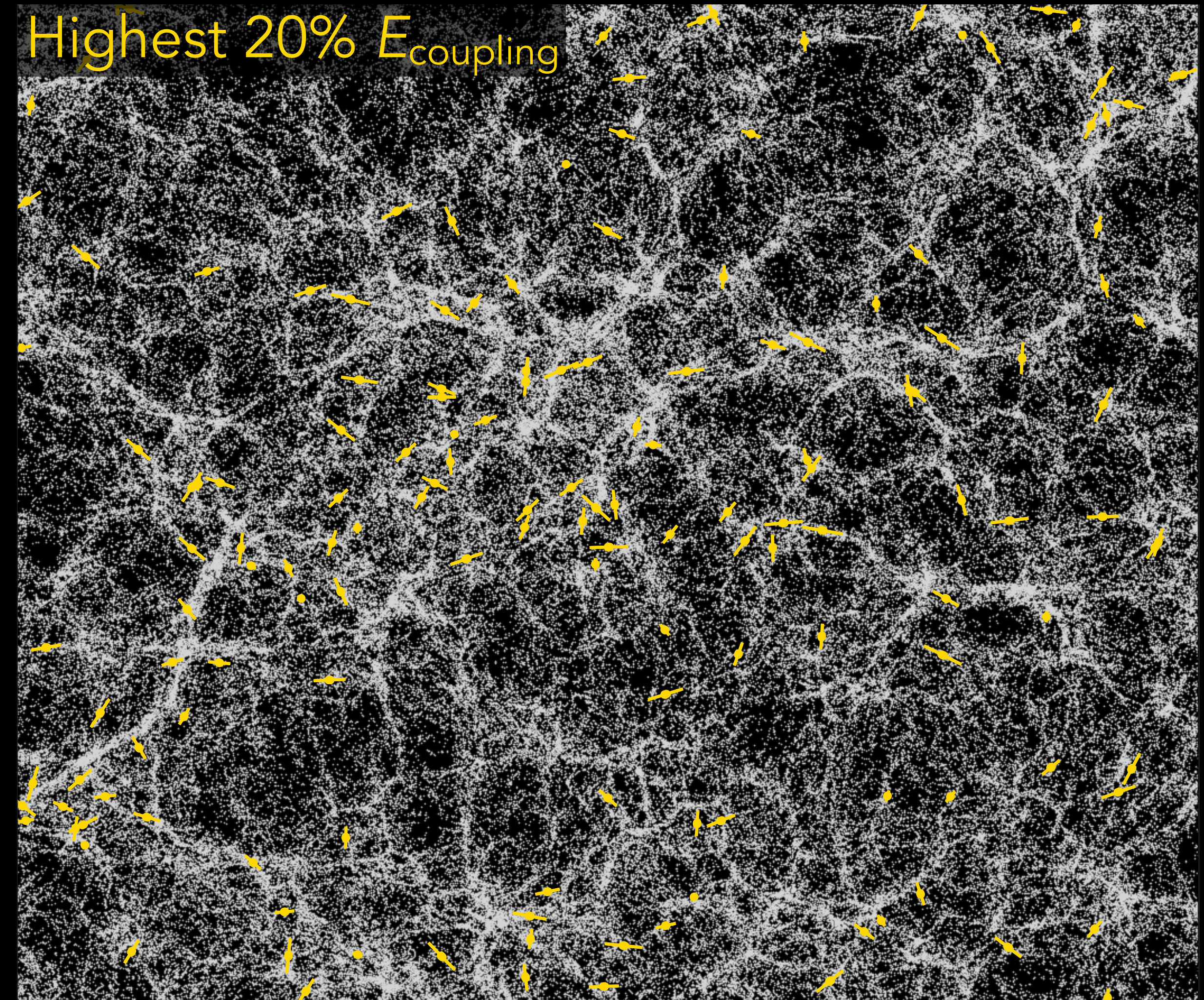
~80%

Halo assembly bias for formation time only differs by ~60%.

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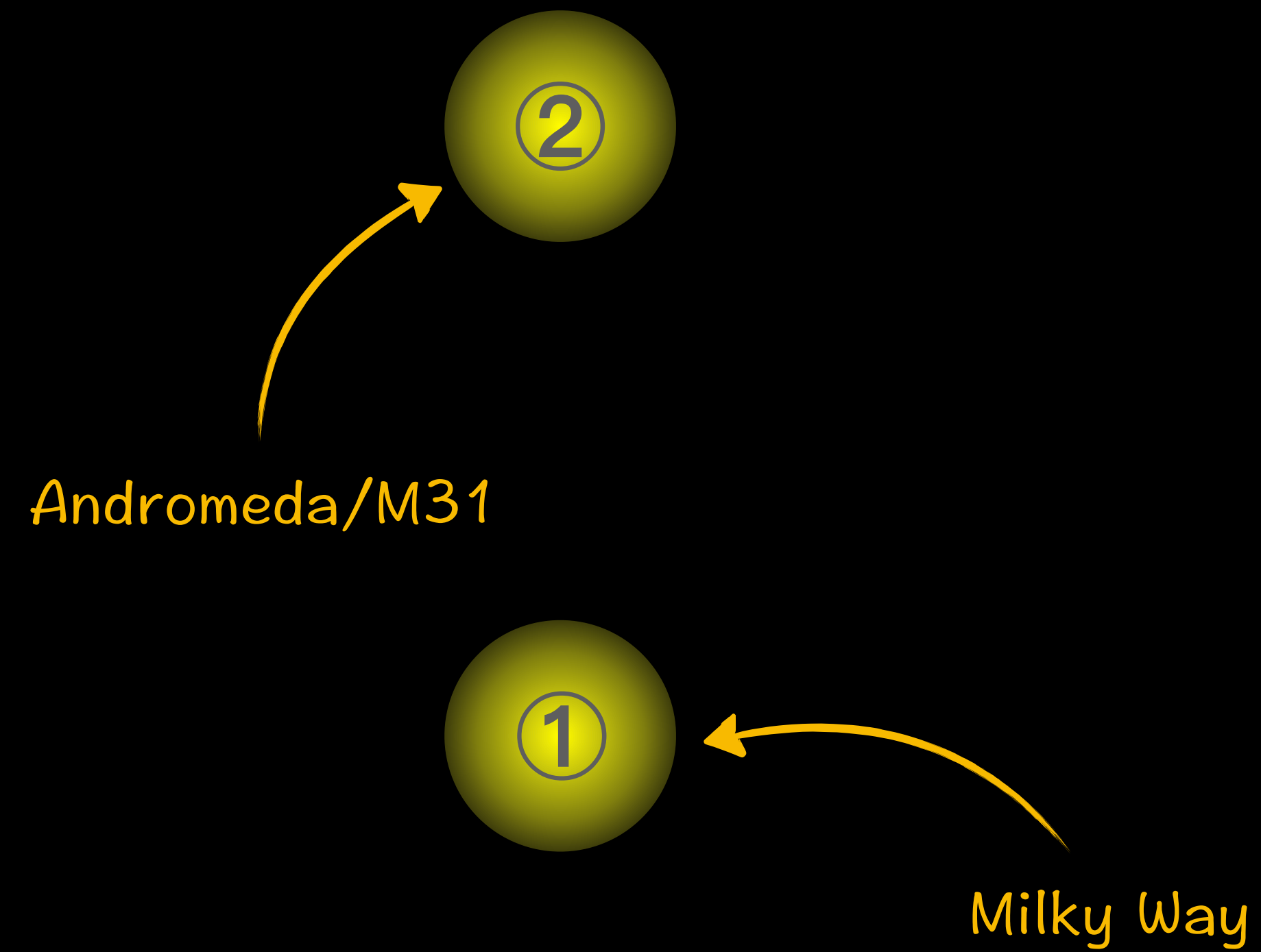


Aligned to filament

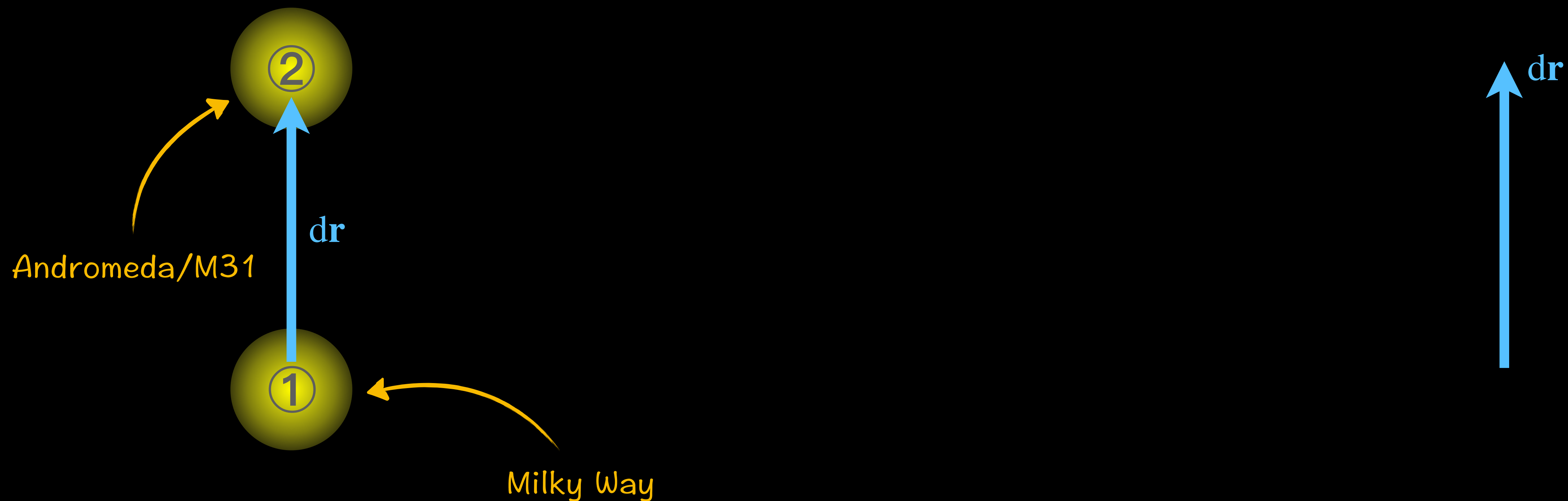


Random alignment

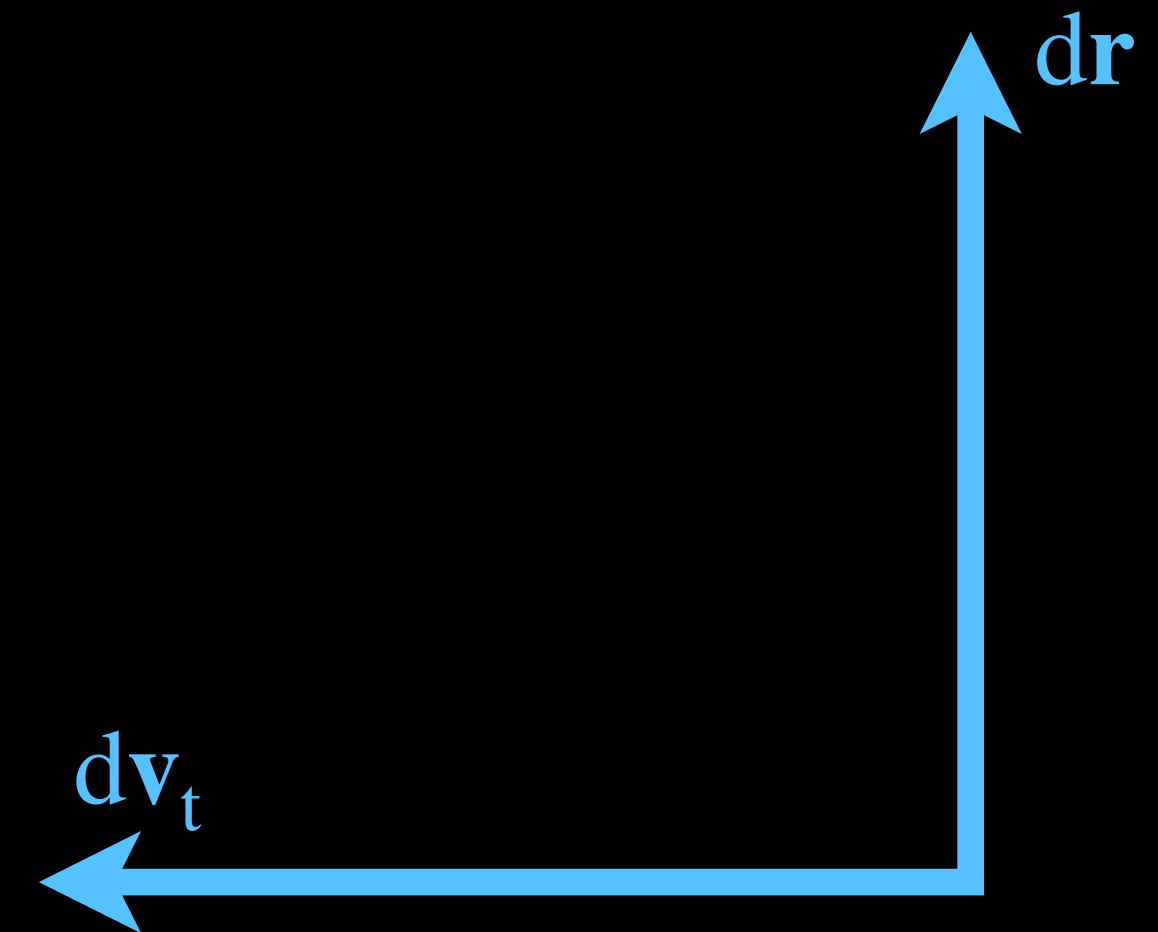
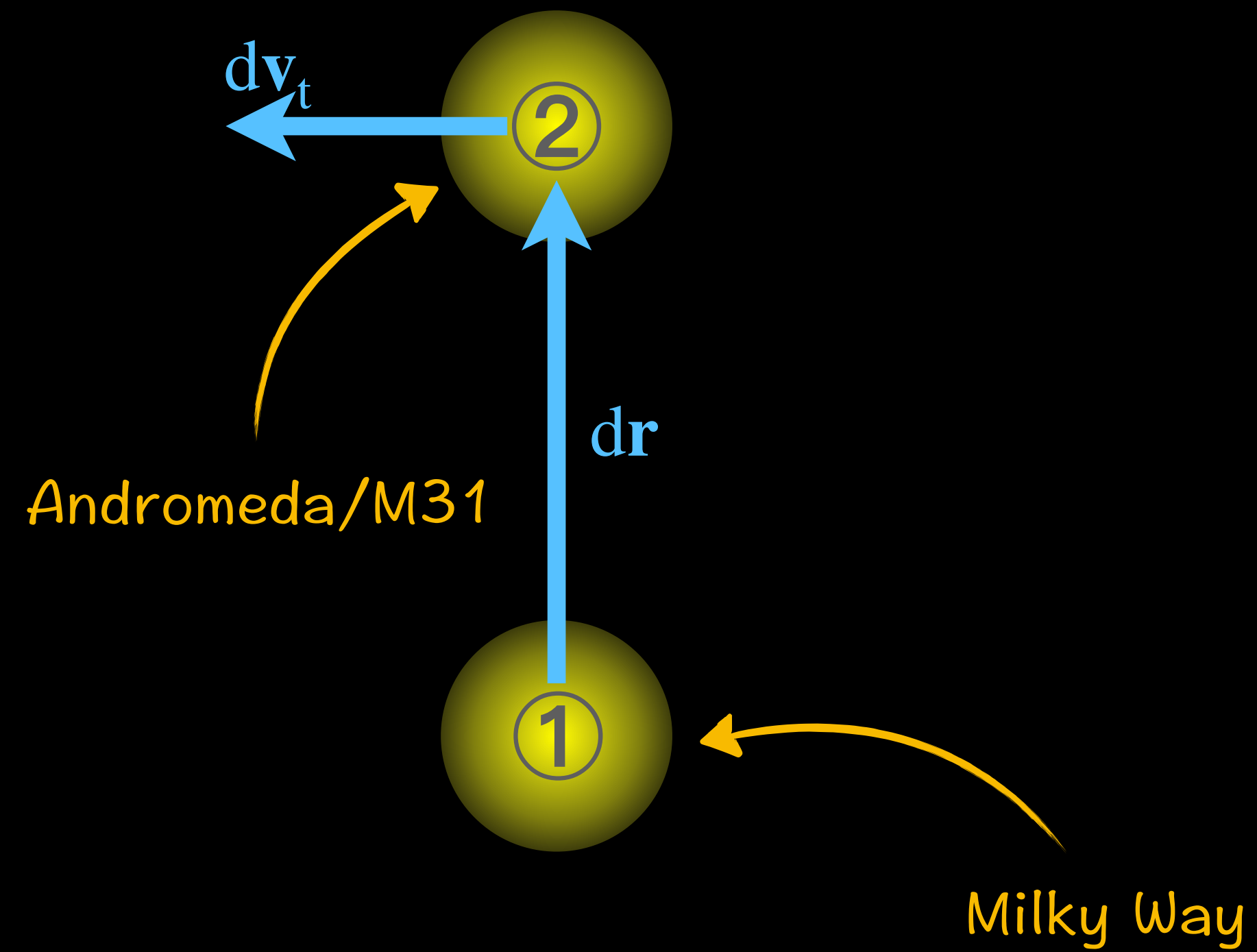
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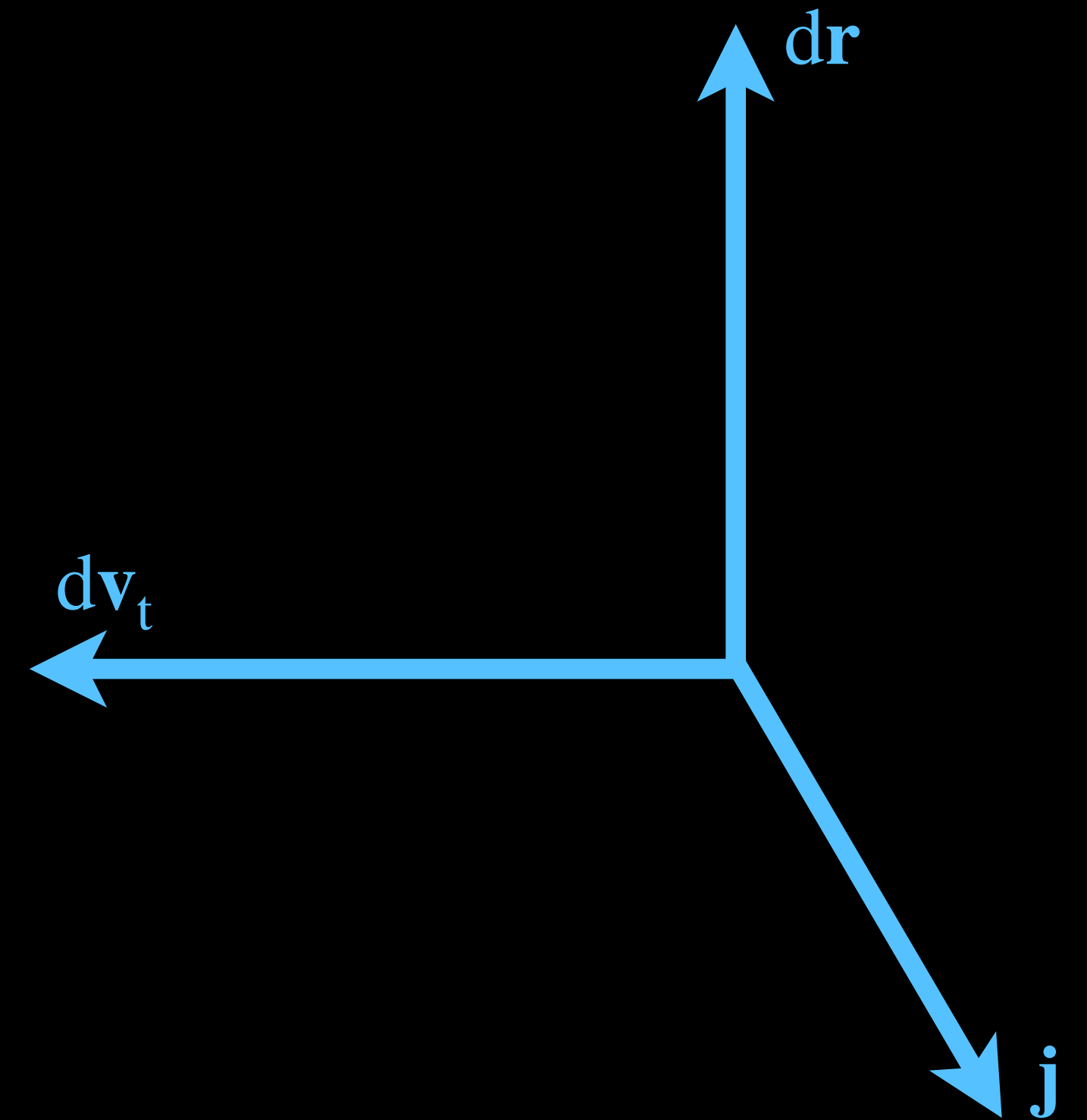
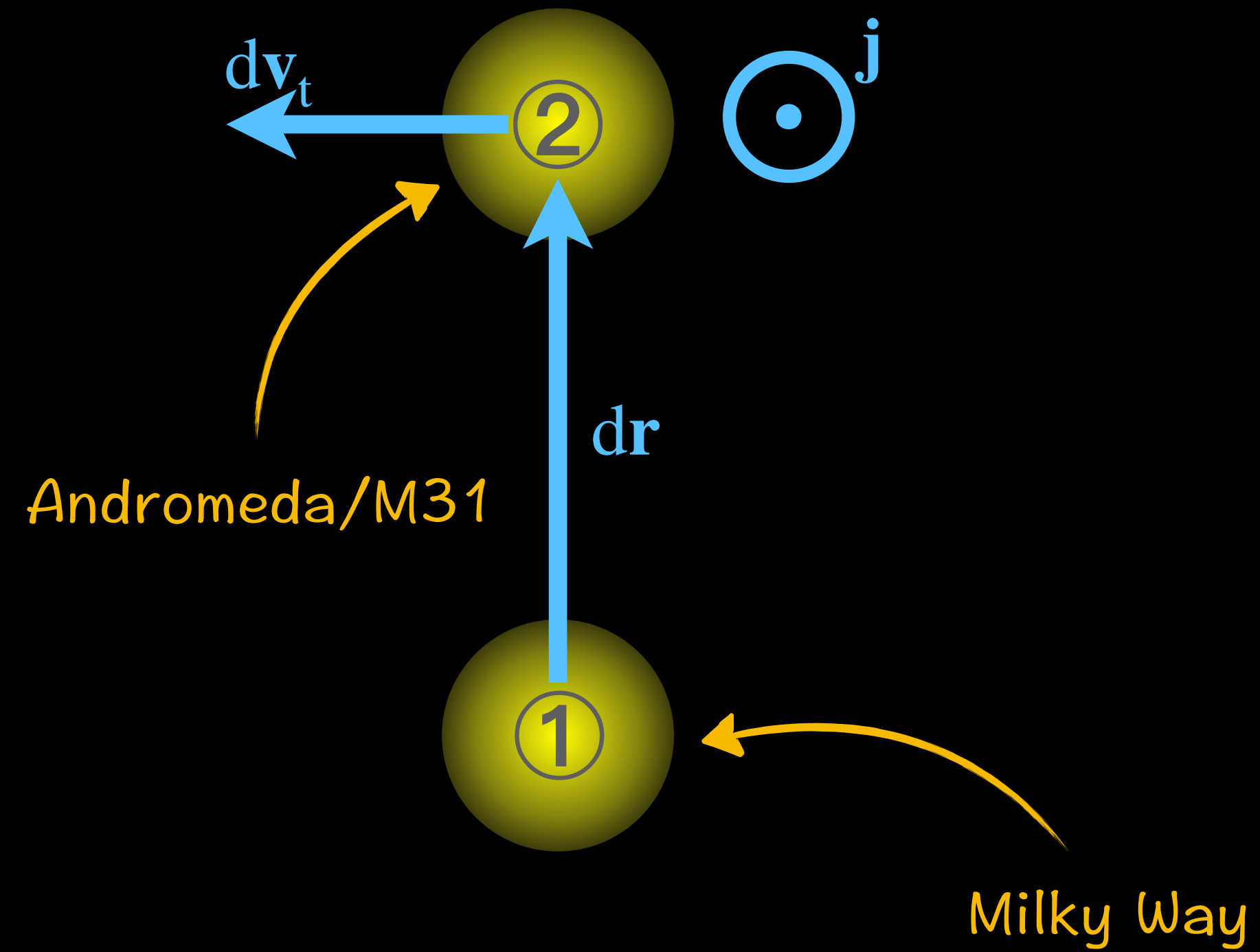
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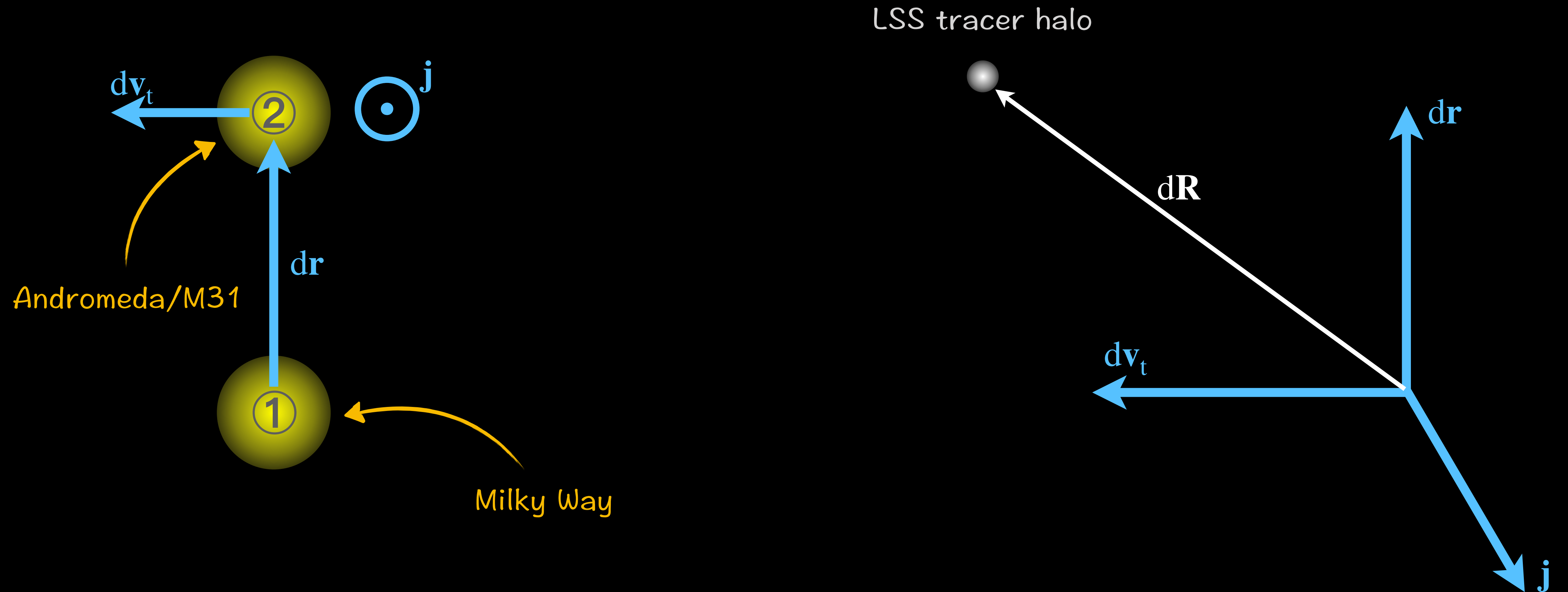
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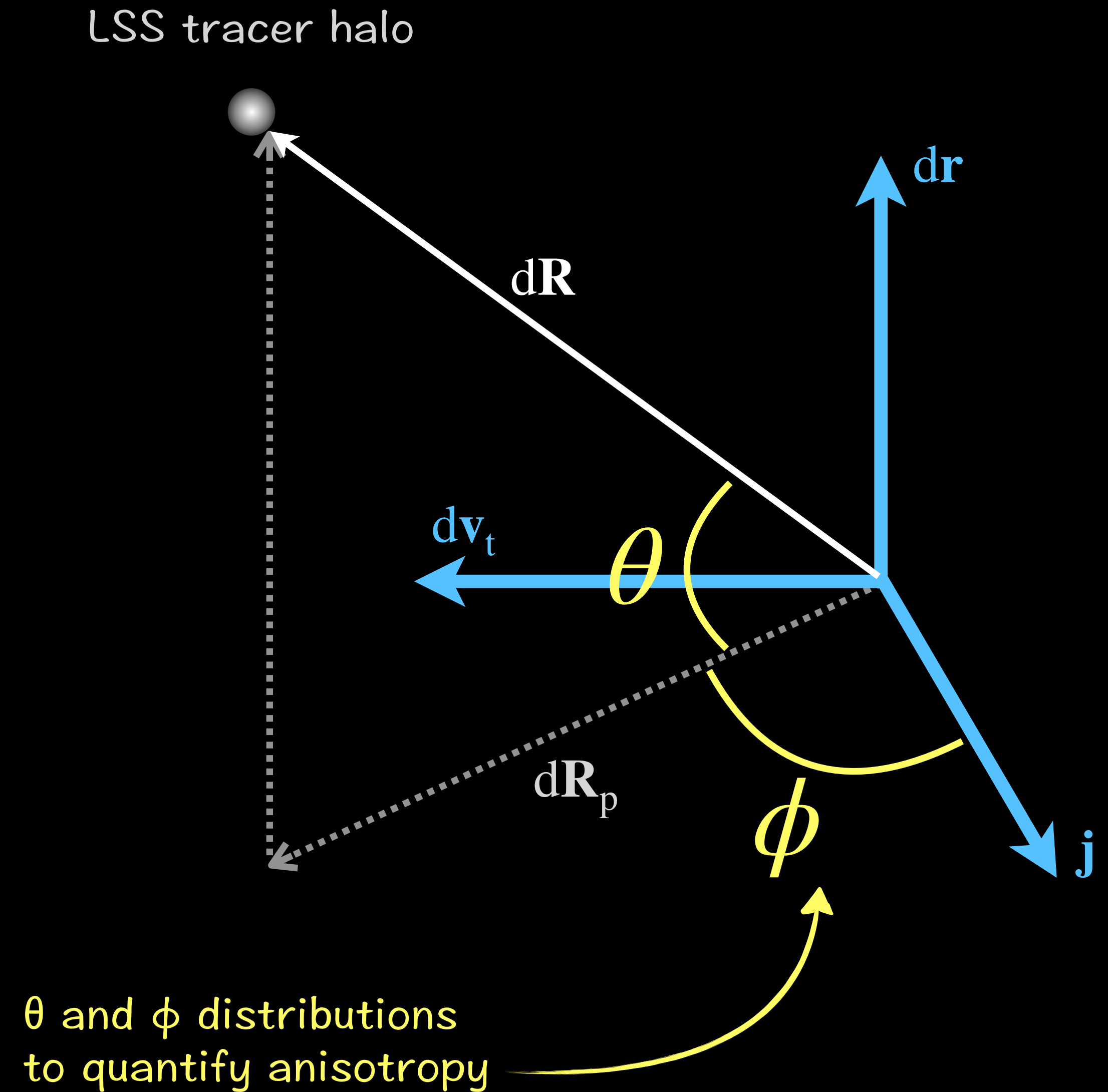
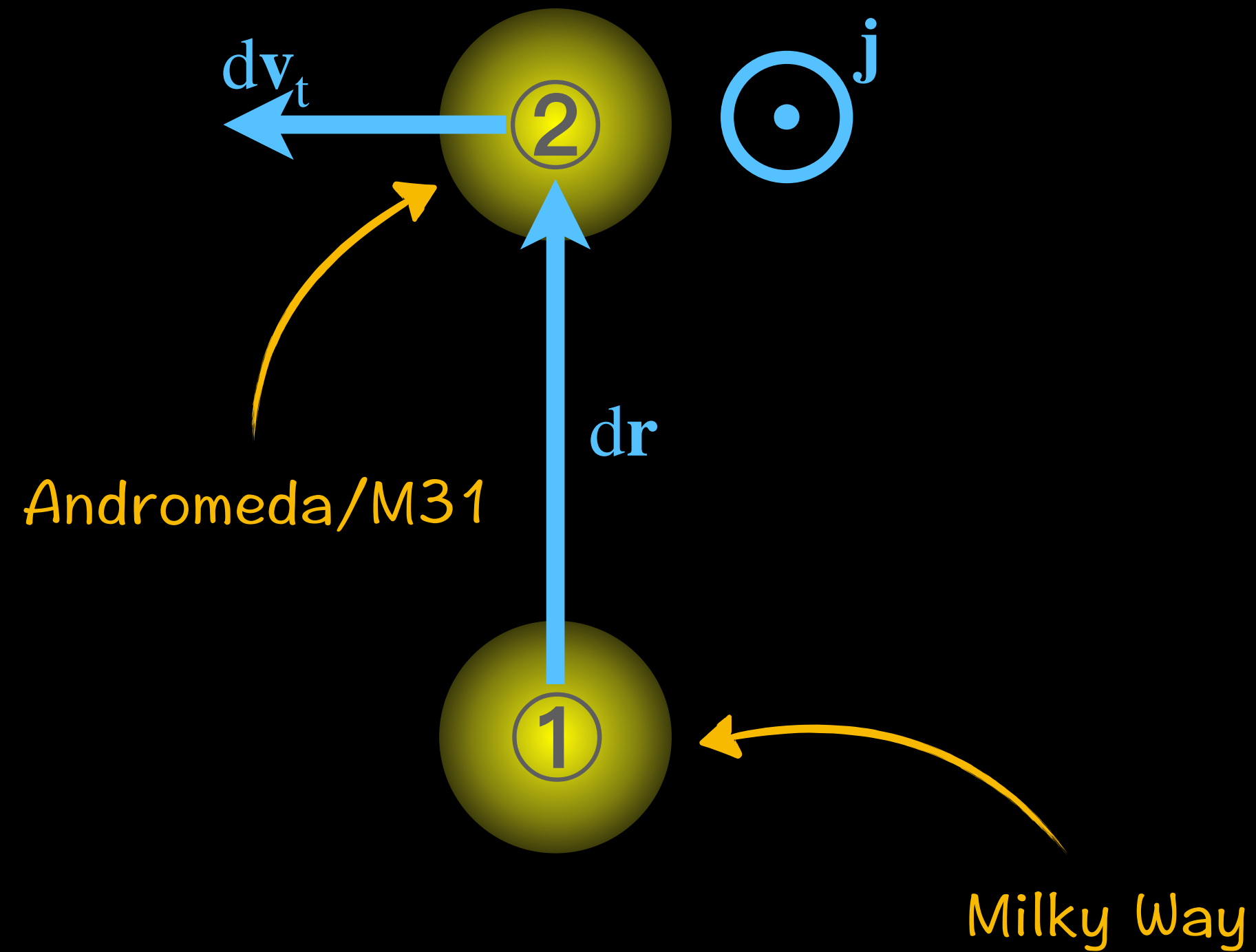
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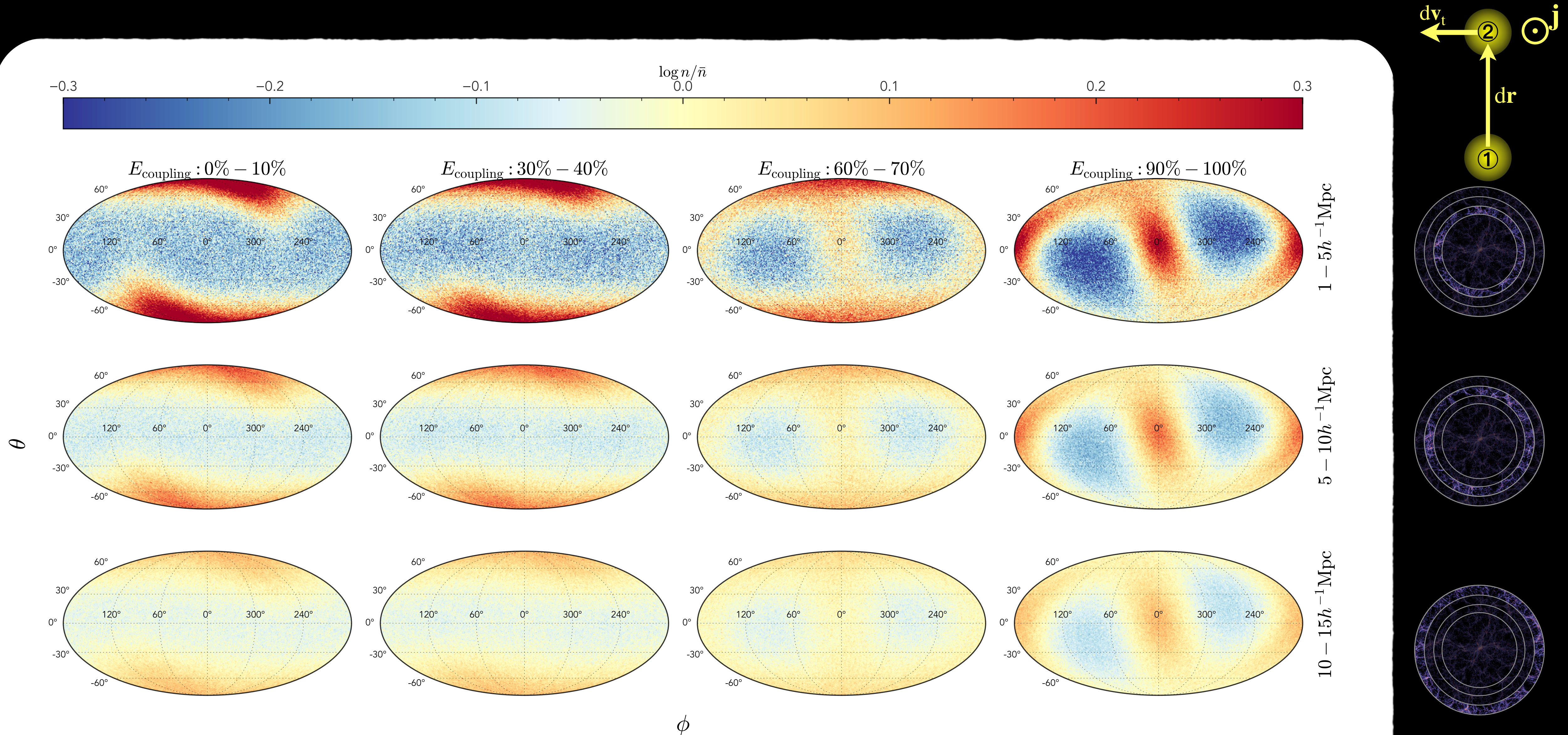
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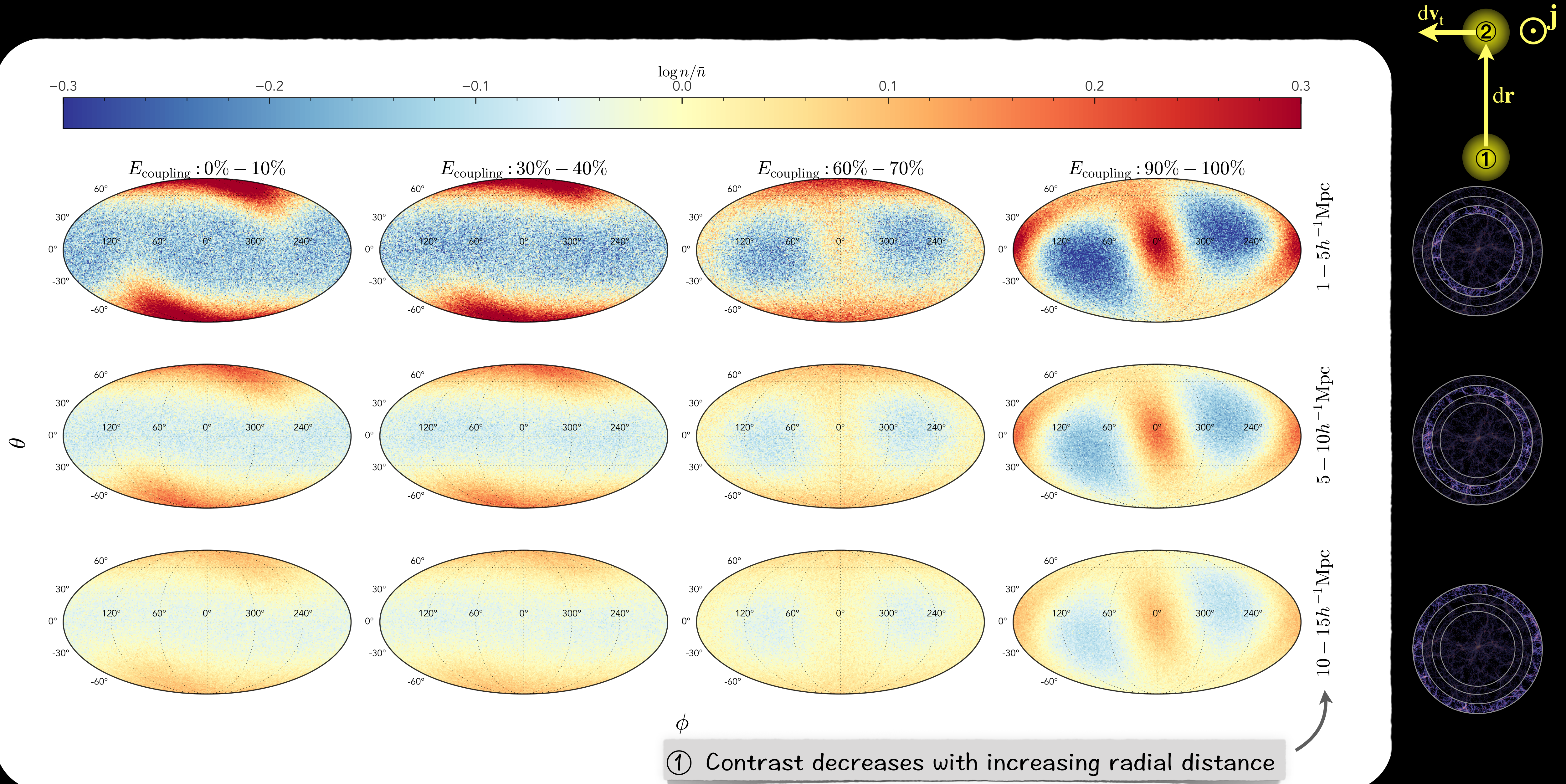
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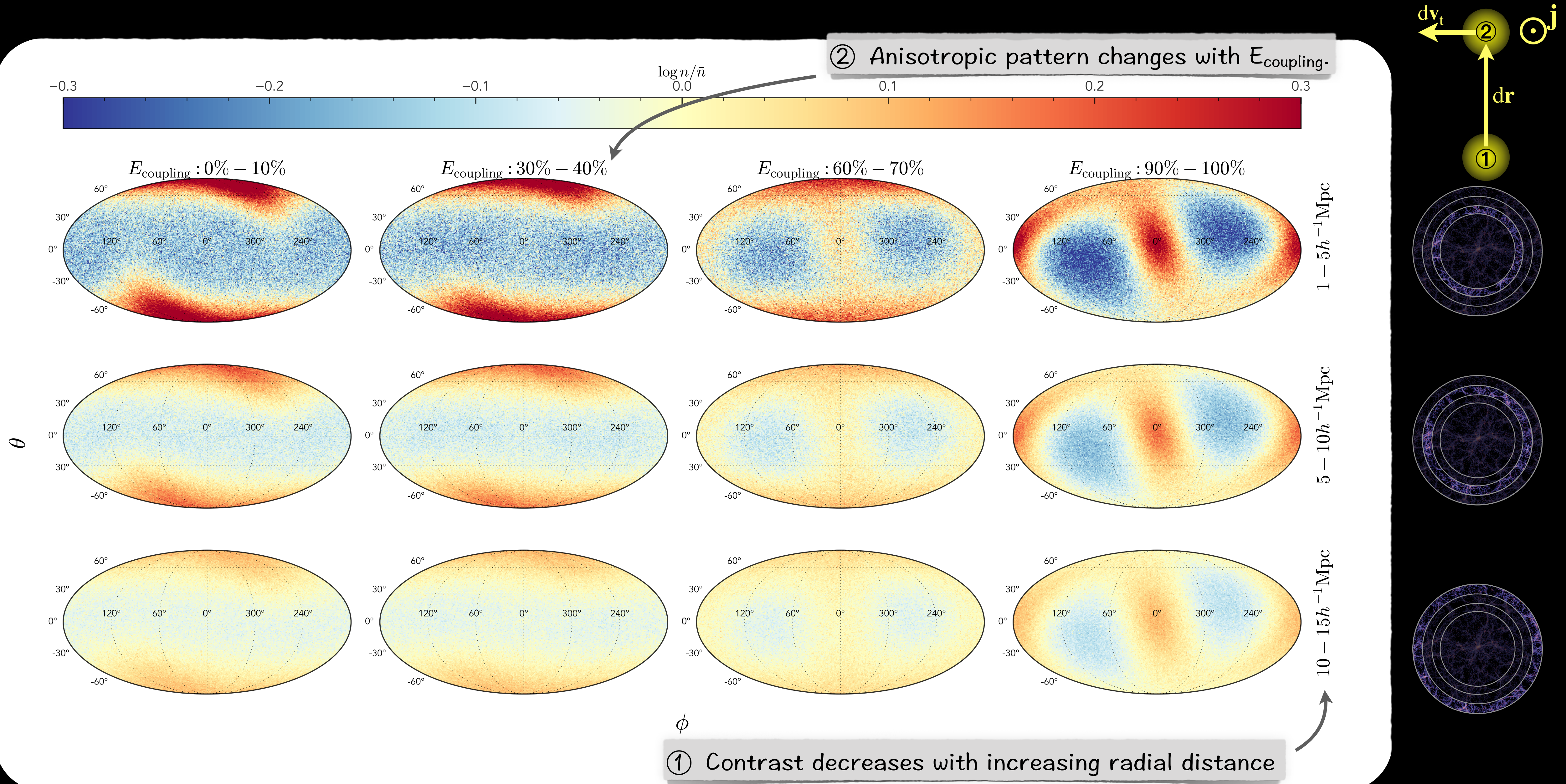
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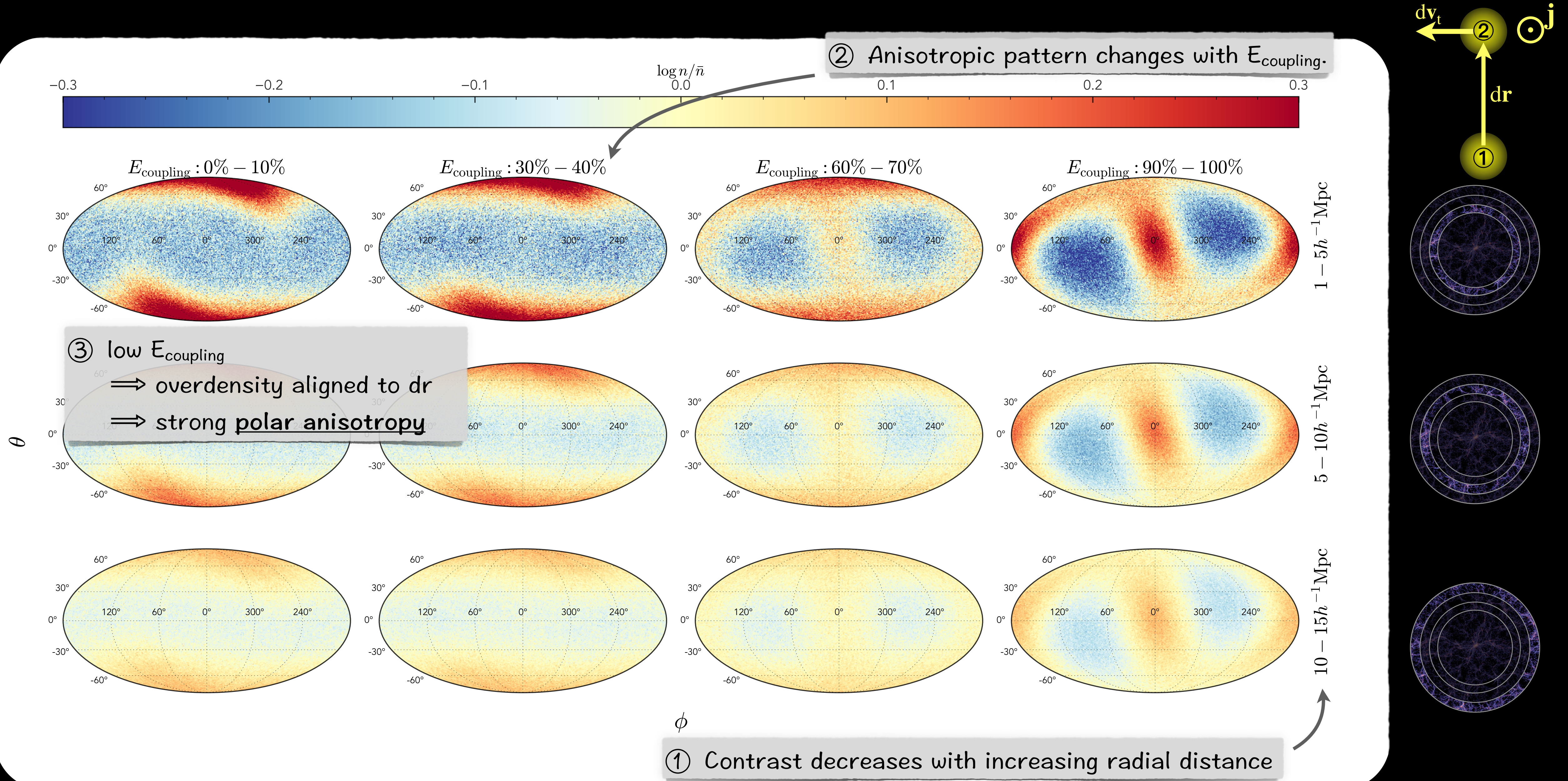
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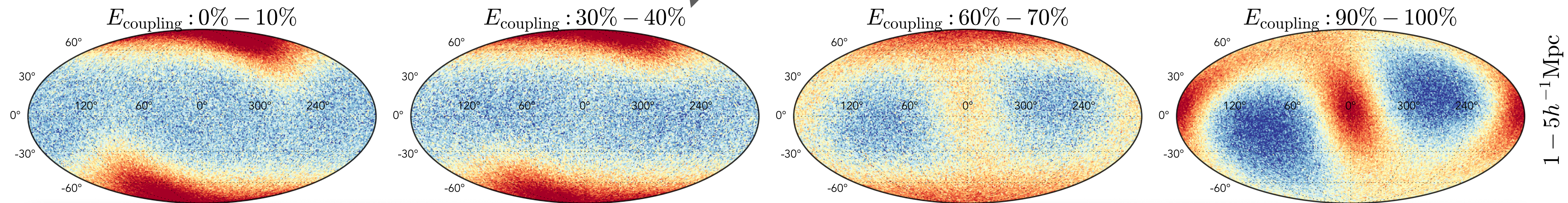
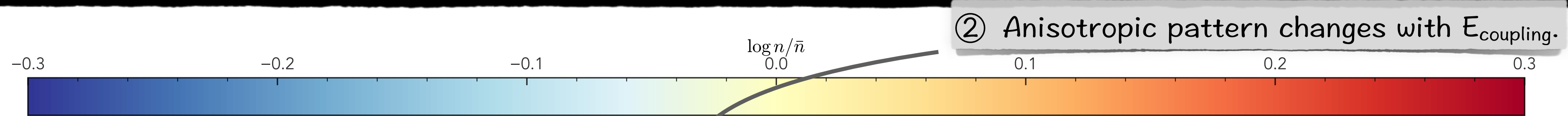
Correlations between Local Group Analogue and Cosmic Web



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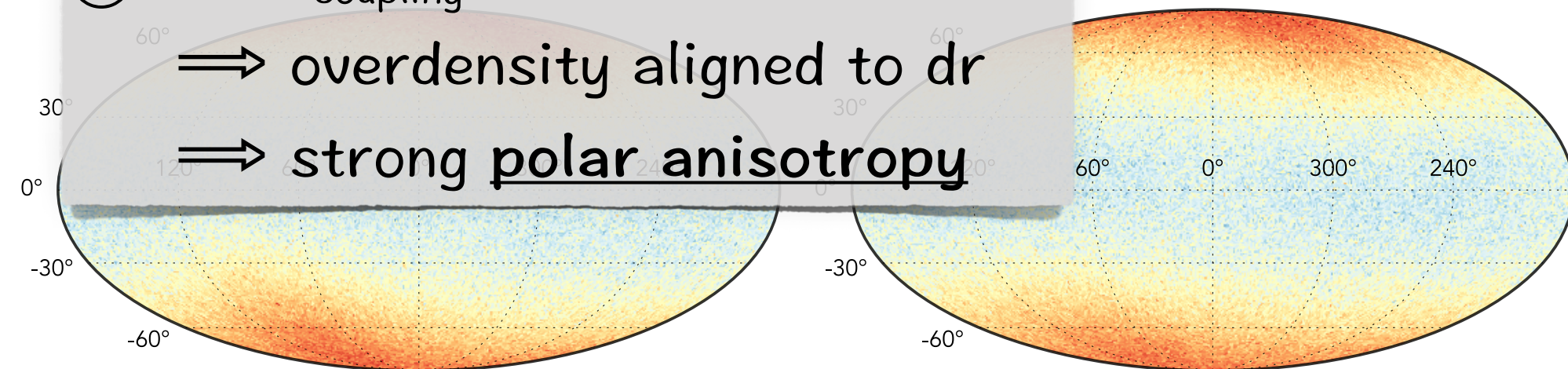


Correlations between Local Group Analogue and Cosmic Web



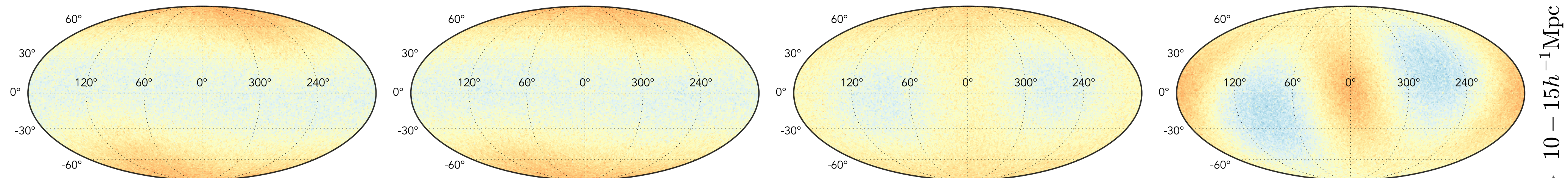
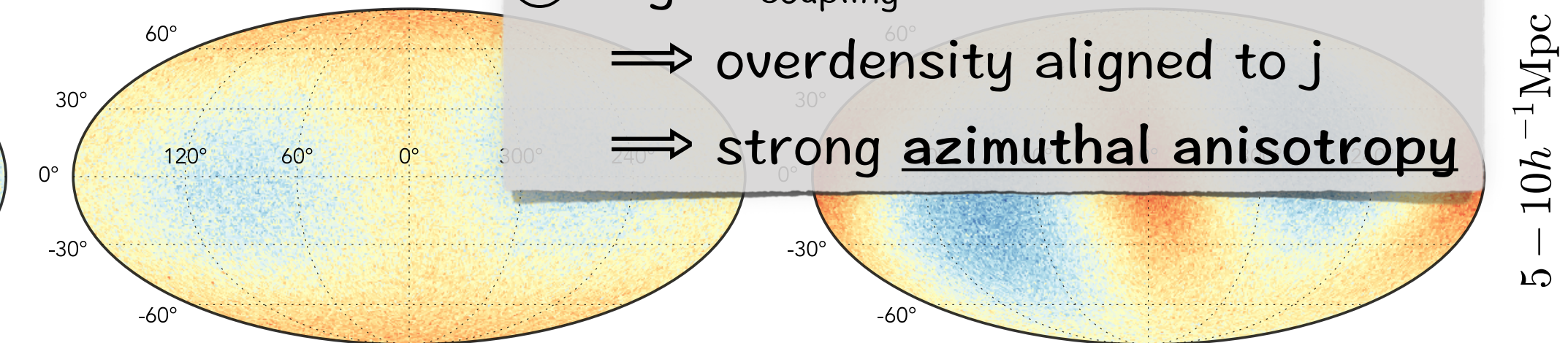
③ low E_{coupling}

$\Rightarrow$ overdensity aligned to dr
 $\Rightarrow$ strong polar anisotropy

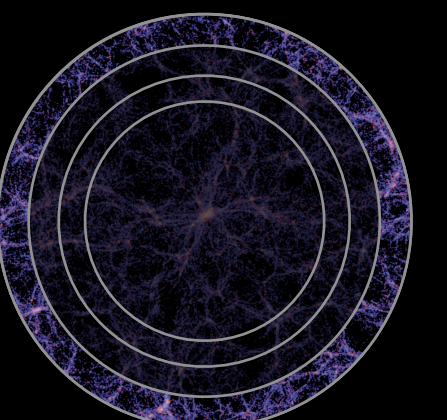
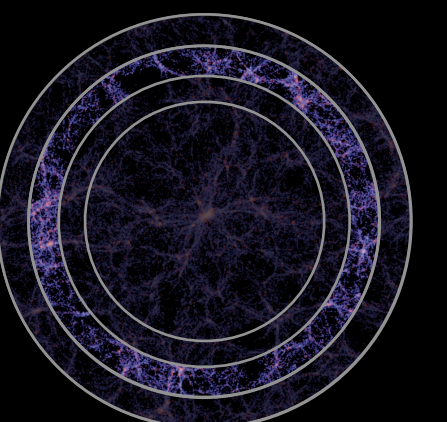
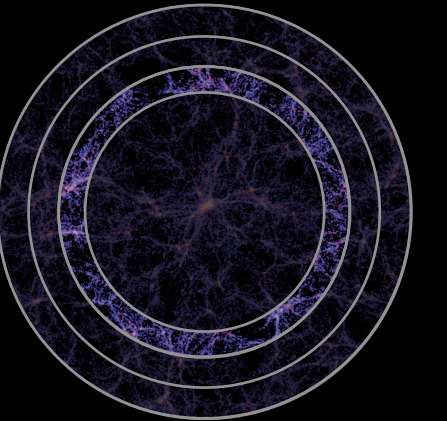
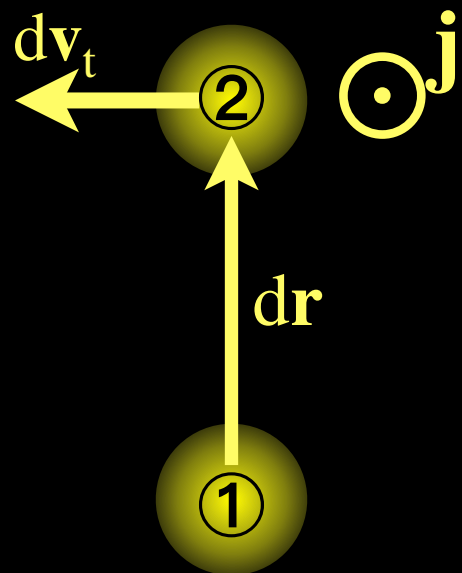


④ high E_{coupling}

$\Rightarrow$ overdensity aligned to j
 $\Rightarrow$ strong azimuthal anisotropy



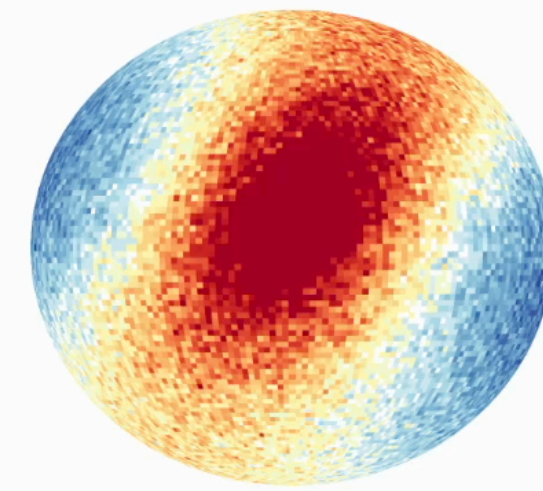
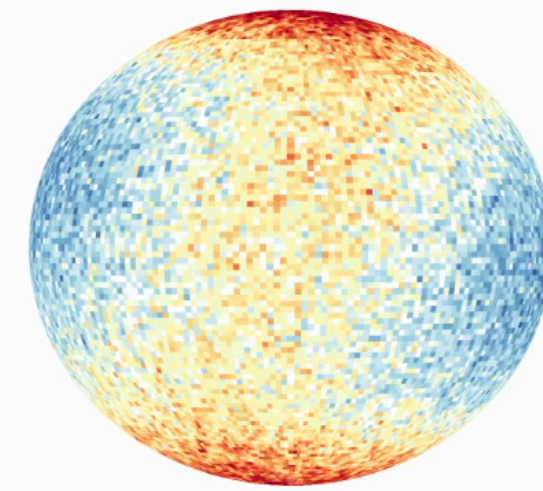
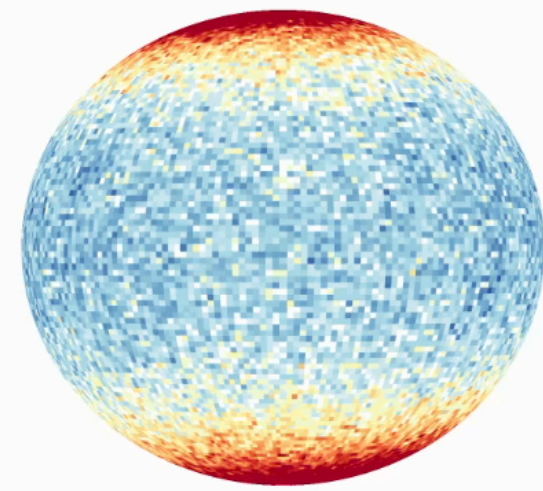
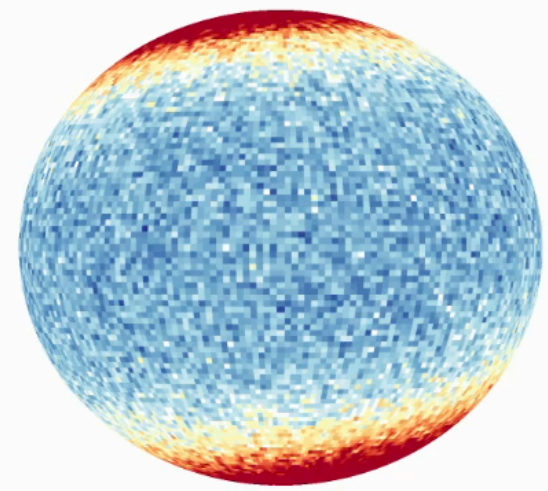
① Contrast decreases with increasing radial distance



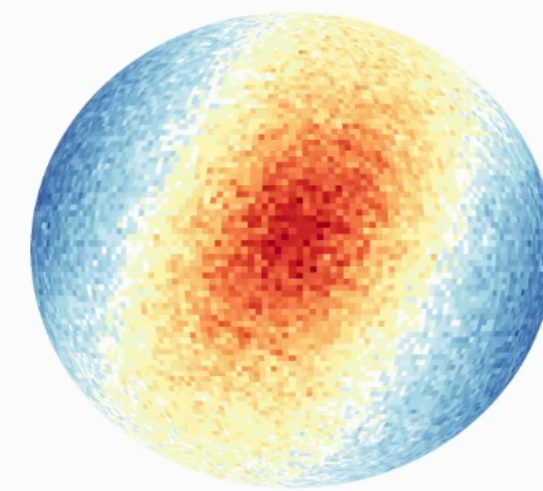
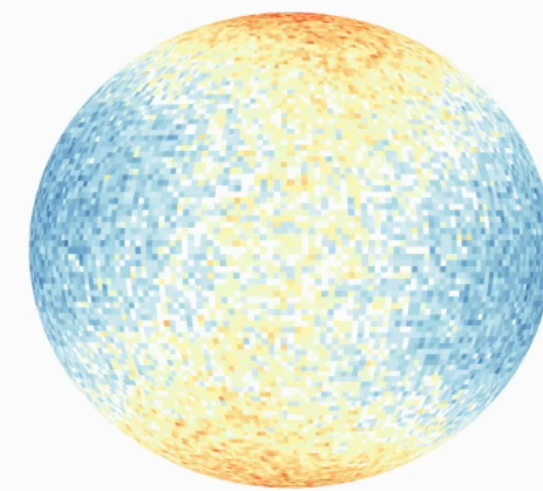
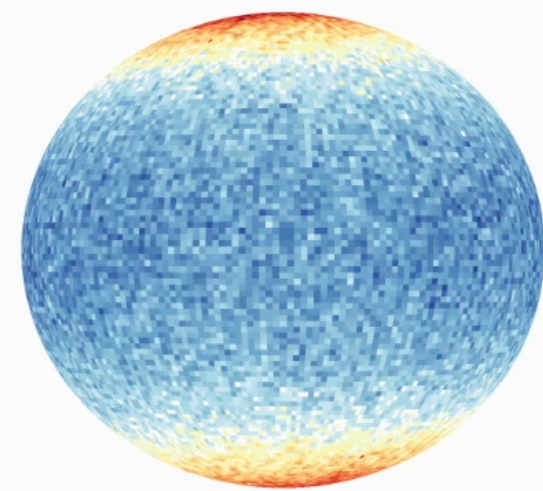
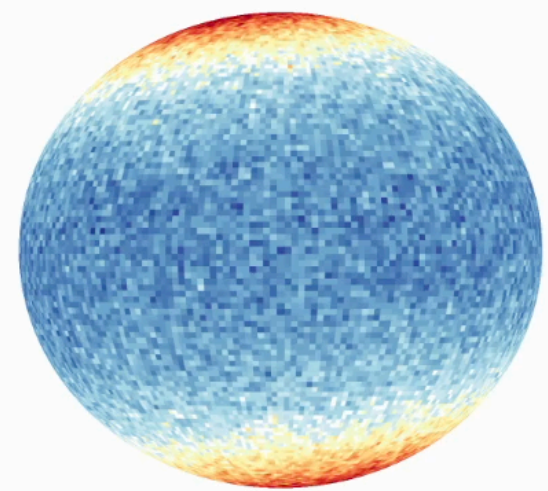
- Correlations between Local Group Analogue and Cosmic Web

Low E_{coupling}

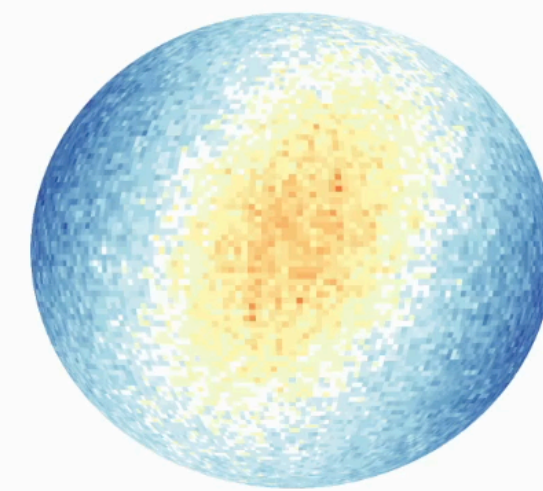
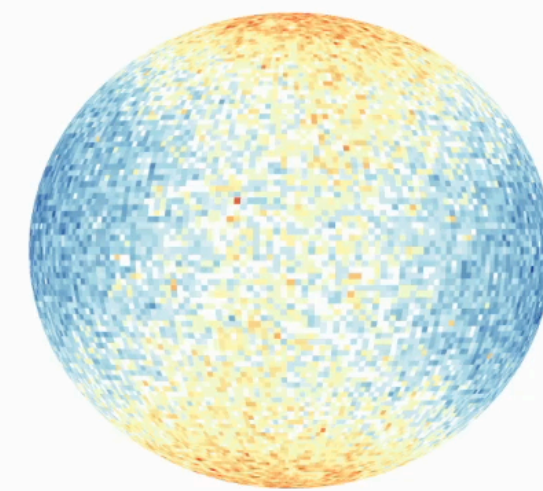
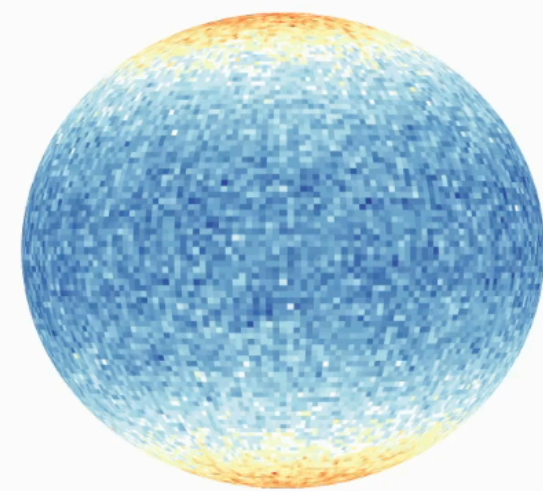
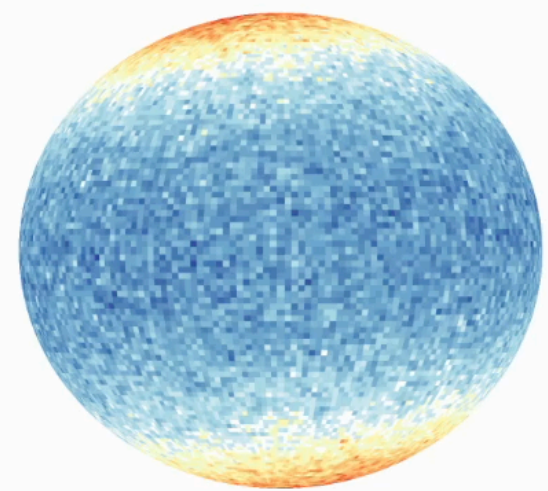
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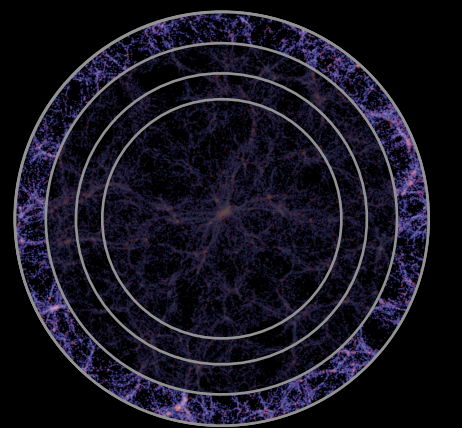
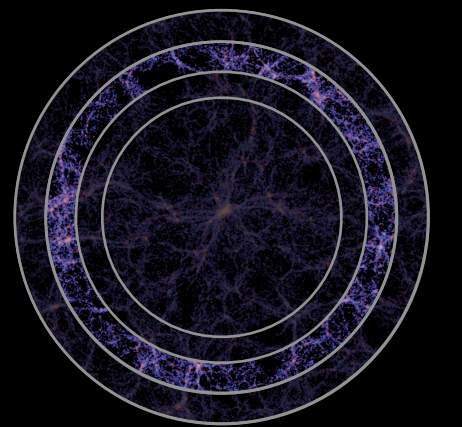
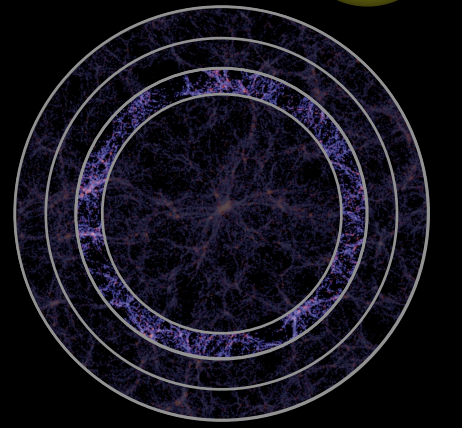
$1 - 5h^{-1}\text{Mpc}$



$5 - 10h^{-1}\text{Mpc}$



$10 - 15h^{-1}\text{Mpc}$



②

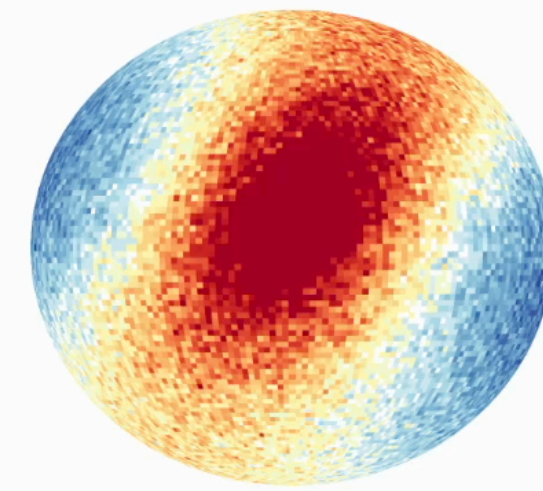
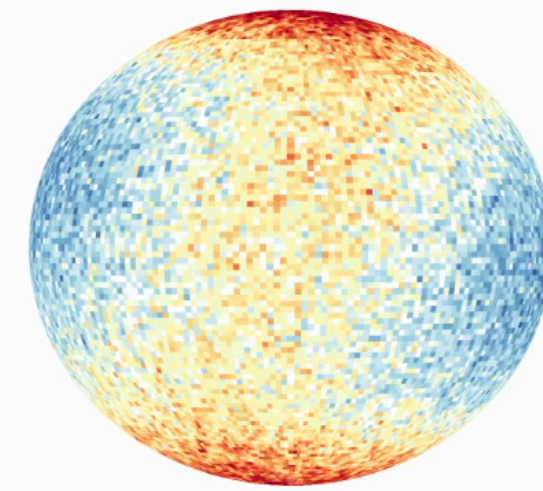
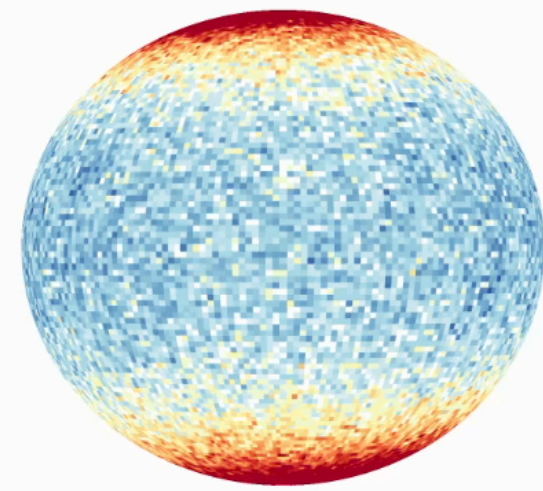
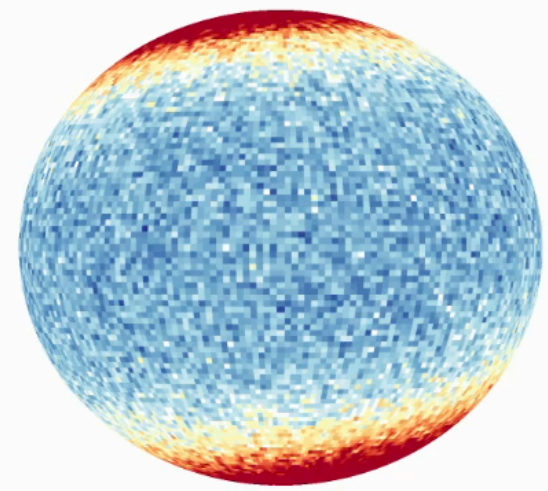
dr

①

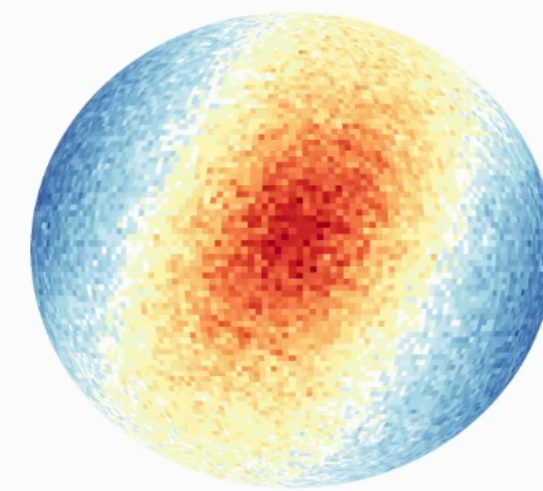
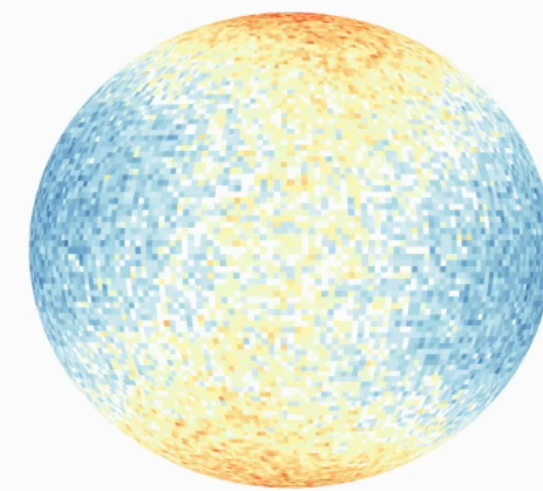
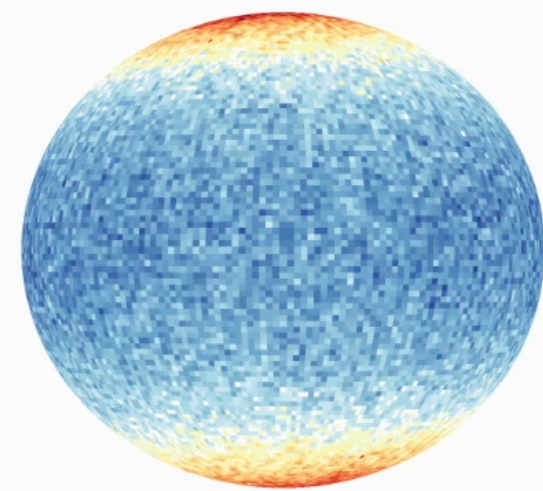
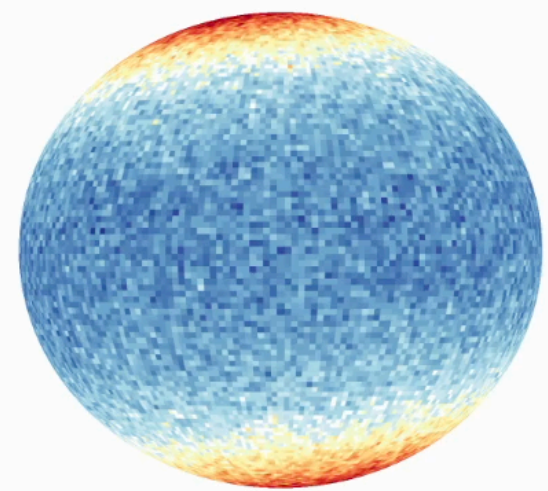
- Correlations between Local Group Analogue and Cosmic Web

Low E_{coupling}

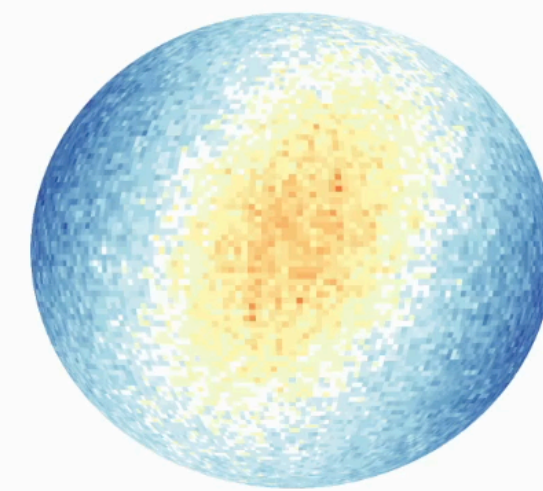
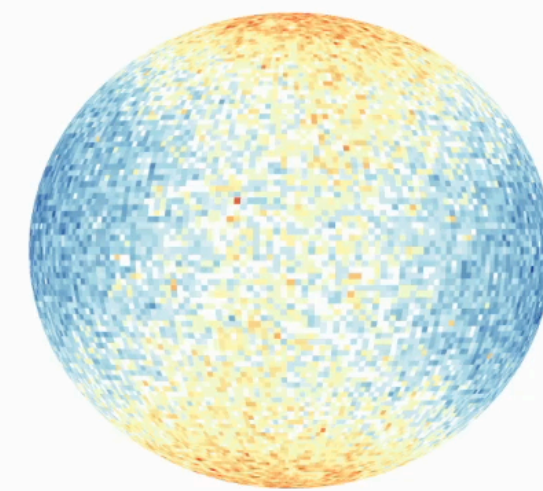
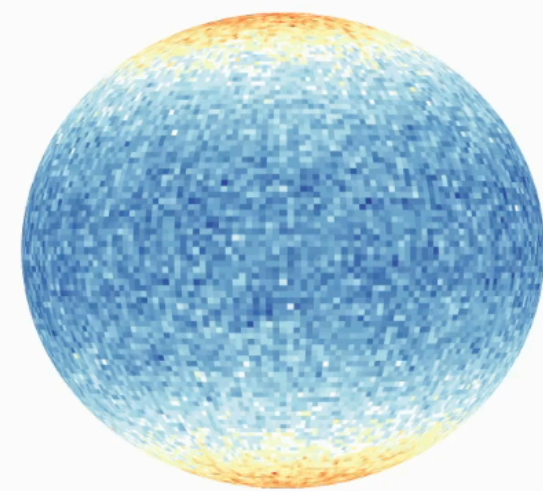
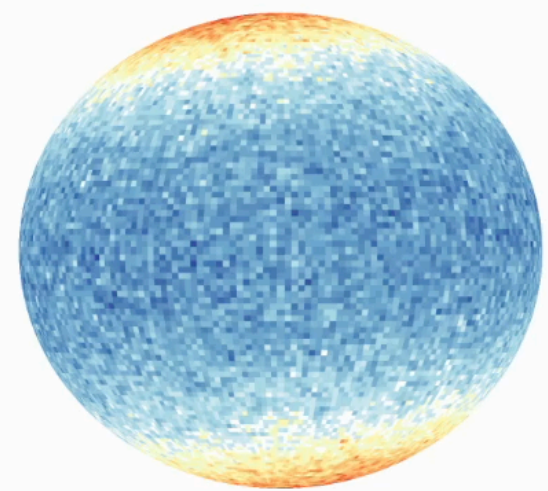
High E_{coupling}



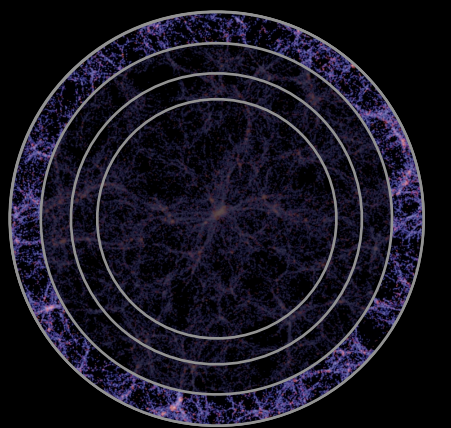
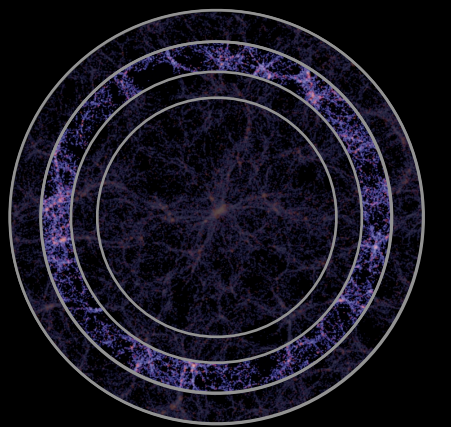
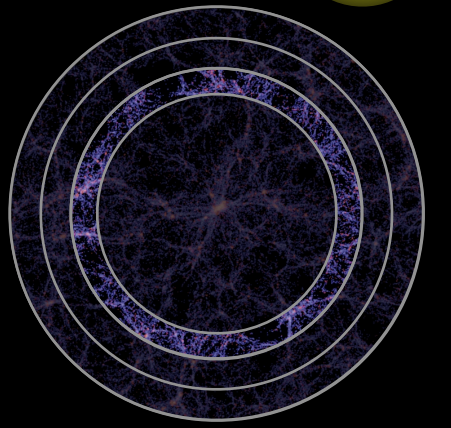
1 – 5 h^{-1} Mpc



5 – 10 h^{-1} Mpc



10 – 15 h^{-1} Mpc



②

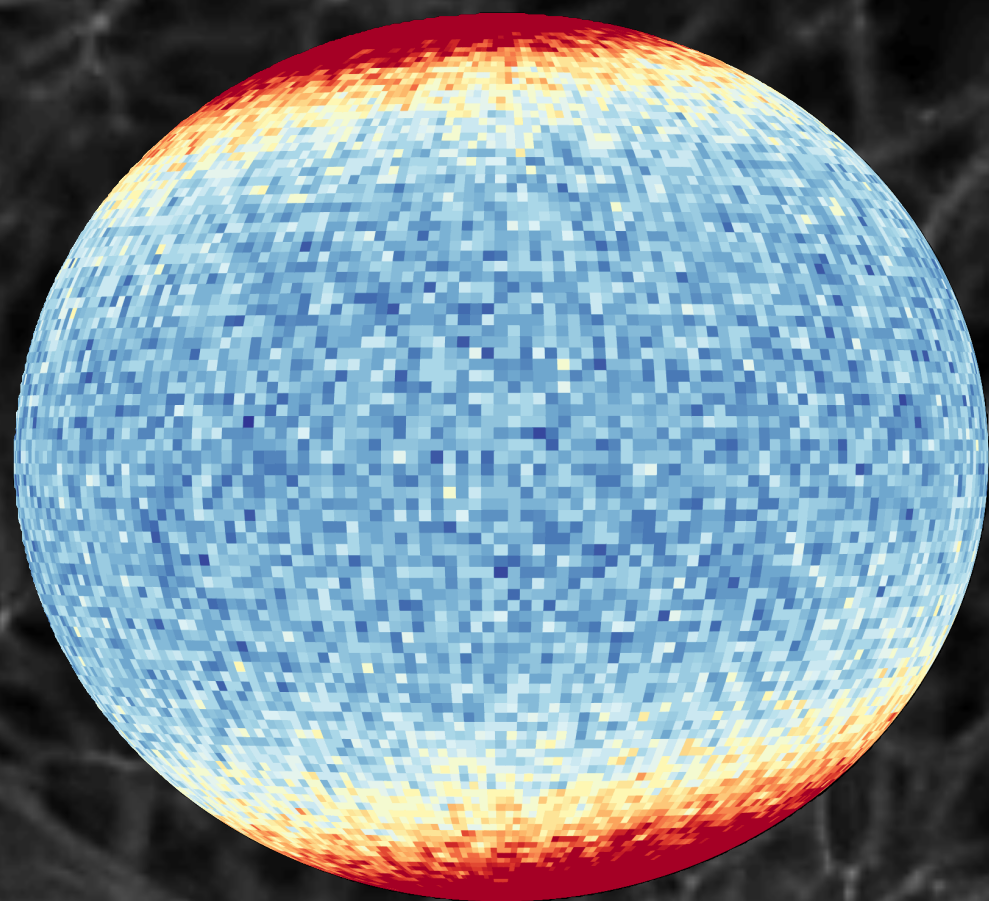
dr

①

Take home: We found Local Group Analogues with ...

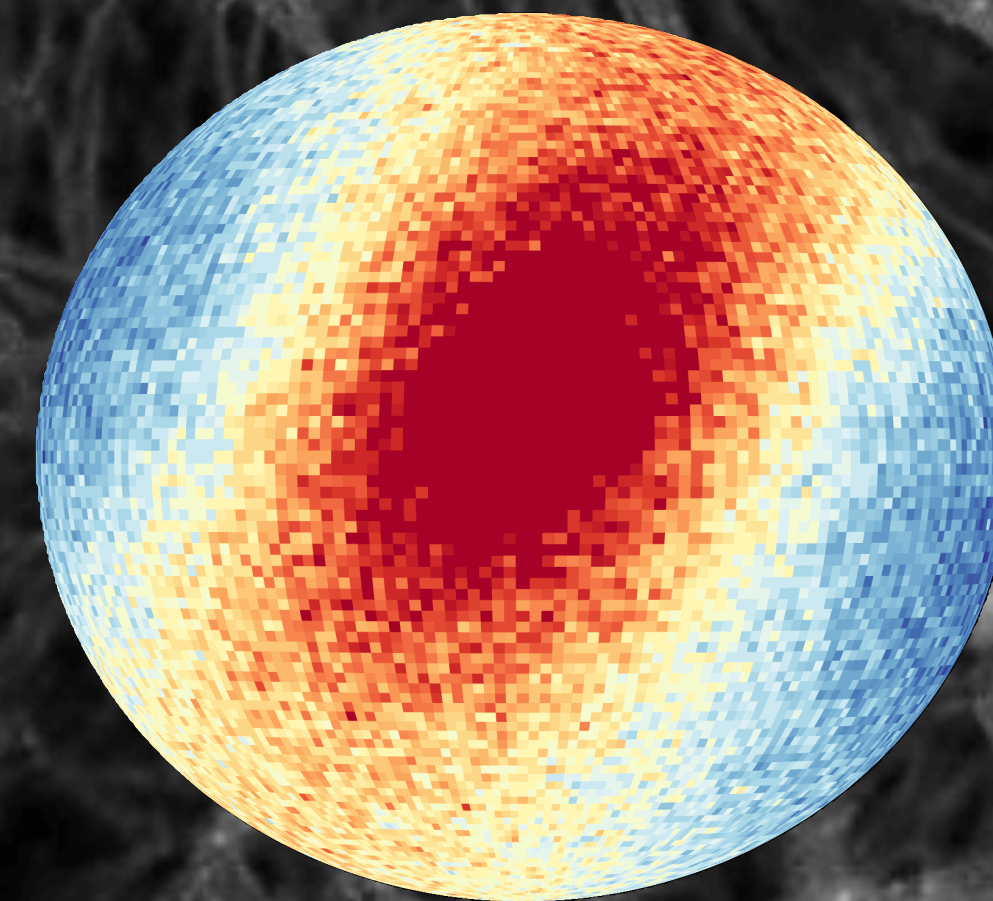
Low E_{coupling}

- live in low-density regions
- displacement vec aligned to LSS
- strong polar anisotropy
- weak azimuthal anisotropy



High E_{coupling}

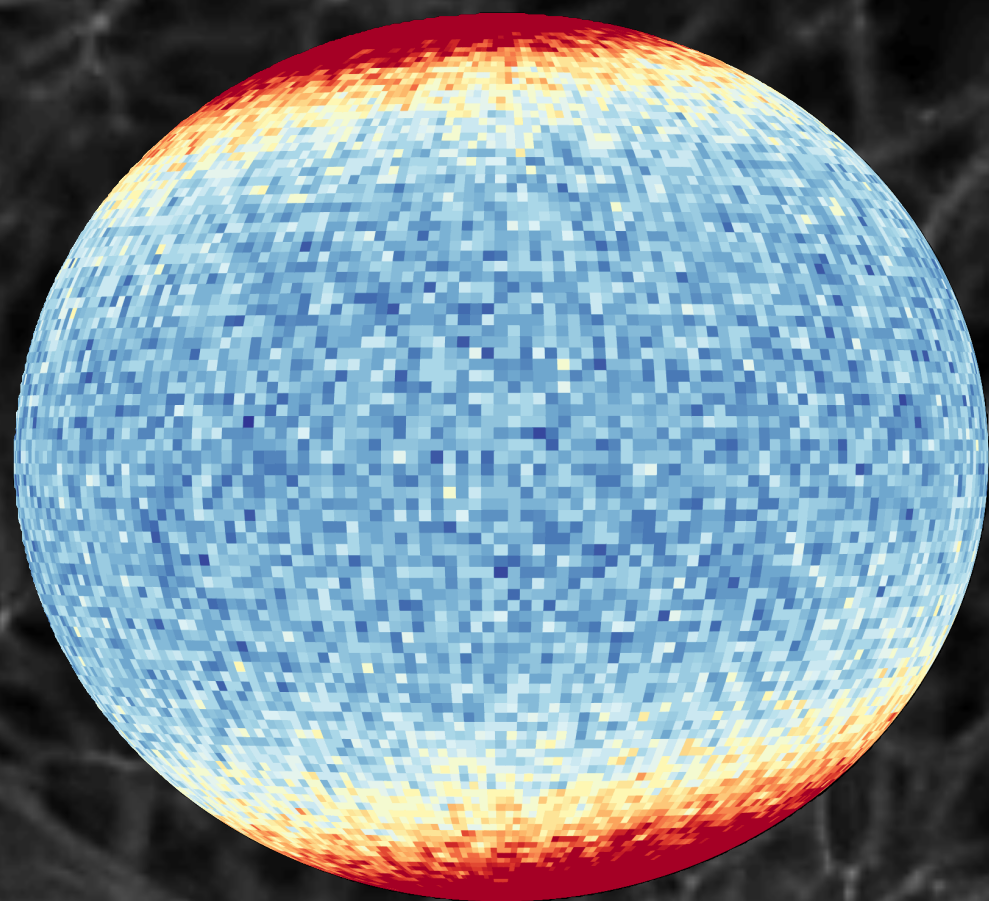
- live in high-density regions
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Take home: We found Local Group Analogues with ...

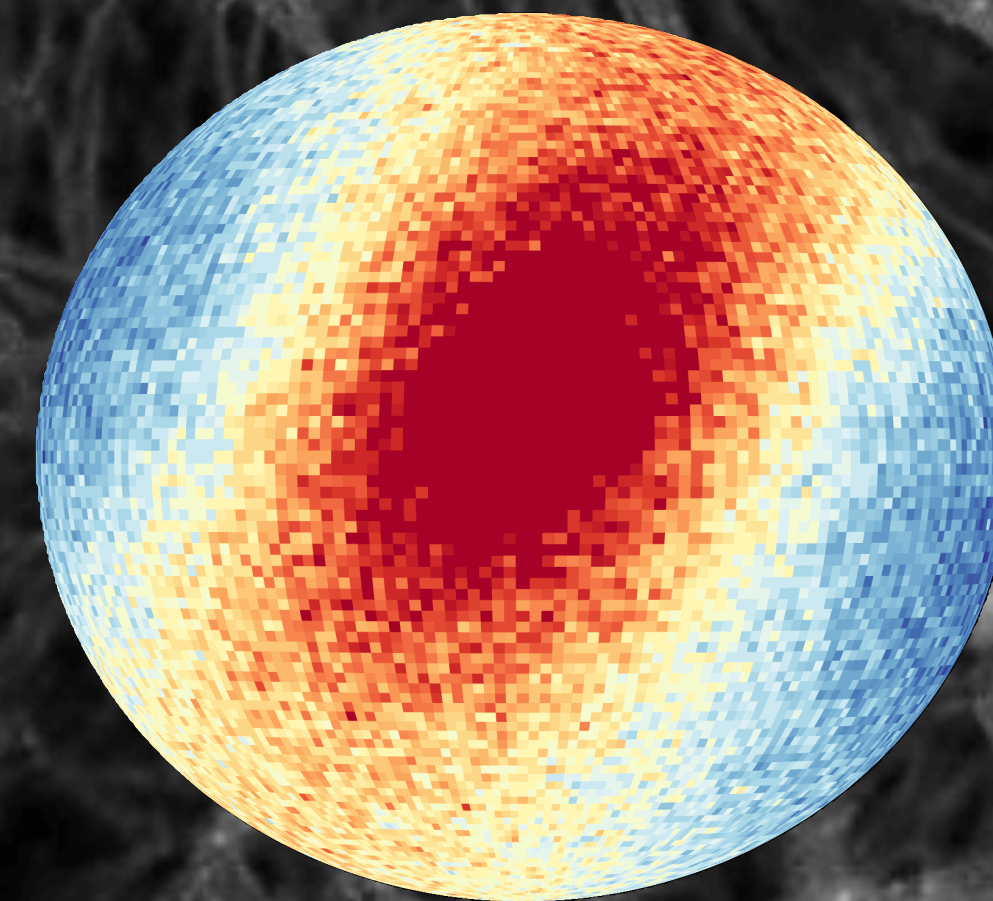
Low E_{coupling}

- live in low-density regions
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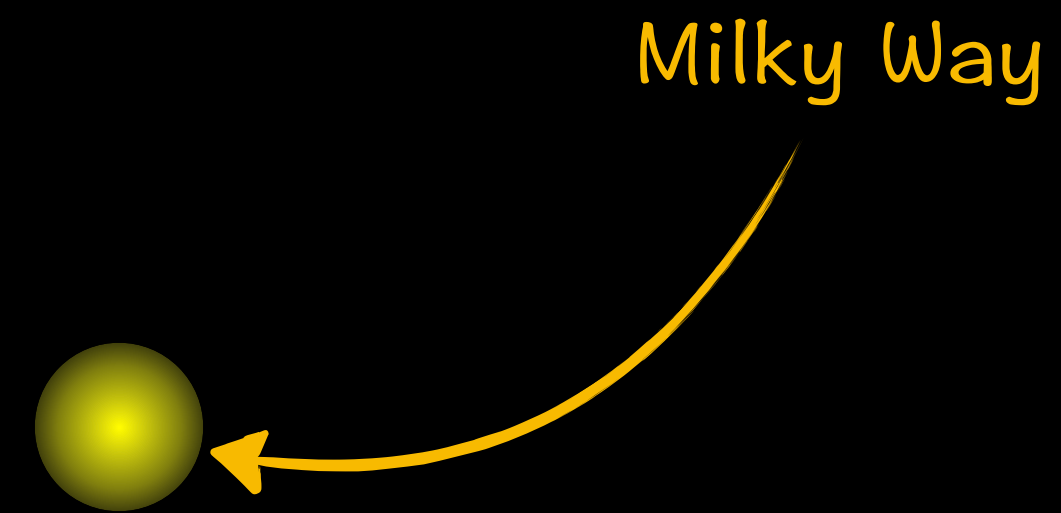


High E_{coupling}

- live in high-density regions
- orbital spin aligned to LSS
- weak polar anisotropy
- strong azimuthal anisotropy



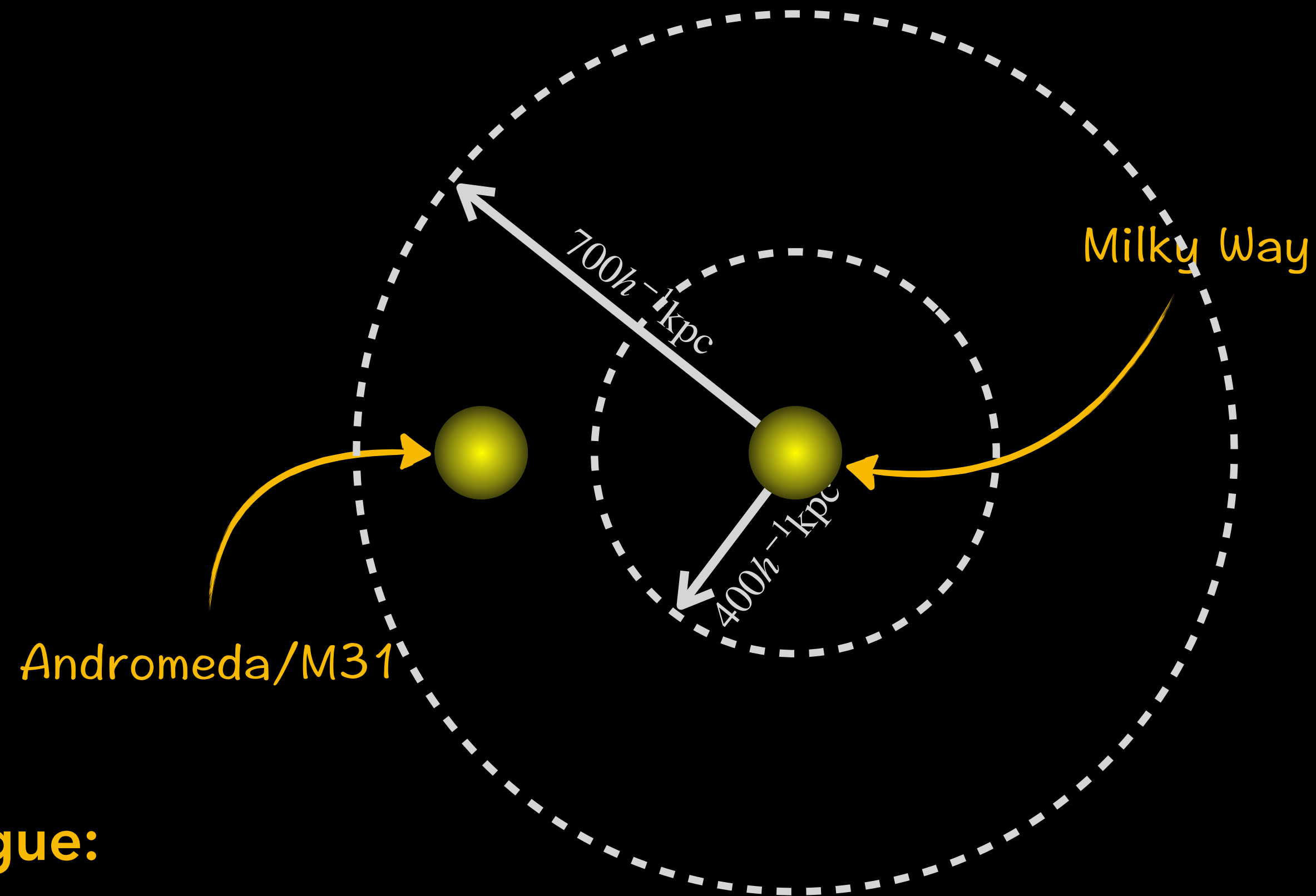
- **Local Group Analogue**



Local Group Analogue:

- One MW with $M_{\text{halo}} \sim 10^{12 \pm 0.3} M_{\odot}/h$

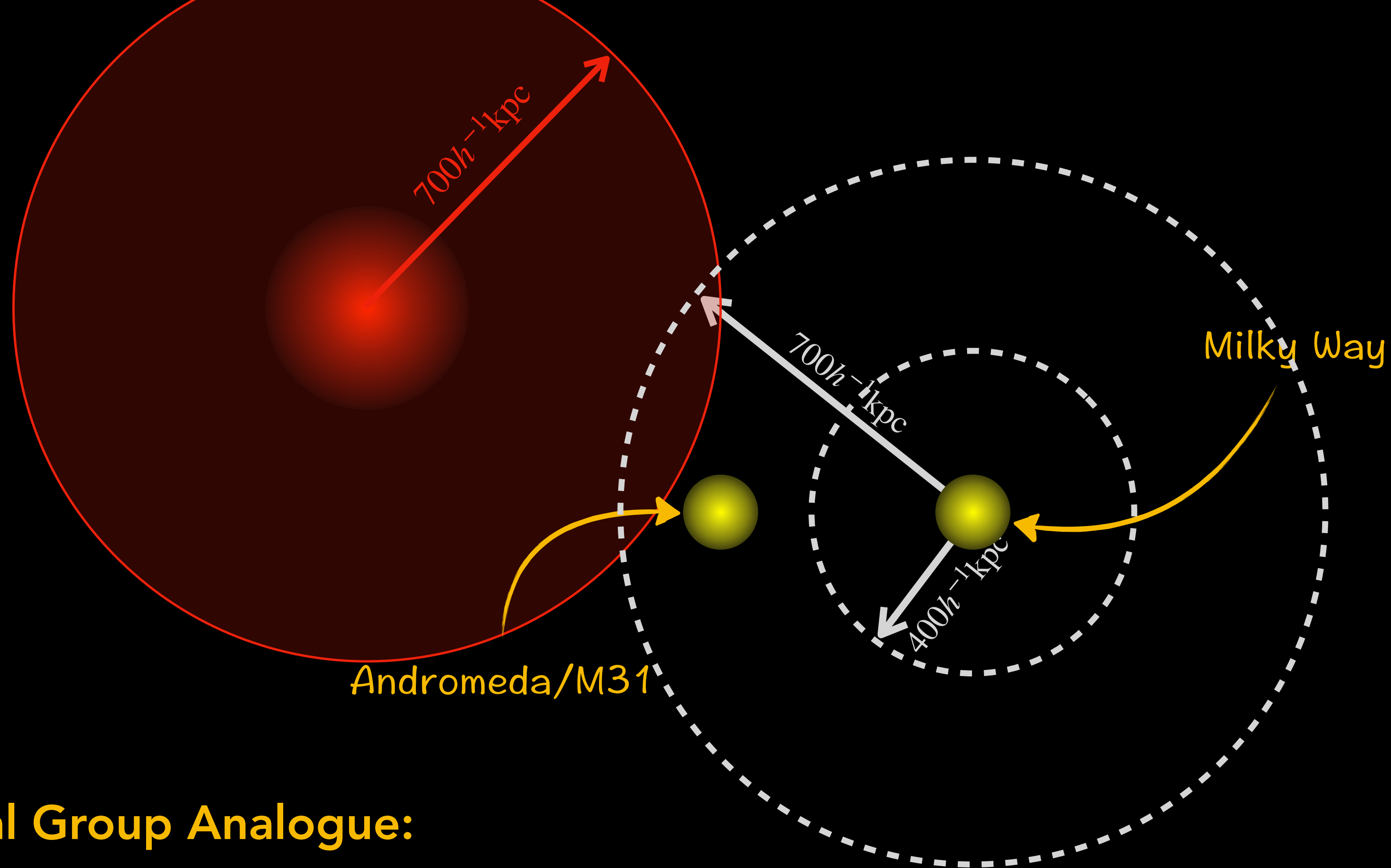
- **Local Group Analogue**



Local Group Analogue:

- One MW with $M_{\text{halo}} \sim 10^{12 \pm 0.3} M_{\odot}/h$
- One MW-like-companion within $[400, 700] \text{ kpc}/h$

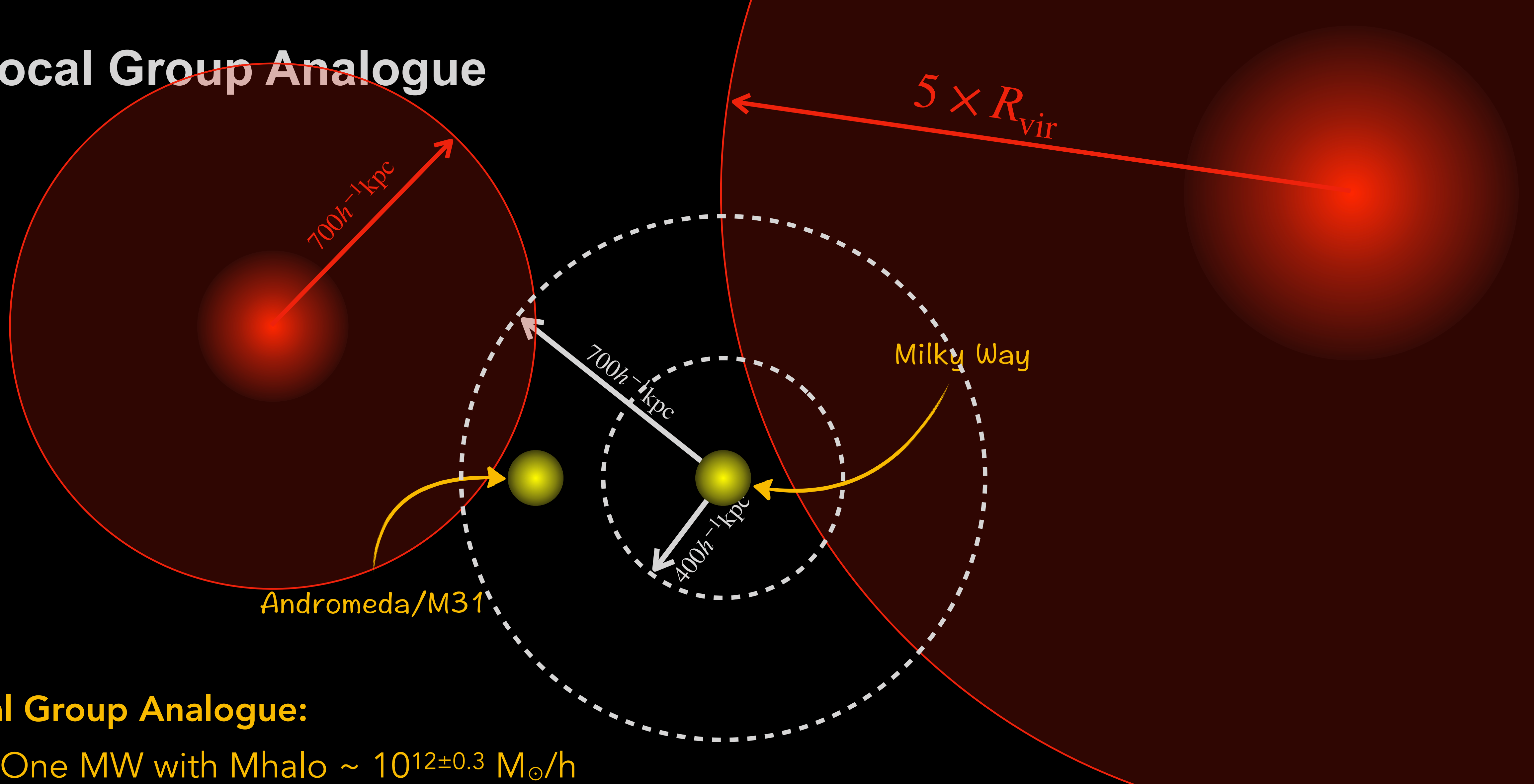
- **Local Group Analogue**



Local Group Analogue:

- One MW with $M_{\text{halo}} \sim 10^{12 \pm 0.3} M_{\odot}/h$
- One MW-like-companion within $[400, 700] \text{ kpc}/h$
- No other massive halos within $700 \text{ kpc}/h$

• Local Group Analogue



Local Group Analogue:

- One MW with $M_{\text{halo}} \sim 10^{12 \pm 0.3} M_{\odot}/h$
- One MW-like-companion within $[400, 700] \text{ kpc}/h$
- No other massive halos within $700 \text{ kpc}/h$
- $> 5R_{\text{vir}}$ away from other massive halos

ABACUSSUMMIT simulation:

- contains 25 $(2\text{Gpc}/h)^3$ boxes
- particle mass of $2 \times 10^9 M_{\text{sun}}/h$

- **Correlations between Local Group Analogue and Cosmic Web**

- coupling energy:

$$E_{\text{coupling}} = \frac{1}{2} \frac{m_1 m_2}{m_1 + m_2} V^2 - \frac{G m_1 m_2}{\|\mathbf{dr}\|}$$

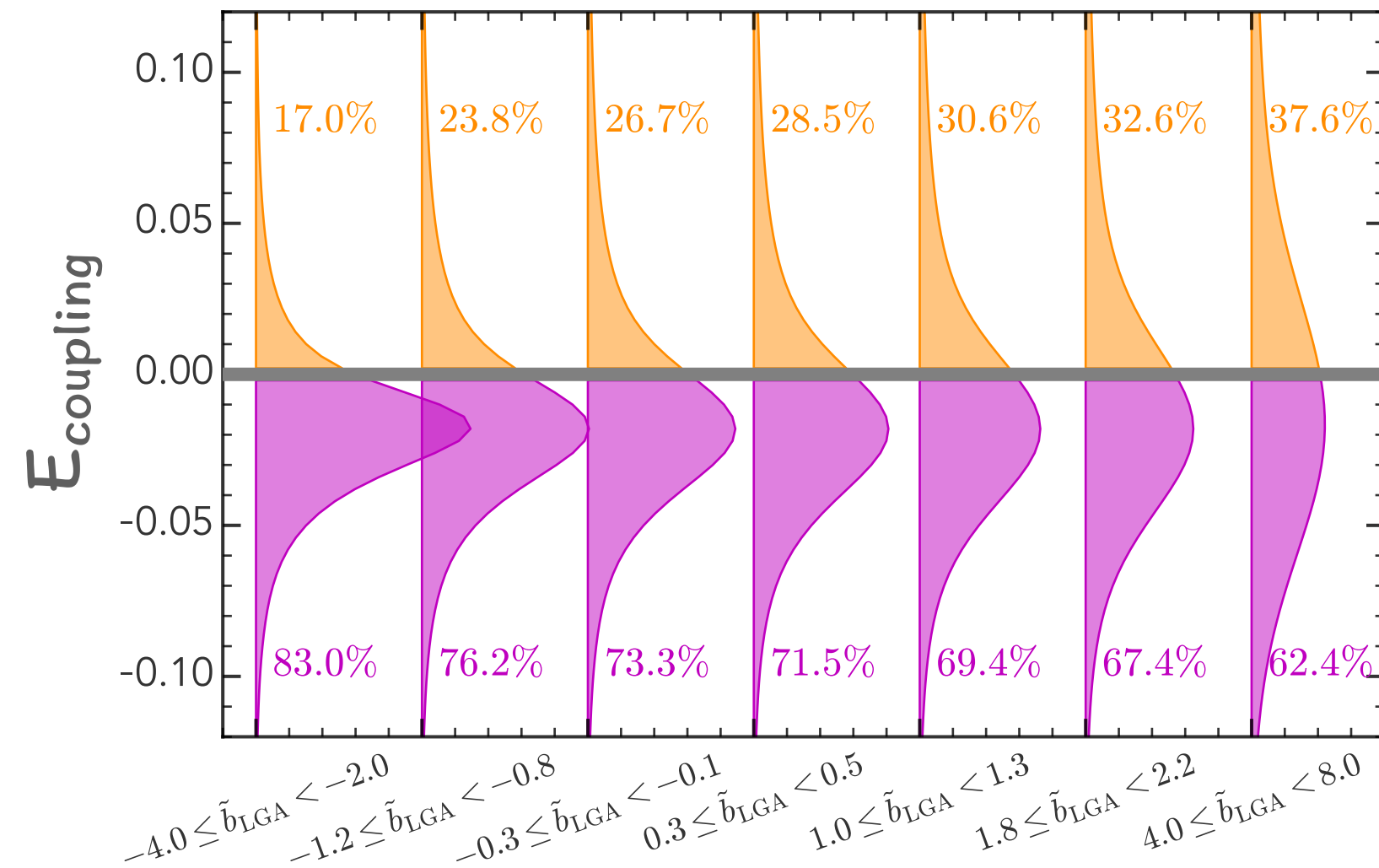
- large-scale bias:

$$\tilde{b}_{\text{LGA}}(r) = \delta(r) / \xi_{\text{mm}}(r)$$

traced by halos above $10^{11.6} M_{\odot}/h$

• Correlations between Local Group Analogue and Cosmic Web

low-density env. $\Longrightarrow$ high-density env.



- coupling energy:

$$E_{\text{coupling}} = \frac{1}{2} \frac{m_1 m_2}{m_1 + m_2} v^2 - \frac{G m_1 m_2}{\|\mathbf{dr}\|}$$

- large-scale bias:

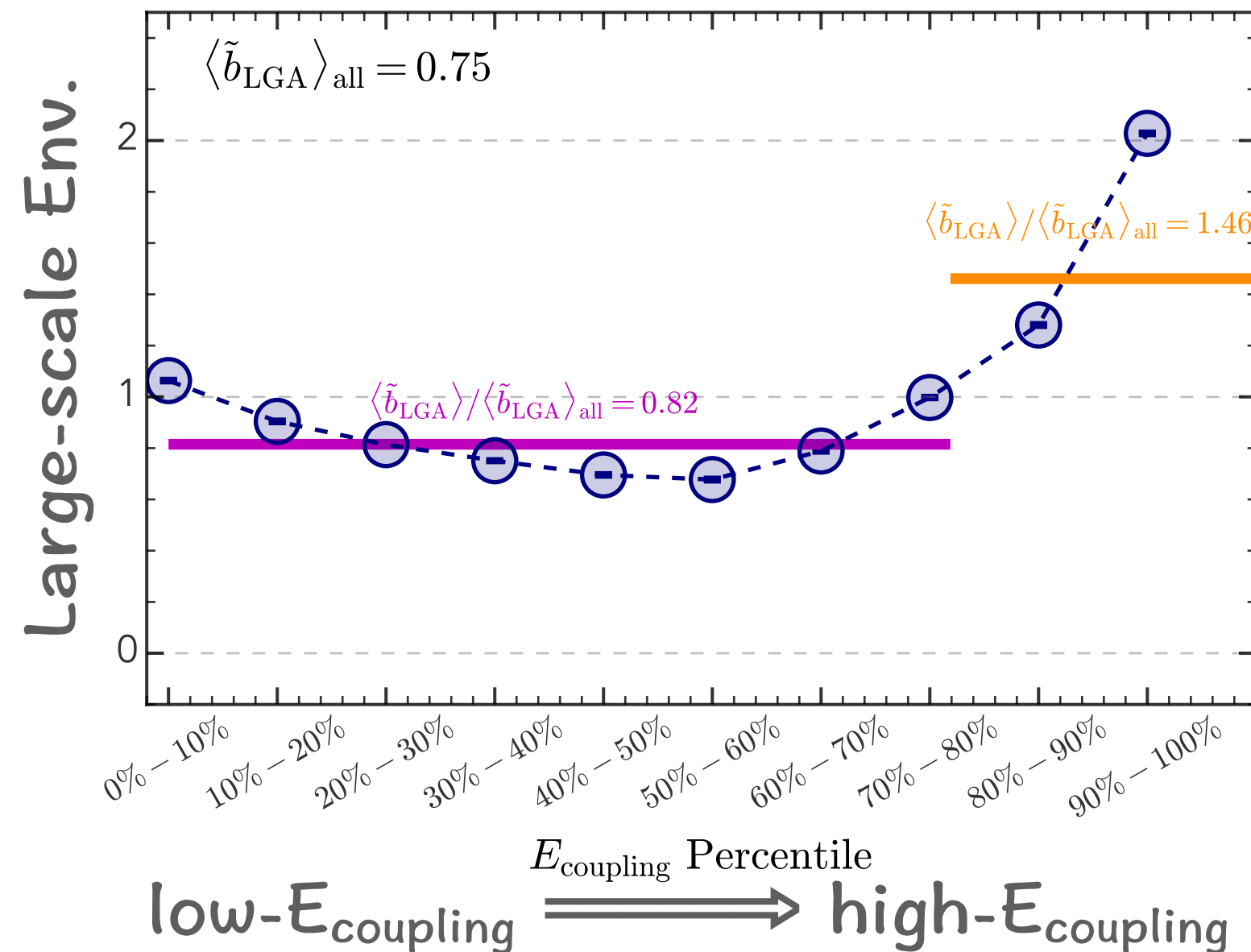
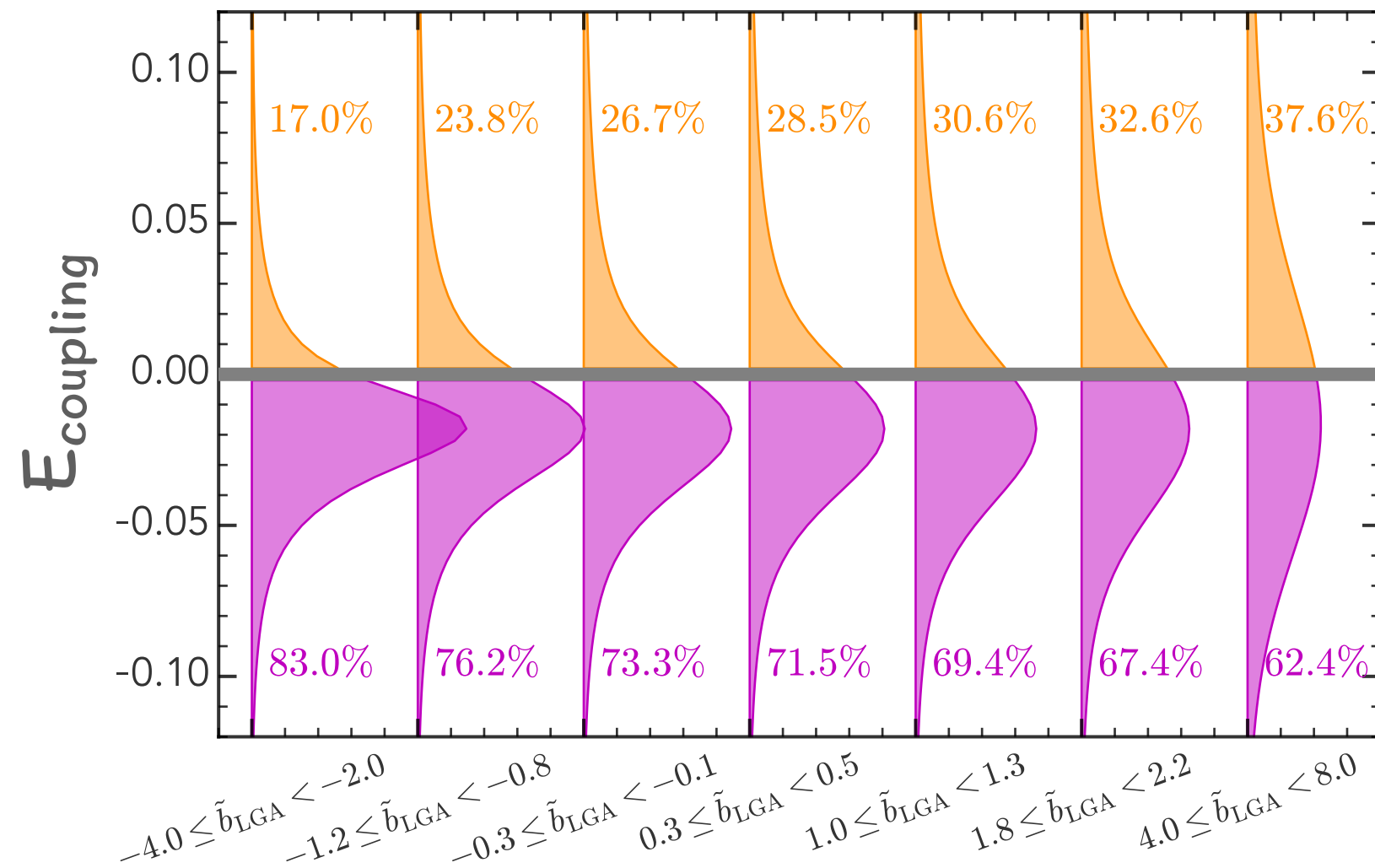
$$\tilde{b}_{\text{LGA}}(r) = \delta(r) / \xi_{\text{mm}}(r)$$

traced by halos above $10^{11.6} M_{\odot}/h$

- LGAs in higher-density regions have higher E_{coupling} .

• Correlations between Local Group Analogue and Cosmic Web

low-density env. $\Longrightarrow$ high-density env.



- coupling energy:

$$E_{\text{coupling}} = \frac{1}{2} \frac{m_1 m_2}{m_1 + m_2} v^2 - \frac{G m_1 m_2}{\|\mathbf{dr}\|}$$

- large-scale bias:

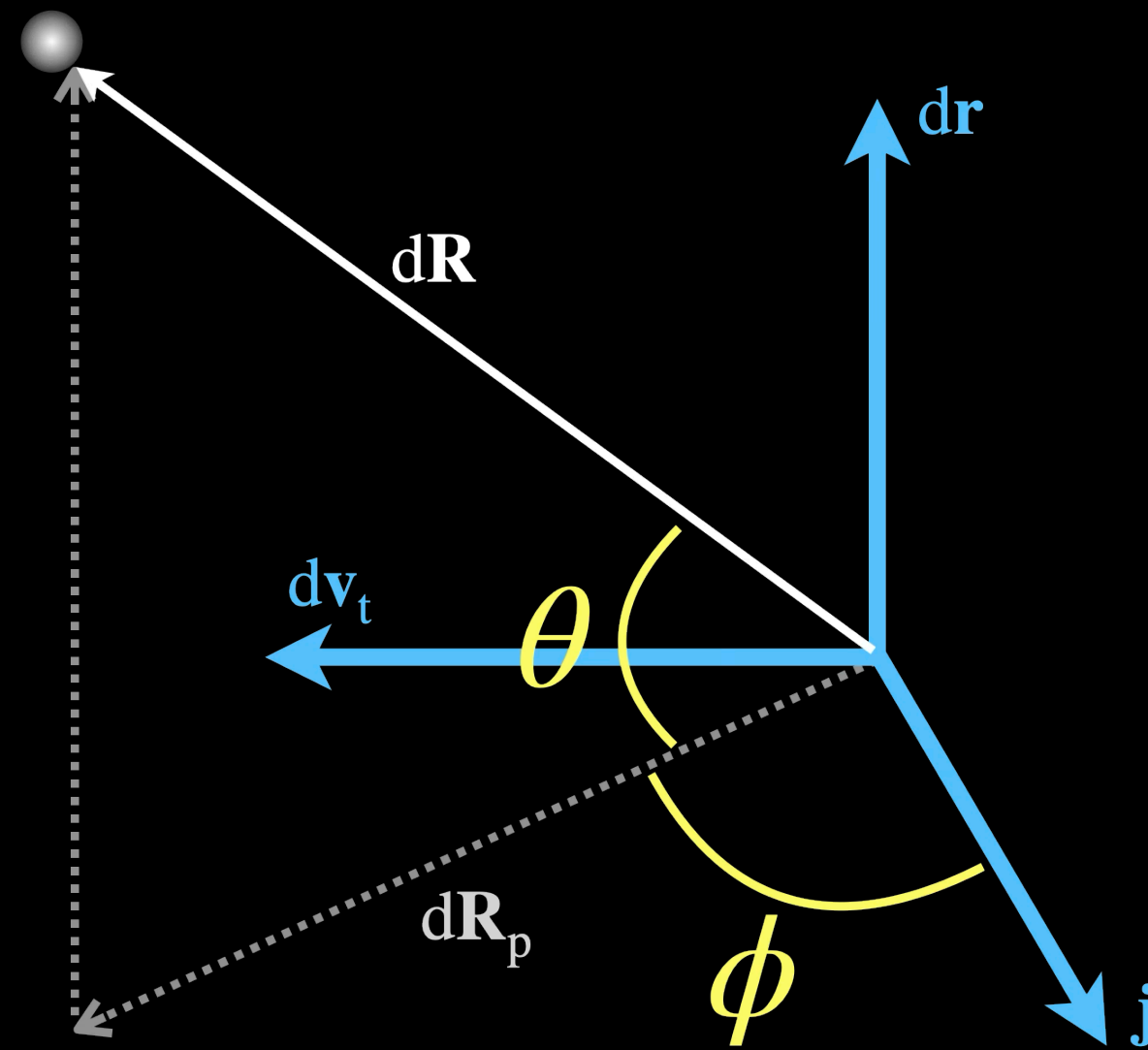
$$\tilde{b}_{\text{LGA}}(r) = \delta(r) / \xi_{\text{mm}}(r)$$

traced by halos above $10^{11.6} M_{\odot}/h$

- LGAs in higher-density regions have higher E_{coupling} .
- Higher- E_{coupling} LGAs live in denser regions and the large-scale bias differs by $\sim 80\%$.
 $\Longrightarrow$ strong "secondary halo bias" effect

Correlations between Local Group Analogue and Cosmic Web

LSS tracer halo



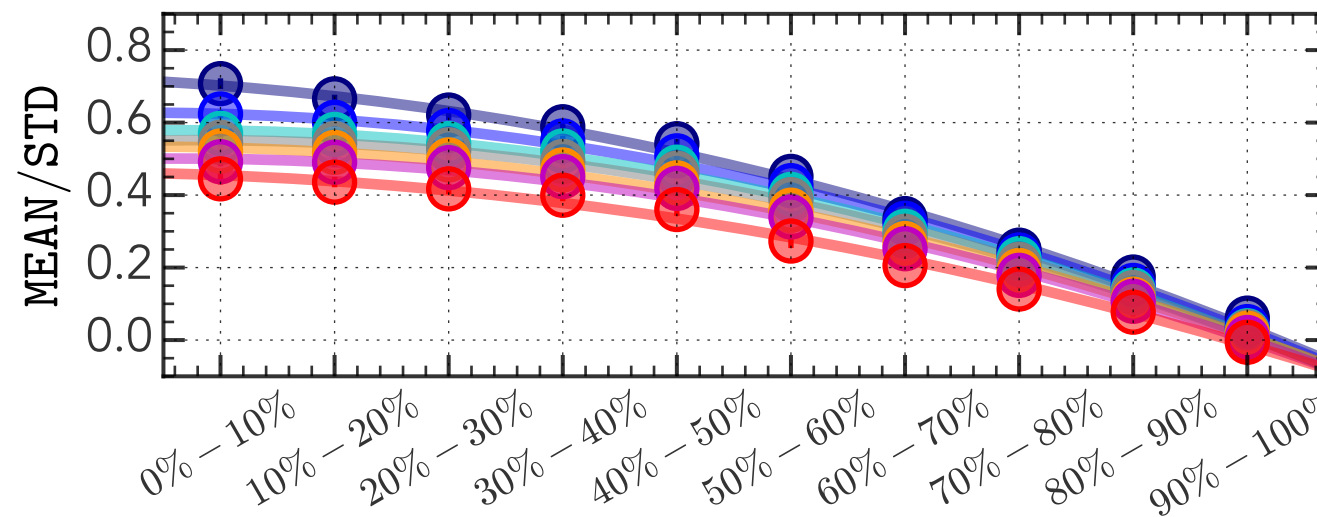
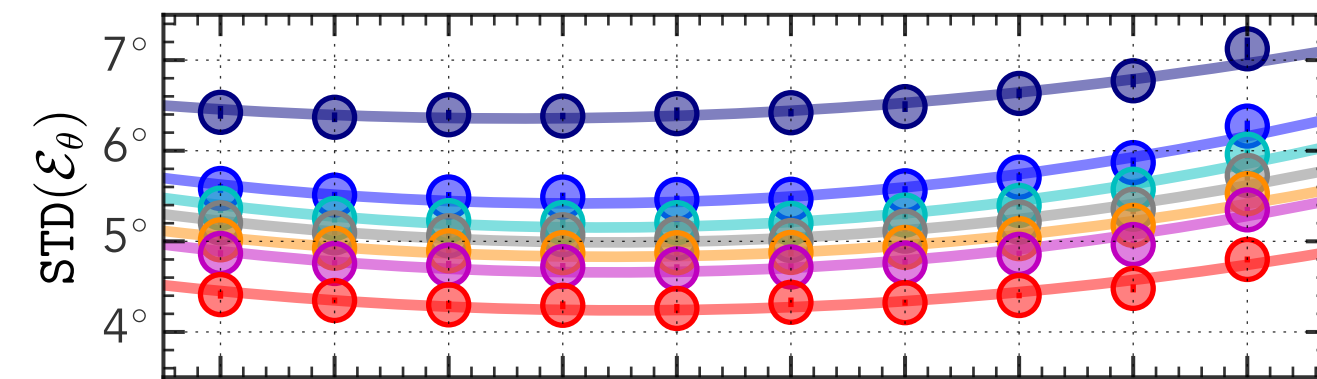
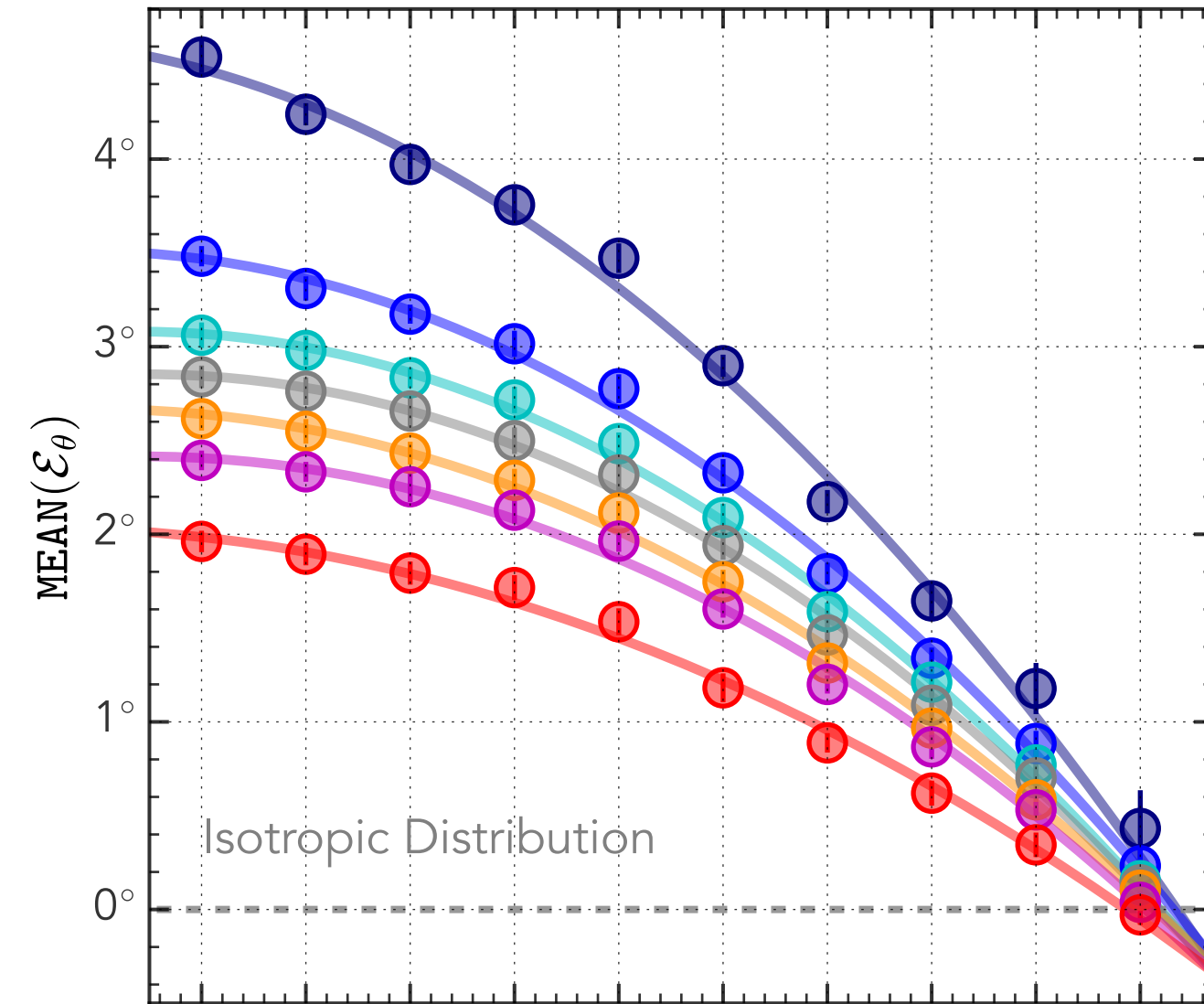
deviation from isotropy:

$$\mathcal{E}_\theta = \langle |\theta_i| \rangle_{\text{tracer}} - \langle |\theta| \rangle_{\text{isotro}}$$

$$\mathcal{E}_\phi = \left\langle \left| |\phi_i| - \frac{\pi}{2} \right| \right\rangle_{\text{tracer}} - \left\langle \left| |\phi| - \frac{\pi}{2} \right| \right\rangle_{\text{isotro}}$$

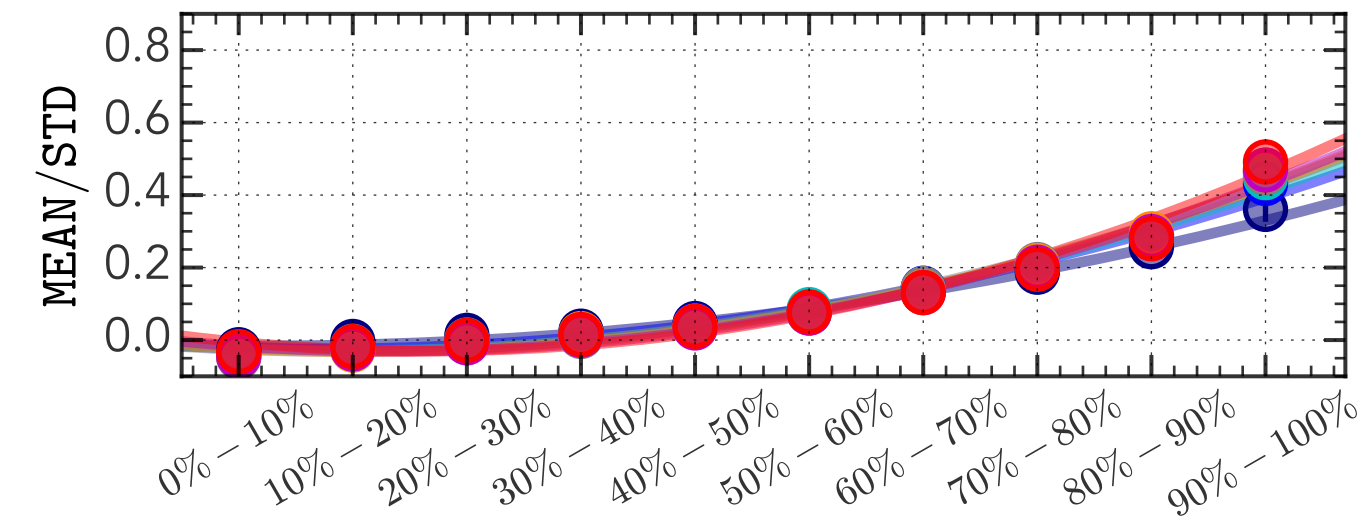
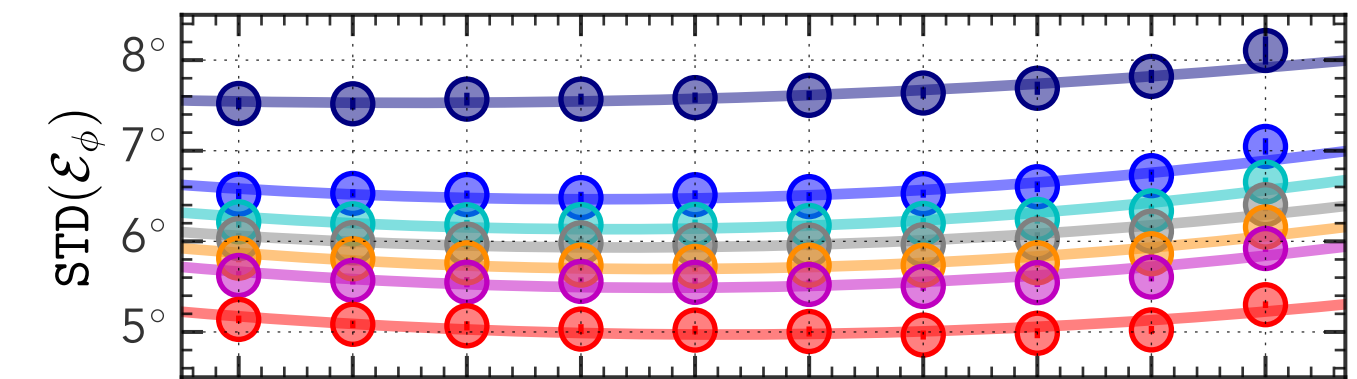
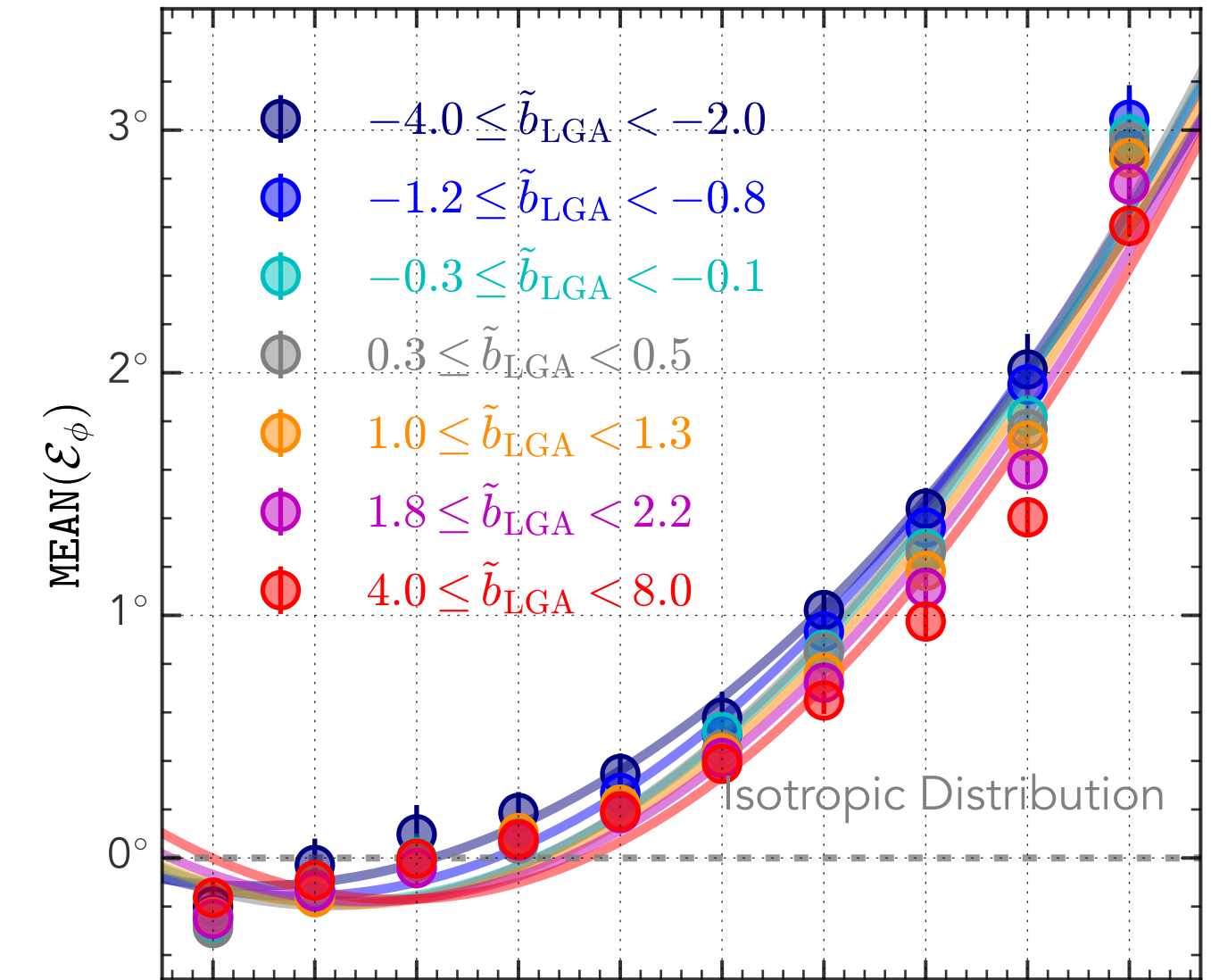
within the 5-15Mpc/h shell

Polar Anisotropy



low- $E_{\text{coupling}} \implies$ high- E_{coupling}

Azimuthal Anisotropy



low- $E_{\text{coupling}} \implies$ high- E_{coupling}

- Anisotropic environment of our Local Group

